

A GRID OF NLTE LINE-BLANKETED MODEL ATMOSPHERES OF EARLY B-TYPE STARS

THIERRY LANZ

Department of Astronomy, University of Maryland, College Park, MD 20742; lanz@astro.umd.edu

AND

IVAN HUBENY

Steward Observatory, University of Arizona, Tucson, AZ 85721; hubeny@as.arizona.edu

Received 2006 August 24; accepted 2006 November 28

ABSTRACT

We have constructed a comprehensive grid of 1540 metal line-blanketed, NLTE, plane-parallel, hydrostatic model atmospheres for the basic parameters appropriate to early B-type stars. The BSTAR2006 grid considers 16 values of effective temperatures, $15,000 \text{ K} \leq T_{\text{eff}} \leq 30,000 \text{ K}$ with 1000 K steps, 13 surface gravities, $1.75 \leq \log g \leq 4.75$ with 0.25 dex steps, six chemical compositions, and a microturbulent velocity of 2 km s^{-1} . The lower limit of $\log g$ for a given effective temperature is set by an approximate location of the Eddington limit. The selected chemical compositions range from twice to one-tenth of the solar metallicity and metal-free. Additional model atmospheres for B supergiants ($\log g \leq 3.0$) have been calculated with a higher microturbulent velocity (10 km s^{-1}) and a surface composition that is enriched in helium and nitrogen and depleted in carbon. This new grid complements our earlier OSTAR2002 grid of O-type stars (our Paper I). The paper contains a description of the BSTAR2006 grid and some illustrative examples and comparisons. NLTE ionization fractions, bolometric corrections, radiative accelerations, and effective gravities are obtained over the parameter range covered by the grid. By extrapolating radiative accelerations, we have determined an improved estimate of the Eddington limit in absence of rotation between 55,000 and 15,000 K. The complete BSTAR2006 grid is available at the TLUSTY Web site.

Subject headings: methods: numerical — radiative transfer — stars: atmospheres — stars: early-type

Online material: color figures

1. INTRODUCTION

This paper is the second of a series dealing with new grids of non-LTE (NLTE), metal line-blanketed model atmospheres of hot stars. Lanz & Hubeny (2003, hereafter Paper I) was devoted to O-type stars (OSTAR2002), covering a range of effective temperatures from 27,500 to 55,000 K and 10 chemical compositions from metal-rich relative to the Sun to metal-free. The new grid presented in this paper complements this initial grid by extending it to cooler temperatures ($15,000 \text{ K} \leq T_{\text{eff}} \leq 30,000 \text{ K}$). Together, these two grids provide a full set of model atmospheres necessary to construct composite model spectra of young clusters of massive stars, OB associations, and starburst galaxies. We plan a third forthcoming paper to further extend these two grids of NLTE line-blanketed model atmospheres to late B-type and early A-type stars.

Advances in numerical schemes applied to stellar atmosphere modeling and large increases in computational resources during the last decade have provided the ability to calculate essentially “exact,” fully blanketed NLTE hydrostatic model stellar atmospheres. Earlier model atmospheres of B-type stars neglected either departures from local thermodynamic equilibrium (LTE; Kurucz 1993) or the effect of line opacity from heavier species such as the initial H-He NLTE model atmospheres (Mihalas & Auer 1970). Major improvements are thus expected from the inclusion of a detailed treatment of line opacity without assuming LTE, since these two ingredients are important in hot stellar atmospheres. These advances have been well documented with applications of the OSTAR2002 grid to the analysis of O stars in the Galaxy and the Small Magellanic Cloud (SMC; e.g., Bouret et al. 2003; Heap et al. 2006; T. Lanz et al. 2007, in preparation). The new model atmospheres predict and consistently match lines of

different ions and species, hence robustly supporting the newly derived properties of O stars such as effective temperatures that are systematically lower than those derived using unblanketed model atmospheres. Furthermore, because the new model atmospheres incorporate the contribution of all significant species to the total opacity, NLTE abundance studies will include all background opacities in contrast to earlier line-formation studies based on NLTE model atmospheres that incorporated only hydrogen and helium besides the studied species. These new NLTE model atmospheres should therefore predict ionization balances with significantly better precision, hence enhancing the accuracy of abundance determinations from lines of several ions.

In Paper I, we discussed extensively the reasons why hydrostatic model atmospheres are relevant for O star studies despite neglecting the stellar wind. We argued in particular that the basic properties of massive stars are determined on much safer grounds using selected lines formed in the quasi-static photosphere rather than from lines formed in the supersonic wind. Indeed, many uncertainties still remain in current wind models such as wind clumping (Bouret et al. 2005) and wind ionization by X-rays (e.g., Martins et al. 2005). OSTAR2002 models are frequently adopted now to describe the density structure of the inner, quasi-static layers of unified models of O stars, such as the CMFGEN models (Hillier & Miller 1998).

On the main sequence, the radiatively driven stellar winds observed in O-type stars are vanishing very quickly in early B-type stars (around B0–B1 spectral type; that is, $T_{\text{eff}} \lesssim 30,000 \text{ K}$). Using hydrostatic model atmospheres to study B dwarfs is therefore fully appropriate. The case of B supergiants may however differ because of their large spherical extension and wind. Dufton et al. (2005) analyzed several luminous B supergiants in the SMC. They derived similar results from the hydrostatic TLUSTY model

atmospheres and from the unified FASTWIND wind models (Santolaya-Rey et al. 1997; Herrero et al. 2002). Their careful discussion thus provides adequate grounds for using TLUSTY models in these spectral analyses, keeping within the same limits as those set for O-type stars.

The paper is organized as follows. Section 2 recalls the basic assumptions made in TLUSTY, the numerical methods used, and recent changes implemented in the program. The bulk of atomic data used in the model calculations are the same as those adopted for the OSTAR2002 models, and § 3 is restricted to only document the adopted changes or improvements in model atoms. We describe briefly the grid in § 4, and we present some illustrative results in § 5.

2. TLUSTY MODEL ATMOSPHERES

The computer program TLUSTY (Hubeny 1988; Hubeny & Lanz 1995) has been developed for calculating model stellar atmospheres assuming an atmosphere with a plane-parallel geometry, in hydrostatic equilibrium, in radiative equilibrium (or in radiative-convective equilibrium for cool stellar atmospheres), and allowing for departures from LTE for an arbitrary set of chemical species represented by model atoms of various complexity. TLUSTY has been designed so that it allows for a total flexibility in selecting and treating chemical species and opacities. Because of limited computational resources, the original idea was to incorporate explicitly in the NLTE calculation only the most important opacity sources (in hot stellar atmospheres, H and He mainly; C, N, and O, in later stages). However, the introduction of the hybrid complete linearization/accelerated lambda iteration (Hubeny & Lanz 1995) method has enabled us to build model atmospheres accounting for all significant species and opacities (continua as well as lines) without assuming LTE—the so-called NLTE fully blanketed model atmospheres. Lanz & Hubeny (2003) have presented the first extensive grid of such model atmospheres for the range of stellar parameters covered by O-type stars. In Paper I we described the physical assumptions and numerical methods used to calculate these model atmospheres. In the present study, we use the same methods and, in particular, we use opacity sampling (OS) to represent the complicated frequency-dependent iron line opacity. Because of the smaller thermal Doppler width and the lower microturbulent velocity (2 km s^{-1}), sampling the spectrum with the same characteristic frequency step (0.75 fiducial Doppler width) thus requires many more frequencies over the whole spectrum than in the OSTAR2002 models—typically, the BSTAR2006 models consider about 380,000 frequencies. For more details on the methods, the reader is referred to Paper I.

The OSTAR2002 models have been calculated with TLUSTY, version 198, while we have used a later version, 201, for the BSTAR2006 grid. The newer version of TLUSTY is a unified program that can calculate accretion disk models as well as model stellar atmospheres (two separate codes existed for ver. 198). Other additions consist in a treatment of Compton scattering, opacities at high energies, improvements for convection in cool models, and minor changes and fixes. These changes are however not directly relevant to the modeling of O- and B-type stars.

3. ATOMIC DATA

We have closely followed the approach adopted in Lanz & Hubeny (2003) regarding the treatment and inclusion of atomic data in the model atmosphere calculations. In particular, the bulk of the data are taken from Topbase, the Opacity Project database (Opacity Project Team 1995, 1997).¹ The level energies have

been systematically updated with the more accurate experimental energies extracted from the Atomic and Spectroscopic Database at NIST.² Finally, we use the extensive Kurucz iron line data (Kurucz 1994).³ For a detailed description, refer to § 4 of Paper I.

Compared to the OSTAR2002 grid, we consider lower ions explicitly in the NLTE calculations because we anticipate lower ionization from the lower effective temperatures. For instance, detailed models of some neutral atoms (C I, N I, O I, Ne I) are included in BSTAR2006 models, while some of the highest ions such as O VI are skipped. The ionization balance of the different species (see § 5) provides a good justification of this choice. We have used the same model atoms for most ions in common with the OSTAR2002 grid. In some instances, however, we have expanded the model atoms to include all levels below the ionization limit as they are listed in the Opacity Project (OP; hence, completion up to $n = 10$ as in the OP). The following subsections provide the relevant details of new or updated model atoms—refer to Paper I for model atoms not discussed here.

Table 1 summarizes the atomic data included in the model atmospheres, as well as the references to the original calculations collected in Topbase. Data files and Grotrian diagrams may be retrieved from the TLUSTY Web site.⁴

3.1. Hydrogen and Helium

In TLUSTY, version 201, we have introduced a new default for collisional excitation in hydrogen. Collisional rates are now computed following Giovanardi et al. (1987) with expressions valid in a wide temperature domain ($3000 \text{ K} \leq T \leq 500,000 \text{ K}$). This modification introduces small changes in hydrogen level populations with respect to the older collisional rates (Mihalas et al. 1975). We compare overlapping models from the OSTAR2002 and BSTAR2006 grids in § 4 and we attribute small changes in Balmer line profiles to differences in hydrogen collisional excitation. With this exception, hydrogen and helium are treated as described in Paper I.

3.2. Carbon

We have added a detailed 40 level model atom for neutral carbon and expanded the C III model atom. The new C I model atom includes all 239 levels found in Topbase with level energies below the ionization limit. Levels up to the $4p$ levels ($E < 83,000 \text{ cm}^{-1}$), that is, 13 singlet levels, 14 triplet levels, and $2s2p^3 \ ^5S^0$, are included as individual levels. The 211 levels with higher excitation are grouped into 12 superlevels, 6 in the singlet system and 6 in the triplet system. All the other quintet levels have energies above the ionization limit and are unaccounted for in this model atom. The original C III model atom included levels only up to $n = 6$. We have examined in Paper I the consequences of neglecting more excited levels. While this choice has little effect on the atmospheric structure, we found that the total C^{+3} to C^{+2} recombination rate is underestimated by a factor 2–3. We have therefore opted to update the C III model atom and include all levels listed in Topbase (below the ionization limit). The updated 46 level model atom includes 34 levels with the lower excitation individually ($E < 342,000 \text{ cm}^{-1}$). Levels with higher excitation are grouped into 12 superlevels, 6 in the singlet system and 6 in the triplet system.

² See <http://physics.nist.gov/PhysRefData/ASD/index.html>.

³ See <http://kurucz.harvard.edu/cdroms.html>.

⁴ See <http://nova.astro.umd.edu>.

¹ See <http://vizier.u-strasbg.fr/topbase>.

TABLE 1
ATOMIC DATA INCLUDED IN THE NLTE MODEL ATMOSPHERES

Ion	(Super)Levels	Individual Levels	Lines	References
H I	9	80	172	...
H II	1	1
He I	24	72	784	1
He II	20	20	190	...
He III	1	1
C I	40	239	3201	2
C II	22	44	238	3
C III	46	95	738	4
C IV	25	55	330	5
C V	1	1
N I	34	115	785	6
N II	42	247	3396	2
N III	32	68	549	3
N IV	48	126	1093	4
N V	16	55	330	5
N VI	1	1
O I	33	69	418	7
O II	48	218	3484	6
O III	41	267	3855	2
O IV	39	94	922	3
O V	6	6	4	4
O VI	1	1
Ne I	35	191	2715	8
Ne II	32	167	2301	7
Ne III	34	149	1354	7
Ne IV	12	18	38	6
Ne V	1	1
Mg II	25	53	306	9
Mg III	1	1
Al II	29	81	536	10
Al III	23	53	306	9
Al IV	1	1
Si II	40	64	392	11
Si III	30	105	747	10
Si IV	23	53	306	9
Si V	1	1
S II	33	226	4166	13
S III	41	235	3452	12
S IV	38	100	909	11
S V	25	131	1171	10
S VI	1	1
Fe II	36	10 921	1 264 969	14, 15
Fe III	50	12 660	1 604 934	14, 16
Fe IV	43	13 705	1 776 984	14, 17
Fe V	42	11 986	1 008 385	14, 18
Fe VI	1	1

REFERENCES.—(1) <http://physics.nist.gov/PhysRefData/ASD/index.html>; (2) Luo & Pradhan 1989; (3) Fernley et al. 1999; (4) Tully et al. 1990; (5) Peach et al. 1988; (6) V. M. Burke & D. J. Lennon 2007, in preparation; (7) K. Butler & C. J. Zeippen 2007, in preparation; (8) Hibbert & Scott 1994; (9) K. T. Taylor 2007, in preparation; (10) Butler et al. 1993; (11) Mendoza et al. 1995; (12) Nahar & Pradhan 1993; (13) K. Butler et al. 2007, in preparation; (14) Kurucz 1994; (15) Nahar 1997; (16) Nahar 1996; (17) Bautista & Pradhan 1997; (18) Bautista 1996.

3.3. Nitrogen

We have added a detailed 34 level model atom for neutral nitrogen and expanded the N II and N IV model atoms. The new N I model atom includes the lower 27 levels up to $4p\ ^2P^0$ ($E < 108,000\text{ cm}^{-1}$). The next 88 levels, up to $n \approx 7$, are grouped into 7 superlevels (3 in the doublet system, 4 in the quartets). Levels from the sextet system are all above the ionization limit and are neglected. The original N II model atom did not incorporate any

high levels, $n \geq 6$. The updated N II model atom includes 32 individual levels, up to $4s$ ($E < 200,000\text{ cm}^{-1}$), and 3 additional singlets and the 4 listed quintet levels below the ionization limit. The remaining 326 levels below the limit are grouped into 10 superlevels. The NIST database also lists several excited levels in intermediate coupling; these levels are skipped because the OP calculations assumed *LS*-coupling. The N IV model atom was updated similarly to C III, including the lower 34 individual levels ($E < 525,000\text{ cm}^{-1}$), and grouping the 92 levels with higher excitation ($n \leq 10$) into 14 superlevels.

3.4. Oxygen

The model atmospheres incorporate five ions of oxygen, O I to O V, and the ground state of O VI. We have added a detailed 33-level model atom for neutral oxygen. The new O I model atom includes the lower 23 levels up to $5p\ ^3P$ ($E < 104,000\text{ cm}^{-1}$). All remaining 46 levels, which are listed in Topbase as being below the ionization limit, are grouped into 10 superlevels evenly split between the triplet and quintet systems. O II and O III have a relatively rich structure of levels. We have thus updated these two model atoms to include the high excitation levels. The updated O II model atom includes the lower 34 individual doublet and quartet levels, up to $4p\ ^2P^0$ ($E < 250,000\text{ cm}^{-1}$), and 2 sextet levels ($3s\ ^6S^0$ and $3p\ ^6P$). The higher 182 levels are grouped into 12 superlevels, 6 in the doublet and 6 in the quartet systems. Levels in intermediate coupling listed in the NIST database are skipped. The updated O III model atom includes the first 28 individual levels, up to $3s\ ^3P$ ($E \leq 350,000\text{ cm}^{-1}$), and groups the higher 239 levels below the limit into 13 superlevels in the singlet (5), triplet (5), and quintet (3) systems. The model atom of O IV is the same as in Paper I. Because of lower temperatures, we adopted simplified model atoms for the highest ions. O V model atom only includes the first 6 levels (up to $2p^2\ ^1S$), and excludes the higher, very excited levels ($E > 540,000\text{ cm}^{-1}$). The highest level is the ground state of O VI, while O VII is neglected altogether.

3.5. Neon

Simple model atoms have been used in Paper I. For the BSTAR2006 grid, we have added a new detailed 35 level model atom for neutral neon, and we have extended the Ne II and Ne III model atoms. The Ne I model atom includes the lower 23 levels listed in Topbase, and groups all 168 higher levels into 12 superlevels. Following the Opacity Project, this model atom is built assuming *LS*-coupling. However, the level structure of neutral neon is poorly described in *LS*-coupling, and intermediate *JK*-coupling should be preferred (e.g., Sigut 1999). We have constructed a 79 level model atom with fine structure and allowing for intermediate coupling (Cunha et al. 2006). Changes in atmospheric structure are negligible as expected. The predicted Ne I line strengths in the red spectrum are also little changed. While an exhaustive differential study of the two model atoms needs to be performed, we believe thus that the BSTAR2006 model spectra provide a good starting point to analyze Ne I lines. Detailed line formation calculations and accurate abundance determinations would however benefit from using the 79 level model atom. The updated Ne II model atom includes the lower 23 individual levels, up to $4s\ ^2P$ ($E < 284,000\text{ cm}^{-1}$). The higher 144 levels are grouped into 9 superlevels. The Ne III model atom includes the first 7 singlet and 10 triplet levels, up to $3p\ ^3P$ ($E < 400,000\text{ cm}^{-1}$), and the first 5 quintet levels individually. All 127 higher levels below the ionization limit are grouped into 12 superlevels. The Ne IV model atom is identical to the model used in Paper I.

3.6. Magnesium

We have added magnesium to the list of explicit species because of the importance of Mg II $\lambda\lambda 2798, 2803, 4481$ lines as diagnostics in spectral analyses of B-type stars. The Mg II model atom includes all individual levels up to $n = 6$ and four superlevels ($7 \leq n \leq 10$). The data extracted from Topbase have been extended to include all levels up to $10l^2L^0$. We have included the two fine-structure levels of $3p^2P^0$ for an exact treatment of the Mg II $\lambda 2800$ resonance doublet, described with depth-dependent Voigt profiles. Because of the high ionization energy of Mg^{2+} (about 80 eV), we have restricted our model of high Mg ions to the ground state of Mg III and neglected higher ions.

3.7. Aluminum

The UV spectrum of B stars also have important diagnostics lines of aluminum, such as the resonance lines Al II $\lambda 1670$ and Al III $\lambda\lambda 1855, 1863$, thus leading us to incorporate aluminum in the BSTAR2006 models. Two ions, Al II and Al III, and the ground state of Al IV are explicitly considered. The Al II model atom includes the first 20 levels individually, up to $5p^1P^0$ ($E < 126,000 \text{ cm}^{-1}$), and groups all 61 higher levels in Topbase into 9 superlevels. The Al III was built similarly to the Mg II model atom. It includes all individual levels up to $n = 6$ and 4 superlevels ($7 \leq n \leq 10$). The data in Topbase have been extended to include all levels up to $10l^2L^0$. We have included the 2 fine-structure levels of $3p^2P^0$ for an exact treatment of the Al III $\lambda 1860$ resonance doublet. The UV resonance lines, Al II $\lambda 1670$ and Al III $\lambda\lambda 1855, 1863$, are described with depth-dependent Voigt profiles, while we assume depth-independent Doppler profiles for all other lines.

3.8. Silicon

We have added a new Si II model atom, while keeping the Si III and Si IV model atoms used in the OSTAR2002 grid. The Si II model atom includes all the levels in Topbase below the ionization limit and 4 levels slightly above the limit, up to $3p^3^2D^0$ ($E < 126,000 \text{ cm}^{-1}$). Excited doublet levels are grouped into 4 superlevels ($7 \leq n \leq 10$). All other levels are treated as individual levels. We have included the 2 fine-structure levels of the ground state, $3p^2P^0$ for an exact treatment of the 6 far-UV resonance doublets, which are described with depth-dependent Voigt profiles. Autoionization represents an important channel for the silicon ionization equilibrium. It results in broad resonances in the OP photoionization cross sections, and autoionization is thus incorporated in our models via the OP cross sections (Lanz et al. 1996).

3.9. Sulfur

The model atmospheres incorporate four ions of sulfur, S II to S V. All the model atoms have been updated and extended. We restricted the highest ion to the ground state of S VI and neglected S VII. The S II model atom includes all the energy levels found in Topbase below the ionization limit, except the 3 highly excited levels in the sextet system. The first 23 levels, up to $4s^2S$ ($E < 137,000 \text{ cm}^{-1}$), are included individually. The more excited levels are grouped into 10 superlevels. All 235 S III levels in Topbase are incorporated in the model atom. The lowest 24 singlet and triplet levels, up to $4p^1S$ ($E < 190,000 \text{ cm}^{-1}$), and the first 5 quintet levels are included individually. The higher levels are grouped into 5 singlet superlevels, 5 triplet superlevels, and 2 quintet superlevels. The S IV model atom accounts for 100 individual levels. The lowest 21 doublet levels, up to $4p^2P$ ($E < 293,000 \text{ cm}^{-1}$), are included individually, and more excited dou-

blets are grouped into 3 superlevels. The model atom also accounts for the fine structure of the S IV ground state, $3p^2P^0$. The first 11 quartet levels ($E < 320,000 \text{ cm}^{-1}$) are treated as individual explicit levels, while higher quartet levels are grouped into 2 quartet superlevels. Finally, the S V model atom is described with 20 individual levels, up to $E < 400,000 \text{ cm}^{-1}$; all the higher levels are grouped into 5 superlevels. All lines of the four ions are assigned depth-independent Doppler profiles.

3.10. Iron

We have shifted the selected NLTE iron ions toward lower ionization, explicitly incorporating four ions (Fe II to Fe V) in the BSTAR2006 model atmospheres. The Fe III, Fe IV, and Fe V model atoms are the same models that we used in Paper I. The Fe II model atom was built in the same way and is based on Kurucz (1994) extensive semiempirical calculations. Levels of same parity and close energies are grouped into superlevels. The Fe II model atom includes 23 even and 13 odd superlevels. The Fe II photoionization cross sections are also extracted from calculations by the Ohio State group (see Table 1). They assumed *LS*-coupling, and we could typically assign theoretical cross sections to observed levels for the lowest 20–30 levels. We have assumed an hydrogenic approximation for higher-excitation levels. The data are then combined to setup cross sections for the superlevels, and resampled using the resonance-averaged photoionization technique (Bautista et al. 1998; Paper I). For each ion, we consider the full linelist from Kurucz (1994). TLUSTY dynamically rejects the weaker lines, based on the line *gf*-values, the excitation energies, and the ionization fractions. With this selection, mostly Fe II and Fe III lines are included in cooler models, while many more Fe IV and Fe V lines are explicitly accounted for in the hotter models. This process selects all the necessary lines that do contribute to the total opacity, out of the list of 5.7 million iron lines. Typically, about 1 million lines are selected. It may select, however, as few as 500,000 lines and up to 2 million lines in some models, with the larger numbers in cooler and high-gravity models. The adopted selection criterion is inclusive enough to ensure an appropriate description of the line blanketing effect that is mainly caused by the strongest 10^4 – 10^5 lines (Lanz & Hubeny 2003).

4. DESCRIPTION OF THE GRID

The BSTAR2006 grid covers the parameter space of early B-type stars in a dense and comprehensive way. We have selected 16 effective temperatures, $15,000 \text{ K} \leq T_{\text{eff}} \leq 30,000 \text{ K}$, with 1000 K steps, 13 surface gravities, $1.75 \leq \log g \leq 4.75$, with 0.25 dex steps, and 6 chemical compositions, from twice to one-tenth the solar metallicity and metal-free models. The effective temperature range covers spectral types B0–B5 (Schmidt-Kaler 1982). The lower limit in $\log g$ is set by the Eddington limit (see § 5.4).

Solar abundances refer to Grevesse & Sauval (1998) Sun98. We have assumed a solar helium abundance, $\text{He}/\text{H} = 0.1$ by number. All other chemical abundances are scaled from the solar values. The microturbulent velocity was set to $V_t = 2 \text{ km s}^{-1}$. On the main sequence, the covered temperature range corresponds to stars with initial masses between 4 and $15 M_{\odot}$. At lower gravities, the BSTAR2006 grid covers the case of more massive stars evolving off the main sequence.

Recent spectroscopic studies reveal that early B supergiants show strongly processed material at their surface (e.g., Dufton et al. 2005; Crowther et al. 2006). In these NLTE analyses, large microturbulent velocities ($V_t \geq 10 \text{ km s}^{-1}$) are derived for most stars. We emphasize here that these high microturbulences do not result from neglecting the wind and from assuming hydrostatic

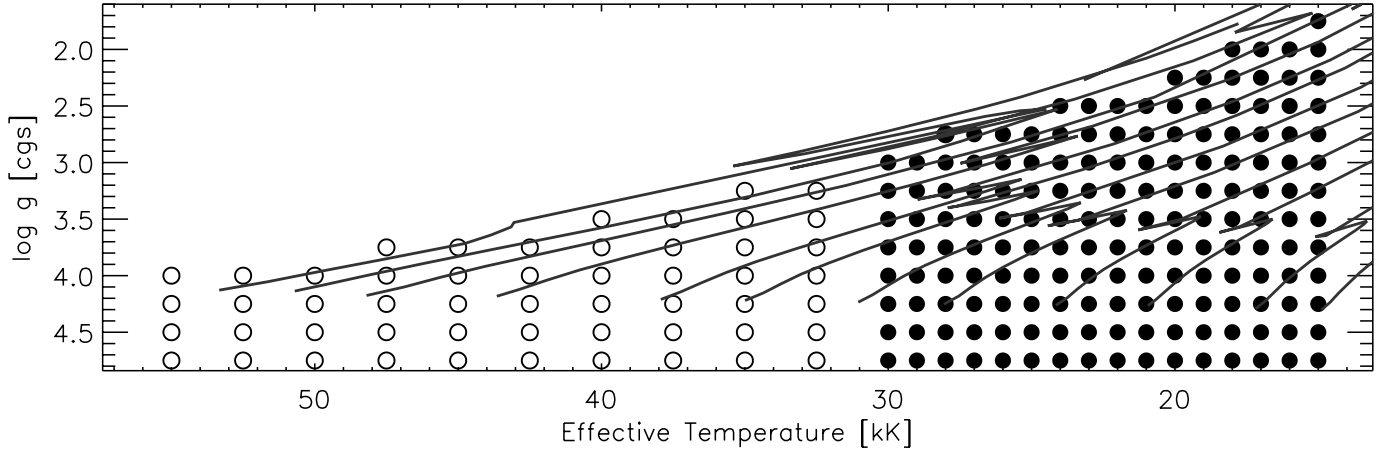


FIG. 1.—Selected BSTAR2006 grid points (*filled symbols*) and OSTAR2002 grid points (*open symbols*) in the $\log g$ vs. T_{eff} plan. The two grids overlap at $T_{\text{eff}} = 30,000$ K. Geneva evolutionary tracks (Schaller et al. 1992) are shown for solar metallicity and correspond to models with initial masses of 120, 85, 60, 40, 25, 20, 15, 12, 9, 7, 5, and $4 M_{\odot}$ from left to right, respectively.

equilibrium. Like the case of O stars, such high microturbulent velocities are obtained both from TLUSTY models and from unified models. Therefore, we decided to supplement the initial grid with two additional sets of model atmospheres suitable for supergiants ($\log g \leq 3.0$). In the first set, we adopted $V_t = 10 \text{ km s}^{-1}$, but kept solar or solar-scaled abundances. In the second set, in addition to the larger microturbulence, we increased the helium abundance to $\text{He}/\text{H} = 0.2$ by number, increased the nitrogen abundance by a factor of 5, and halved the carbon abundance (hence, an order of magnitude increase in the N/C abundance ratio, reflecting CNO-cycle processed material brought to the stellar surface).

For each of the five metallicities, we have thus computed 265 model atmospheres. This includes a full set of 163 models at the lower microturbulence and two sets of 51 models at the higher microturbulence (include only low-gravity models). At the highest metallicities, a few models very close to the Eddington limit could not be converged because of physical and numerical instabilities; these few models are thus skipped. Figure 1 illustrates the sampling of the $\log g$ versus T_{eff} diagram by our two grids of NLTE model atmospheres. The two grids overlap at $T_{\text{eff}} = 30,000$ K, and we show below a comparison.

4.1. Output Products and Availability

The model atmospheres are available at the TLUSTY Web site (see footnote 4). Similarly to the OSTAR2002 grid, each model is characterized by a unique filename describing the parameters of the model, for example BS25000g275v10CN. The first letters indicate the composition, followed by the effective temperature, the gravity, and the turbulent velocity. For models with altered surface composition (supplemental set 2), “CN” is appended to the filename. We have adopted the same key for the models’ overall composition as in the OSTAR2002 grid (see Table 2), that is: twice solar (“BC models”), solar (“BG models”), half solar (“BL models”), one-fifth solar (“BS models”), one-tenth solar (“BT models”), and metal-free (“BZ models”).⁵ Model atmospheres with even lower metallicities like those in the OSTAR2002 grid may be made available at a future time.

⁵ Model BS25000g275v10CN thus corresponds to a model with $T_{\text{eff}} = 25,000$ K, $\log g = 2.75$, $V_t = 10 \text{ km s}^{-1}$, one-fifth solar metallicity, and altered He, C, and N surface abundances.

Each model comes as a set of six files, with an identical filename’s root but a different extension. A complete description of the files’ content and format can be found in the TLUSTY User’s guide (see TLUSTY Web site); for reference, we describe them only very briefly here:

- model.5.—general input data;
- model.nst.—optional keywords;
- model.7.—model atmosphere: temperature, electron density, total density, and NLTE populations as function of depth;
- model.11.—model atmosphere summary;
- model.12.—model atmosphere: Similar to model.7, but NLTE populations are replaced by NLTE b -factors (LTE departure coefficients);
- model.flux.—model flux distribution from the soft X-ray to the far-infrared given as the Eddington flux⁶ H_{ν} (in $\text{ergs s}^{-1} \text{cm}^{-2} \text{Hz}^{-1}$) as a function of frequency.

The model fluxes are provided at all frequency points included in the calculations (about 380,000 points with an irregular sampling for the low-microturbulence models; about 175,000 points for the “v10” models). In addition, we have calculated detailed emergent spectra with SYNSPEC, version 48, in the ultraviolet ($\lambda\lambda 900\text{--}3200$) and in the optical ($\lambda\lambda 3200\text{--}10000$). Filename extensions are model.uv.7, model.vis.7, respectively, and *.17 for the continuum spectra. Additional spectra in other wavelength ranges, with altered chemical compositions, or with different values of the microturbulent velocity can be readily computed using SYNSPEC. The spectrum synthesis requires three input files, model.5, model.7, and model.nst, and the necessary atomic data files (model atoms and the relevant linelist).

⁶ The flux at the stellar surface is $F_{\nu} = 4\pi H_{\nu}$.

TABLE 2
KEY TO THE MODELS’ CHEMICAL COMPOSITIONS

Key	Metallicity (Solar Metallicity)
G.....	1
L.....	1/2
S.....	1/5
T.....	1/10
Z.....	0

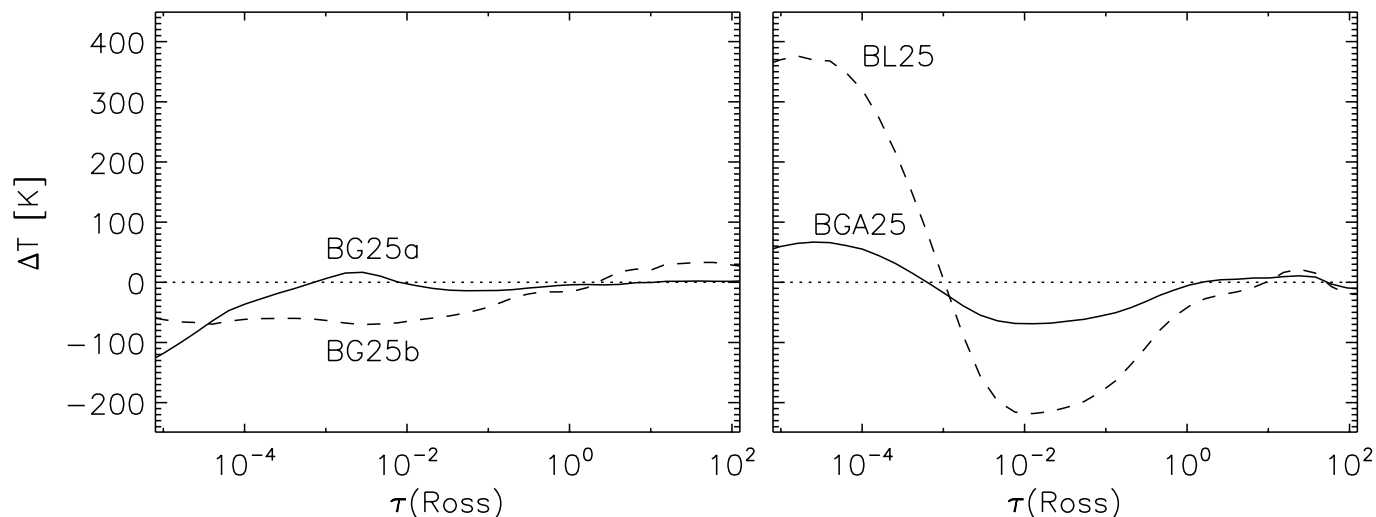


FIG. 2.—Change in the atmospheric temperature structure with respect to the reference model BG25000g400v2 ($T_{\text{eff}} = 25,000$ K, $\log g = 4.0$, solar composition, and $V_t = 2$ km s $^{-1}$). *Left:* The changes due to different assumptions in opacity sampling, using larger frequency sampling steps (BG25a; solid line) or only selecting the strongest Fe II to Fe V lines (BG25b; dashed line). *Right:* The effect of the assumed abundances, with lower CNO abundances and higher Ne abundance (BGA25; solid line), or lowering the overall metallicity by half (BL25; dashed line).

The analysis of individual stars may require to interpolate within the model grid. Interpolation procedures were discussed in § 5.2 of Paper I. Since the BSTAR2006 grid is sampled with a temperature step (1000 K) that is even finer than the step in the OSTAR2002 grid, these procedures can be applied. As the safest method, however, we advise to interpolate the atmospheric structure for the selected parameters, followed by recalculating the spectrum with SYNSPEC. This approach provides a very fast way to apply these fully blanketed NLTE model atmospheres to detailed spectrum analyses of B stars.

4.2. Model Sensitivity

We have investigated the effect of several assumptions made to calculate the BSTAR2006 model atmospheres, including the treatment of iron opacity, the choice of the “standard” solar abundances, and the adopted model atoms. For these tests, we have selected the solar composition model, $T_{\text{eff}} = 25,000$ K, $\log g = 4.0$, $V_t = 2$ km s $^{-1}$, and one solar composition model ($T_{\text{eff}} = 30,000$ K, $\log g = 3.0$) where the two grids overlap.

The first test expands at lower T_{eff} on the tests of iron line-blanking made in Paper I, § 8. On the basis of earlier experience and mostly on the importance of representing blends as accurately as possible,⁷ we preferentially adopt opacity sampling (OS) over the opacity distribution function approach. In order to ensure the best description of line blends, we have used a small frequency sampling step (0.75 fiducial Doppler width) as the standard. We have then tested the resulting effect on the atmospheric structure of a larger sampling step and of a different line strength selection criterion. Compared to the reference solar composition model BG25000g400v2, model BG25a uses a frequency sampling step that is 40 times larger for the iron lines, without any changes for lines of lighter species, hence reducing the total number of frequencies by a factor of about 3. Model BG25b limits the dynamically selected iron lines to the strongest lines, selecting about 85,000 lines of Fe II to Fe V compared to over 1,500,000 lines in the reference model. Figure 2 (*left panel*) displays the change in

the atmospheric temperature structure with respect to the reference model. The temperature differences remain small, quite comparable to the changes seen in similar tests for O star model atmospheres. As a reference, the typical numerical accuracy achieved in conserving the total radiative flux corresponds to uncertainties up to 10–20 K on the local temperature. The limited extent of temperature changes found in these tests implies only minimal changes in the predicted spectra, as exemplified by the test with different standard abundances where the temperature changes are somewhat larger (see below, and Figs. 2 and 3). We believe that the tests that we performed for O and B star model atmospheres are representative to the numerical accuracy of OS in hot model atmospheres. The larger sampling step might thus be quite sufficient for most purposes, hence allowing a substantial saving in computer time (up to a factor of 3). However, because we intend the BSTAR2006 grid to serve as reference NLTE model atmospheres, and since accuracy testing is practicable only in a limited extent, we have elected to keep the finest sampling to ensure the best description of iron line blanketing.

When we initiated the work on the OSTAR2002 and BSTAR2006 grids, we adopted the standard solar abundances from Grevesse & Sauval (1998). In the interim, several studies showed that the carbon, nitrogen, and oxygen abundances in the Sun are significantly lower, about 70% of the earlier standard values (Asplund et al. 2005). The new lower abundances are in good agreement with abundances measured in B stars in the solar neighborhood (e.g., Cunha & Lambert 1994). In addition, Cunha et al. (2006) found that the neon abundance in the Orion B stars is twice the standard value, and they argued that this value might be used as a good proxy for the solar abundance. Accordingly, we investigated the differences in atmospheric structure and predicted spectrum from these updated C, N, O, and Ne abundances, calculating a model atmosphere using $T_{\text{eff}} = 25,000$ K, $\log g = 4.0$, $V_t = 2$ km s $^{-1}$, and the newer abundances. For the comparison, we considered four different model atmospheres: two models in the grid for solar (BG25) and half-solar (BL25) metallicities (since the new CNO values fall in between), the BGA25 model atmosphere with the new abundances, and a model spectrum BGZ25 calculated with the solar composition atmospheric structure and applying the new abundances in the spectrum calculation

⁷ See Najarro et al. (2006) for an example of the importance of Fe IV lines blending He I $\lambda 584$.

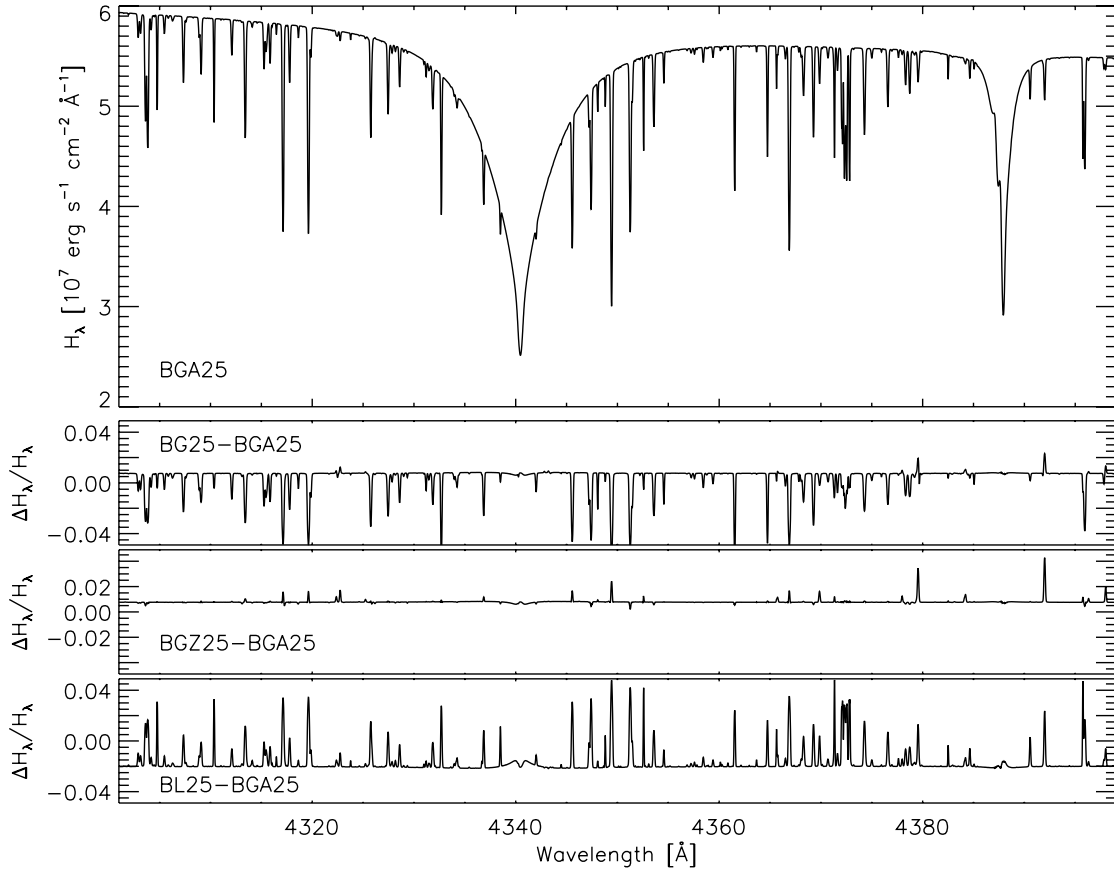


FIG. 3.—Model spectrum around $H\gamma$ for a model atmosphere with updated C, N, O, and Ne abundances ($T_{\text{eff}} = 25,000$ K, $\log g = 4.0$, solar composition with new CNO and Ne abundances, and $V_t = 2$ km s $^{-1}$). The bottom panels show the relative spectrum changes on the same scale for models with various abundances (see text).

step only. Figure 2 (*right panel*) shows the temperature differences between the first three models. The changes resulting from the changes in the light element abundances remain limited ($\Delta T < 90$ K), and they are as much as 4 times smaller than the differences between the solar and half-solar metallicity models. This comparison shows the importance of iron line blanketing in establishing the temperature structure. We then compare the resulting effect of the abundance changes on the predicted spectra. We show in Figure 3 the spectral region around $H\gamma$, also including He I $\lambda 4388$ and many O II lines, and the relative differences between the model spectra. The absolute *continuum* flux level changes by less than 1%, and up to 2% for the half-solar metallicity model. The largest changes, 4% and higher, are seen in *line cores*, mostly in the strength of O II lines, which are directly related to changes in the oxygen abundance in the spectrum calculation. Most importantly, however, the BGZ25 model spectrum is very similar to the BGA25 model spectrum (see Fig. 3, *middle panel of three showing the relative changes*) and, therefore, shows that the largest spectrum changes directly result indeed from abundance changes in the spectrum synthesis step rather than from the indirect change in the atmospheric structure. In most cases, it might therefore be appropriate to use the solar composition model atmospheres and only to recalculate the spectra with the updated abundances of light species.

We finally stress that the highest accuracy will always be reached when tailoring the model atmospheres to the stars studied, using the closest estimates of all stellar parameters and abundances; the BSTAR2006 grid primarily intends to provide the best starting point for these detailed analyses.

Finally, we examine the sensitivity of model stellar atmospheres to the adopted model atoms by comparing a model atmosphere with identical parameters ($T_{\text{eff}} = 30,000$ K, $\log g = 3.0$, solar composition, and $V_t = 10$ km s $^{-1}$) from the two grids, BSTAR2006 and OSTAR2002. Section 3 details the atom models and changes between the two grids. In summary, the newer models include lower ionization ions as well as more complete model atoms, in particular for C III, N II, N IV, O II, O III, and the Ne and S ions. Figure 4 illustrates the very high consistency between the two grids: the very rich ultraviolet and the visible spectra predicted by these two calculations are very much the same. Rotational broadening ($V \sin i = 50$ km s $^{-1}$) has been applied to the spectra to show a comparison that is more relevant to actual spectral analyses. In the visible spectrum, we particularly note on one hand the excellent agreement in $H\gamma$, He I $\lambda 4471$, and He I $\lambda 4388$ in emission, while weak metal lines on the other hand reveal some discrepancies. A closer examination shows that Balmer lines are predicted very slightly weaker in the BSTAR2006 models than in the OSTAR2002 models, which we interpret as resulting from the adoption of different bound-bound collisional rates in hydrogen. This small difference has an insignificant impact on the derived stellar parameters. All metal lines showing changes in the optical spectrum arise from highly excited levels. Hence, these differences are likely a consequence of the newer model atoms that explicitly include higher levels in the NLTE calculations. The gf -values of some lines in the linelist used by SYNSPEC have also been updated from the original Kurucz data to OP data for consistency between the model atmosphere and the spectrum synthesis calculations. We expect the OP values to

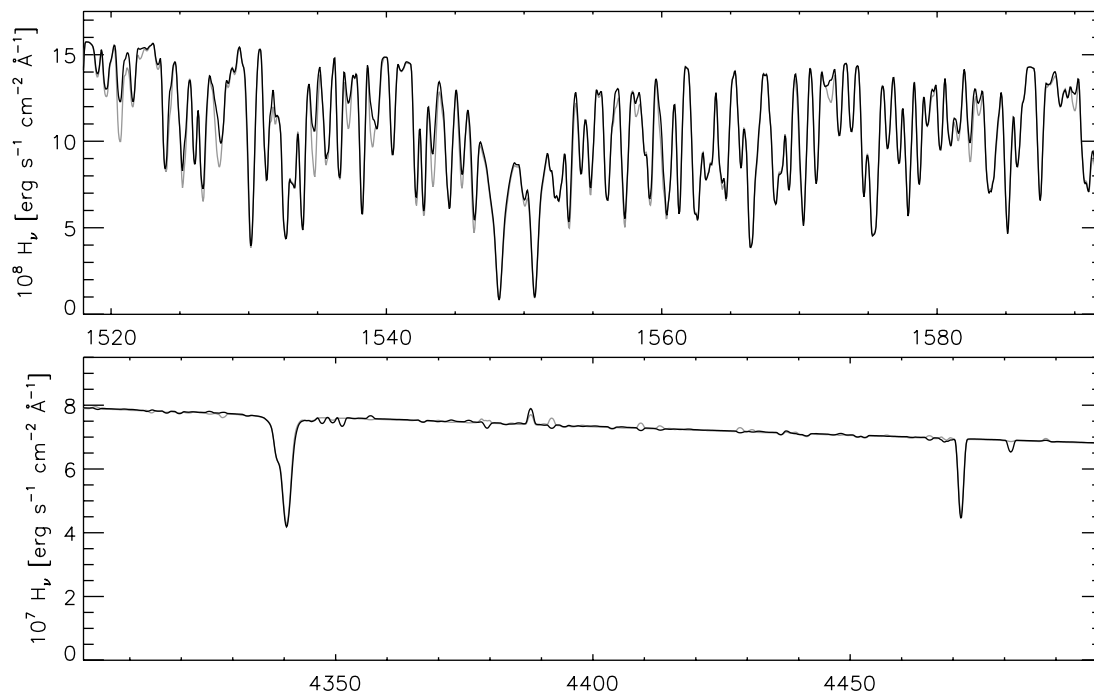


FIG. 4.—Ultraviolet and visible model spectra for model atmospheres with identical stellar parameters ($T_{\text{eff}} = 30,000$ K, $\log g = 3.0$, solar composition, and $V_t = 10$ km s $^{-1}$) from the two grids: BSTAR2006 (black line) and OSTAR2002 (gray line). A rotational broadening of 50 km s $^{-1}$ was applied to the spectra. [See the electronic edition of the Supplement for a color version of this figure.]

be of higher accuracy in general. From this last test, we may therefore conclude that the two model atmosphere grids will yield consistent results. Although these models include a detailed and extensive treatment of line blanketing, we stress that these models should not be applied blindly to the analysis of some weak lines in the visible and infrared range without a line-by-line assessment of the adequacy of the model atoms in each case.

5. REPRESENTATIVE RESULTS

In this section, we show some basic properties and important trends describing model atmospheres of B-type stars, in partic-

ular a comparison to Kurucz (1993) LTE model atmospheres and spectra, ionization structure, bolometric corrections, and radiative accelerations.

5.1. Comparison to LTE Model Atmospheres

We first compare the atmospheric structure of NLTE BSTAR2006 and LTE Kurucz (1993) model atmospheres. Figure 5 displays the temperature stratification of a series of solar composition and one-tenth solar metallicity NLTE and LTE models, $T_{\text{eff}} = 25,000$ K, $\log g = 3.0$ and 4.0. At large depths ($\tau_{\text{Ross}} \gtrsim 1$, corresponding to mass column densities larger than 0.5 g cm $^{-2}$) where the

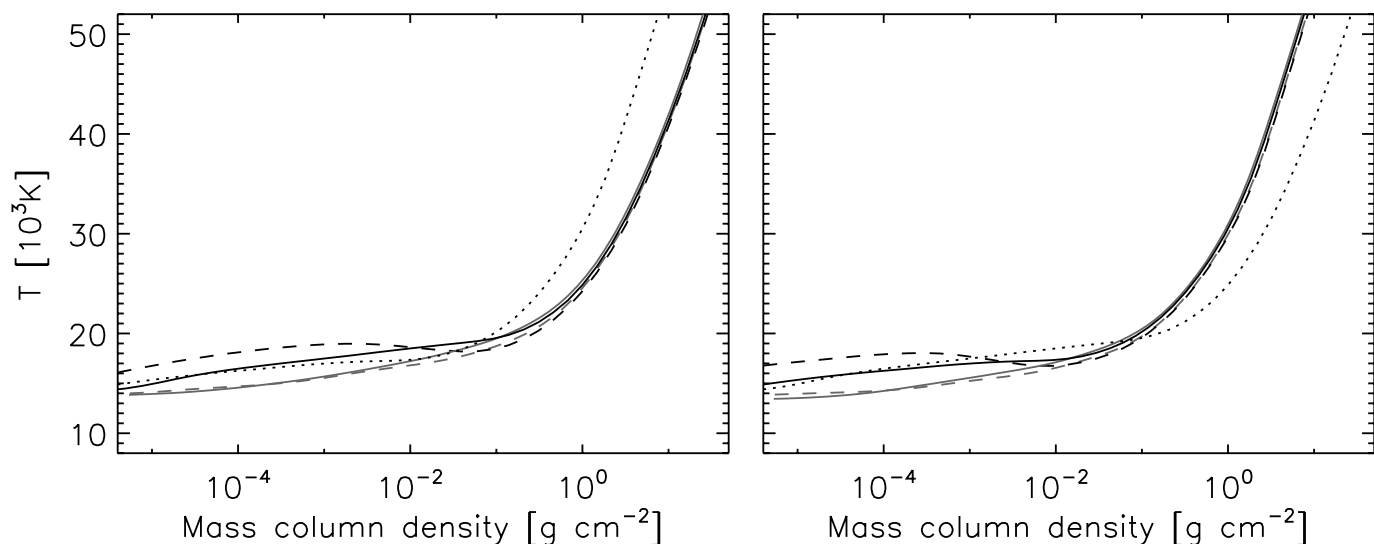


FIG. 5.—Temperature structure of model atmospheres with $T_{\text{eff}} = 25,000$ K, $\log g = 3.0$ (left panel) and $\log g = 4.0$ (right panel). Black lines show the temperature stratification of the NLTE BSTAR2006 models, compared to the LTE Kurucz models (gray lines). Solar composition models (full lines) and metal-poor (1/10 solar) models (dashed lines) are displayed. The dotted lines illustrate the effect of the different surface gravities. [See the electronic edition of the Supplement for a color version of this figure.]

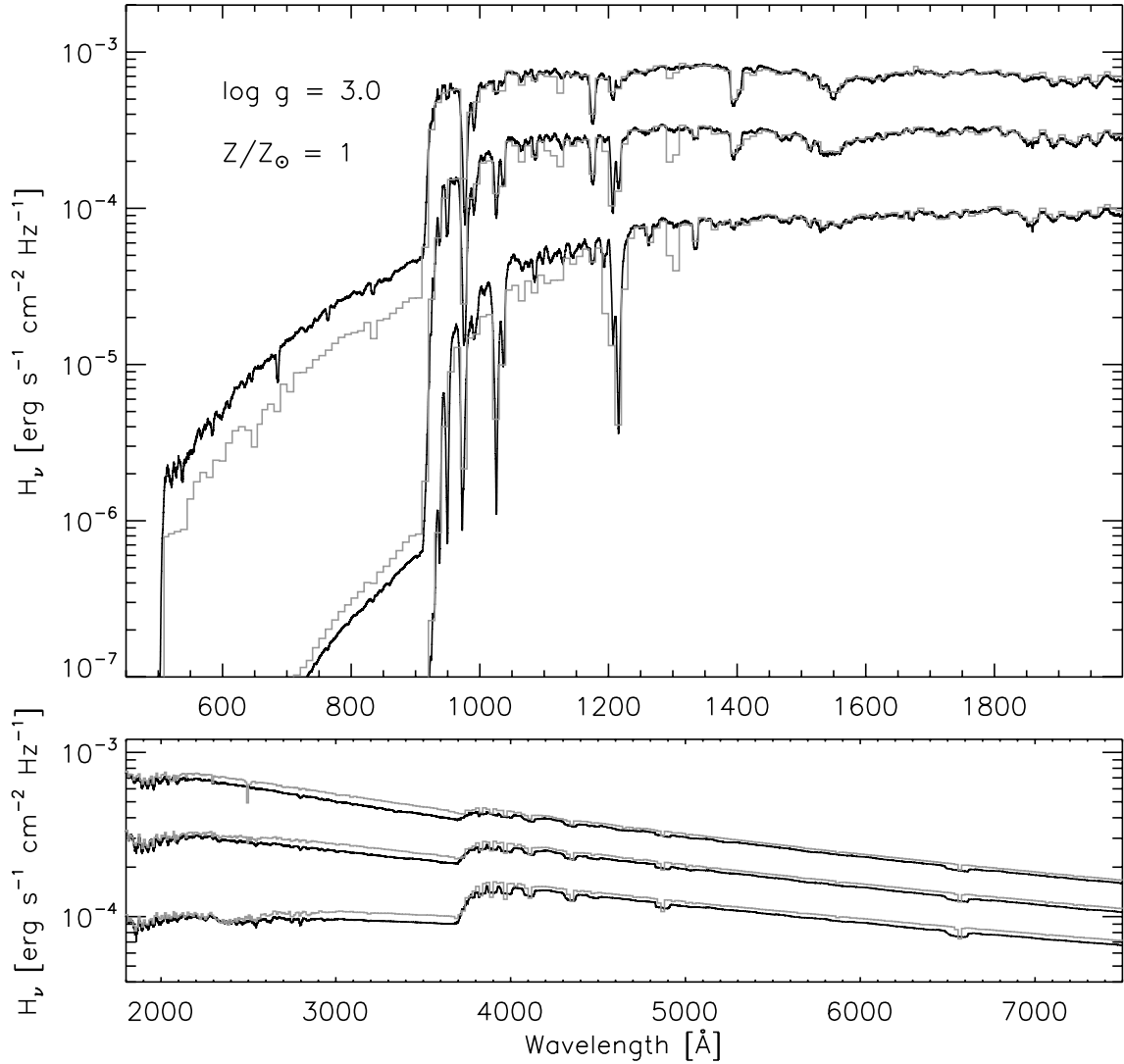


FIG. 6.—Predicted flux for three solar composition model atmospheres with $(T_{\text{eff}}, \log g)$ equal to (25,000 K, 3.0); (20,000 K, 3.0); and (15,000 K, 3.0) (black lines); compared to Kurucz (1993) models with the same parameters (gray histograms). [See the electronic edition of the Supplement for a color version of this figure.]

departures from LTE are small, the LTE and NLTE atmospheric temperature structures are very similar. This agreement thus provides good support to the two independent approaches (Kurucz ODFs and our OS treatment) to incorporate the opacity of millions of lines in model stellar atmospheres. In shallower layers, the local temperature in NLTE models is higher than in the LTE models. In the NLTE models, temperature is basically determined by a balance between heating by the Balmer hydrogen lines and cooling by lines of heavy elements (carbon and heavier elements). Figure 5 illustrates that the classical NLTE temperature rise is not completely removed by the cooling from metal lines, even at solar metallicity, and shows the metallicity dependence of this effect. We may therefore expect from this comparison that LTE and NLTE *continuum* spectra will not differ too much, while the core of strong lines and lines from minor ions will be most affected by departures from LTE.

We then compare the predicted spectra of our NLTE line-blanketed model atmospheres with Kurucz (1993) LTE model spectra. Figure 6 displays a comparison of the predicted LTE and NLTE spectral energy distributions for three solar composition models with $\log g = 3.0$ and $V_t = 2 \text{ km s}^{-1}$. For this comparison, we use the spectra directly calculated by TLUSTY (files model.flux described above), but not detailed spectra com-

puted by SYNSPEC. For clarity, the spectra are smoothed over 800 frequency points; this roughly simulates the 10 Å resolution of Kurucz (1993) model spectra. This figure reveals some differences in the continuum fluxes, most noticeably in the near ultraviolet, where the LTE fluxes are about 10% higher than the NLTE predictions. The lower NLTE fluxes result from the overpopulation of the $H \text{ I } n = 2$ level at the depth of formation of the continuum flux, hence implying a larger Balmer continuum opacity. A smaller difference is seen in the Paschen continuum of the cooler models because of overpopulation of the $n = 3$ level. At higher surface gravities, these differences are still present albeit reduced. To conserve the total flux, the LTE models show a lower flux in the far and extreme ultraviolet. In the hottest model shown ($T_{\text{eff}} = 25,000 \text{ K}$), the LTE prediction in the Lyman continuum is a factor of 2 lower than the NLTE flux. In the cooler models, the LTE flux is lower at wavelengths shorter than $\text{Ly}\alpha$.

In the last step comparing the BSTAR2006 models to LTE models, we examine detailed line profiles calculated with SYNSPEC, using either a TLUSTY model atmosphere or a Kurucz model atmosphere ($T_{\text{eff}} = 20,000 \text{ K}$, $\log g = 3.0$, solar composition, and $V_t = 2 \text{ km s}^{-1}$). We have convolved the spectra with a rotation broadening (10 km s^{-1}) and normalized the spectra to the continuum. To compute the spectra, we have used the same atomic line

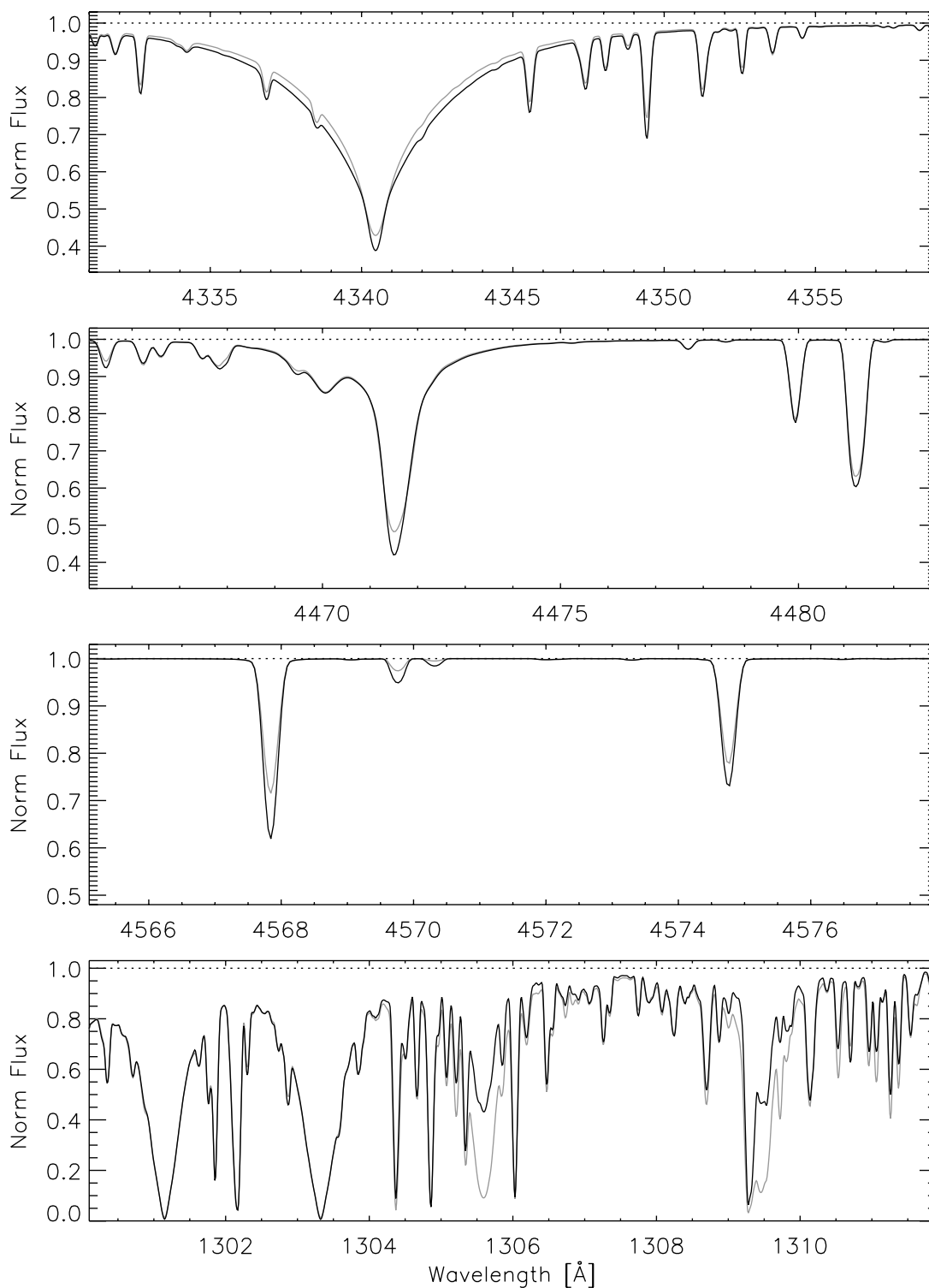


FIG. 7.—Predicted line profiles for a solar composition model atmospheres with $T_{\text{eff}} = 20,000$ K, $\log g = 3.0$, and $V_t = 2$ km s $^{-1}$ (black lines); compared to the Kurucz (1993) model with the same parameters (gray lines). [See the electronic edition of the Supplement for a color version of this figure.]

list, hence all differences result from level populations departing from the LTE value or differences in the atmospheric model structure. Some typical behaviors are illustrated in Figure 7. In the NLTE model, the hydrogen Balmer lines are broader and stronger. For the main part, this is the result of the overpopulation of the $n = 2$ level. Stellar analyses relying on LTE model atmospheres therefore tend to overestimate surface gravities derived from Balmer line wings. Other optical lines differ little (e.g., O II lines,

or Al III $\lambda 4480$) or are somewhat stronger in the NLTE model spectrum (e.g., He I $\lambda 4471$, Mg II $\lambda 4481$, or Si III $\lambda\lambda 4552, 4567, 4575$). In the latter case, this implies that the abundance of some species might be overestimated from LTE predictions; for instance, a LTE analysis of Ne I lines overestimates the neon abundance up to 0.5 dex (Cunha et al. 2006). Finally, Figure 7 shows a comparison in the far ultraviolet around 1300 Å, illustrating the effect of NLTE radiative overionization. Lines of the dominant

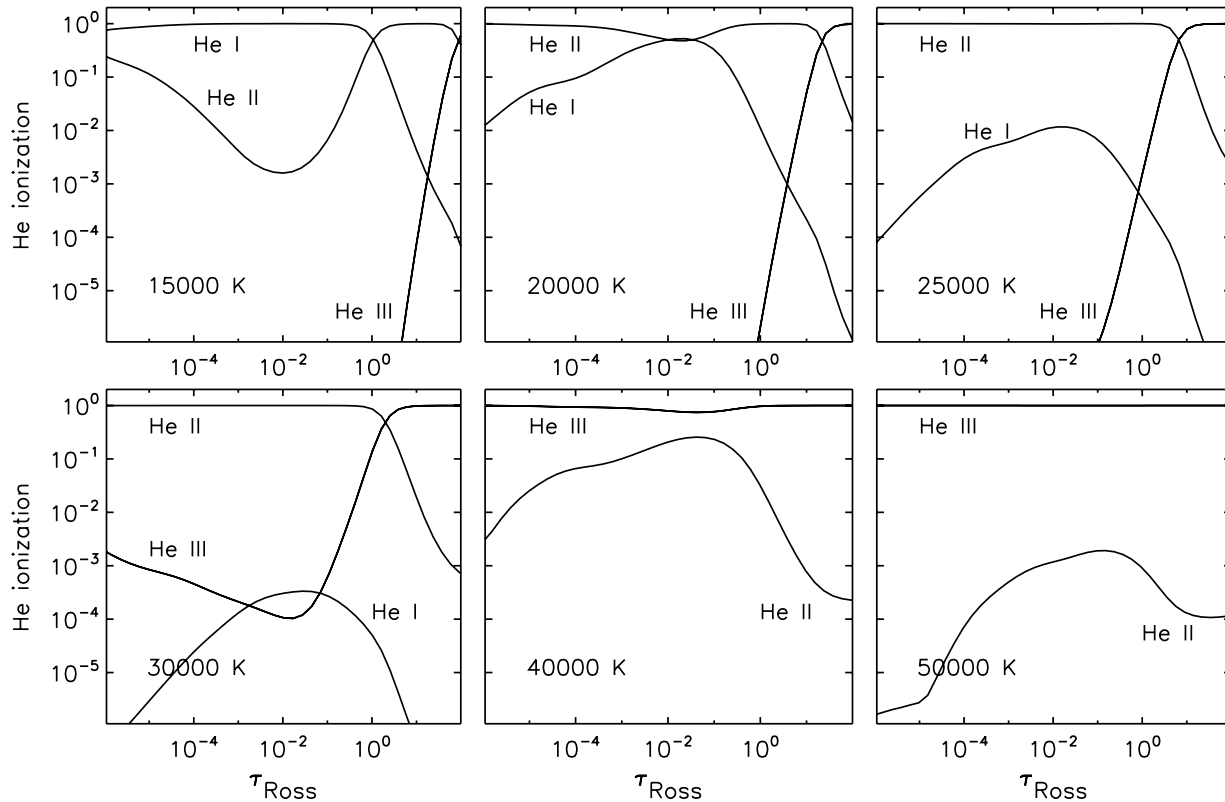


FIG. 8.—Ionization fractions of helium as function of depth in six NLTE model atmospheres, $T_{\text{eff}} = 15,000, 20,000, 25,000, 30,000, 40,000$, and $50,000$ K, $\log g = 4.0$, and solar composition. [See the electronic edition of the Supplement for a color version of this figure.]

ions, such as Si III $\lambda\lambda 1301, 1303$, remain virtually unchanged, while lines of minor ions (e.g., Si II $\lambda\lambda 1305, 1309$) are significantly weaker in the NLTE model. The weak Si II lines mainly result from the typical radiative overionization found in NLTE models; differences in atmospheric structure have a limited influence. While using Kurucz LTE models might have been a reasonable choice for analyzing UV and optical spectra of B-type stars, these comparisons demonstrate that the new BSTAR2006 NLTE line-blanketed model atmospheres represent a significant progress for determining the stellar parameters and the surface chemical composition of these stars. We plan to present a systematic study of various effects and differences between LTE and NLTE models, and their influence on deduced stellar parameters, in a future paper.

5.2. Ionization Fractions

Figures 8–16 display the ionization fractions of all explicit species (with the exception of hydrogen) in model atmospheres spanning the range of temperature of the BSTAR2006 grid. For completion, we have added two hotter O star model atmospheres ($T_{\text{eff}} = 40,000$ and $50,000$ K), thus providing a comprehensive view of the change in ionization from 15,000 to 50,000 K.⁸ We show the ionization structure of solar composition, main-sequence ($\log g = 4.0$) star atmospheres. The six selected model atmospheres roughly correspond to spectral types B5, B2.5, B1, B0, O5, and O2–O3. These figures thus give a straightforward illustration of the change in the expected line strength for various ions along the spectral type sequence, and support our choice of explicit NLTE ions. At depth ($\tau_{\text{Ross}} \gtrsim 10$), the ionization frac-

tions may become incorrect when higher ions have been neglected, most particularly for magnesium and aluminum. We stress, however, that this restriction does not affect the predicted spectra at all. Figure 17 compares the oxygen and iron ionization in model atmospheres with increasing surface gravities. At $T_{\text{eff}} = 20,000$ K, the ionization shifts up by one degree between a main-sequence model ($\log g = 4.25$) and a supergiant model ($\log g = 2.25$). Higher ionization at lower gravities is a well-known result, which follows from the Saha formula and from the typical NLTE radiative overionization. Finally, a comparison of the ionization structure of models with low and high microturbulent velocities shows that ionization is very slightly higher in models with the higher microturbulence. There is virtually no change in the dominant ion; with respect to this ion, lower ions are less populated and higher ions are more populated in models with the higher microturbulence. We interpret this change as a consequence of the stronger line blanketing. The enhanced line blocking results in slightly higher temperature at depth because of backwarming. The ions are thus exposed to a harder radiation field, hence the slight ionization shift.

These ionization fractions are useful to roughly assess the expected strengths of individual spectral lines. However, the ionization fractions should not be taken too literally. In some important cases, lines may originate from very high-lying levels whose populations may be a tiny fraction of the total abundance of a specific ion. In those cases, relying on the ionization fractions displayed in Figures 8–16 alone may be misleading. The interested reader may, however, always access the individual level populations that are stored in files `model.7`, as well as the individual NLTE departure coefficients stored in files `model.12` to estimate better the expected line strengths, or the ratios of line strengths for lines of different ions of the same species.

⁸ Fig. 7 of Paper I displaying the silicon ionization fraction in O stars was unfortunately drawn incorrectly; Fig. 14 shows the correct silicon ionization.

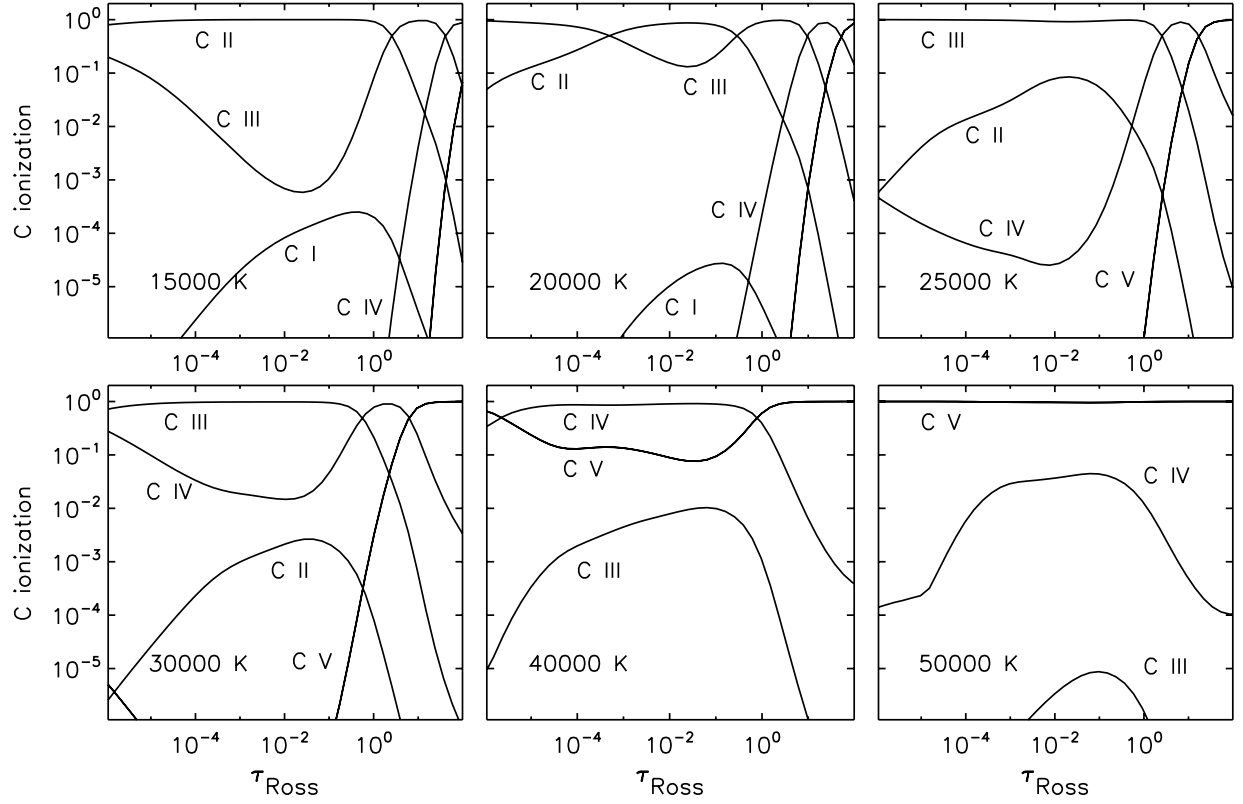


FIG. 9.—Ionization fractions of carbon as function of depth in six NLTE model atmospheres, $T_{\text{eff}} = 15,000, 20,000, 25,000, 30,000, 40,000,$ and $50,000$ K, $\log g = 4.0$, and solar composition. [See the electronic edition of the Supplement for a color version of this figure.]

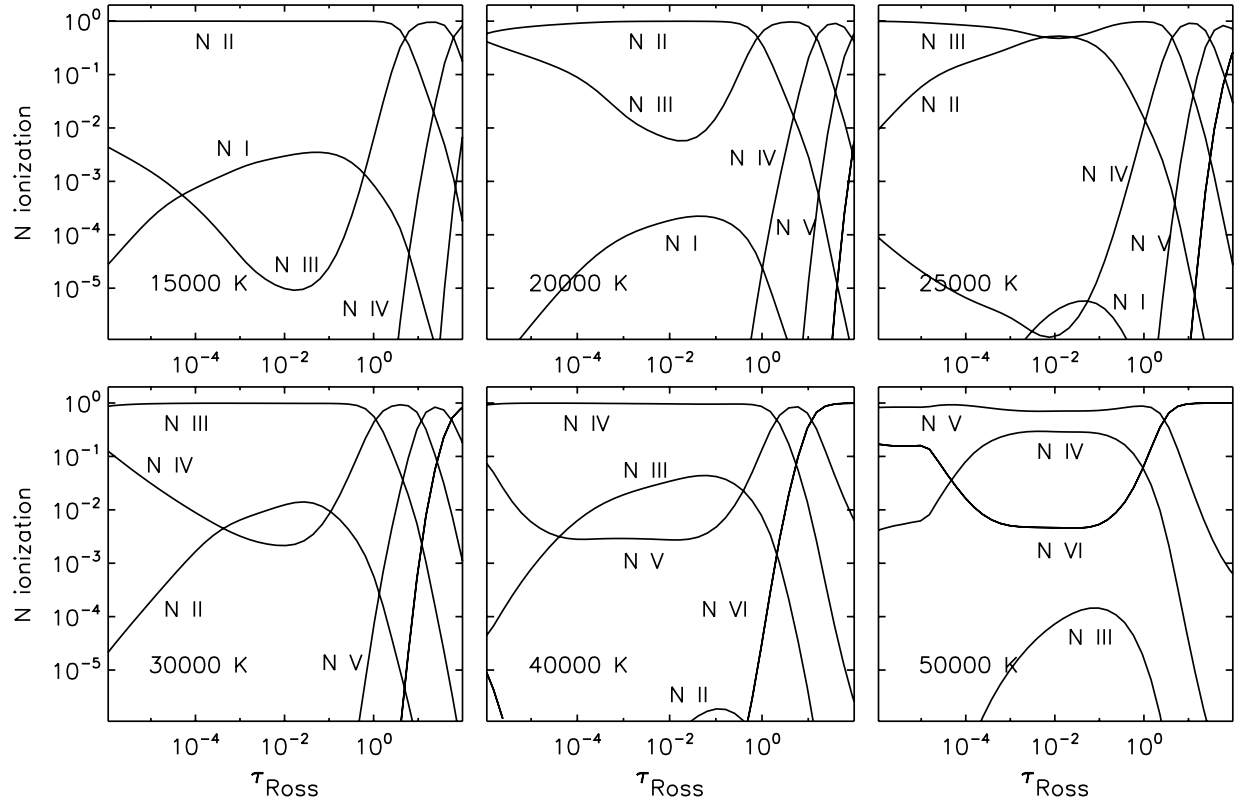


FIG. 10.—Ionization fractions of nitrogen as function of depth in six NLTE model atmospheres, $T_{\text{eff}} = 15,000, 20,000, 25,000, 30,000, 40,000,$ and $50,000$ K, $\log g = 4.0$, and solar composition. [See the electronic edition of the Supplement for a color version of this figure.]

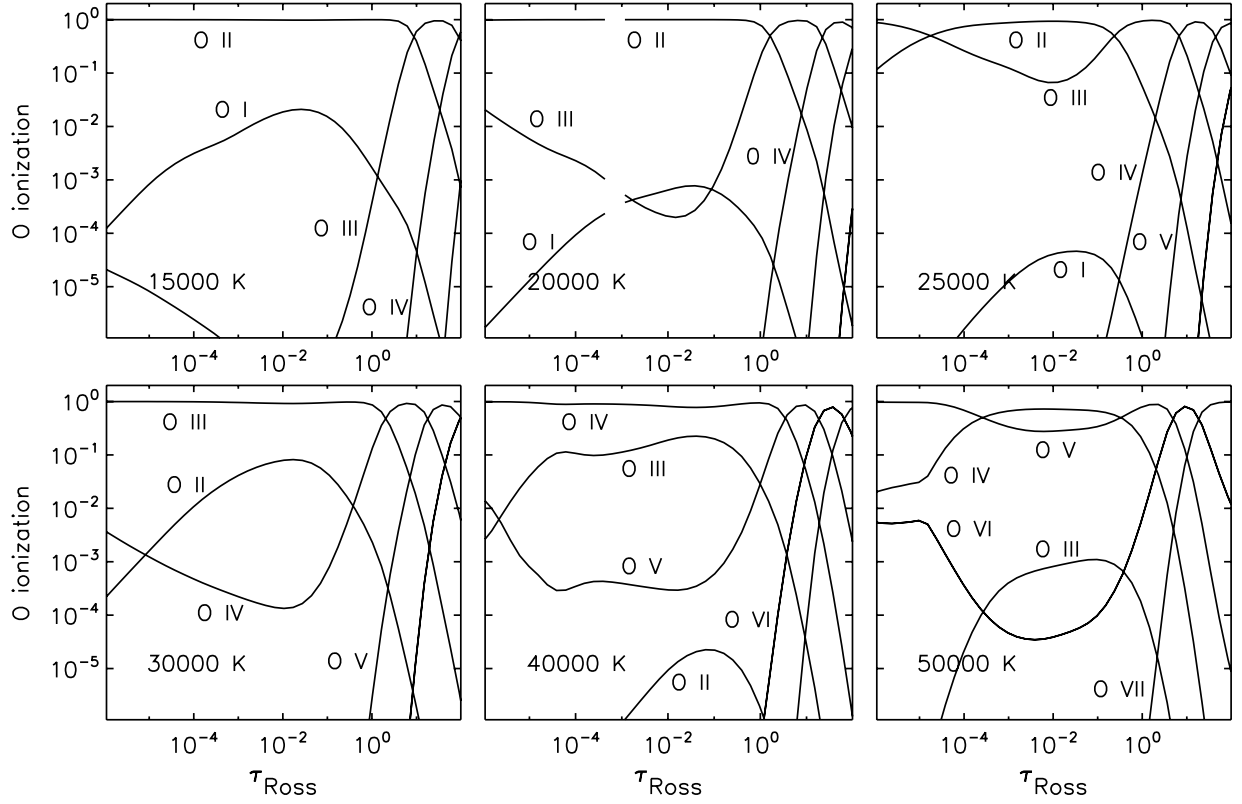


FIG. 11.—Ionization fractions of oxygen as function of depth in six NLTE model atmospheres, $T_{\text{eff}} = 15,000, 20,000, 25,000, 30,000, 40,000,$ and $50,000$ K, $\log g = 4.0$, and solar composition. [See the electronic edition of the Supplement for a color version of this figure.]

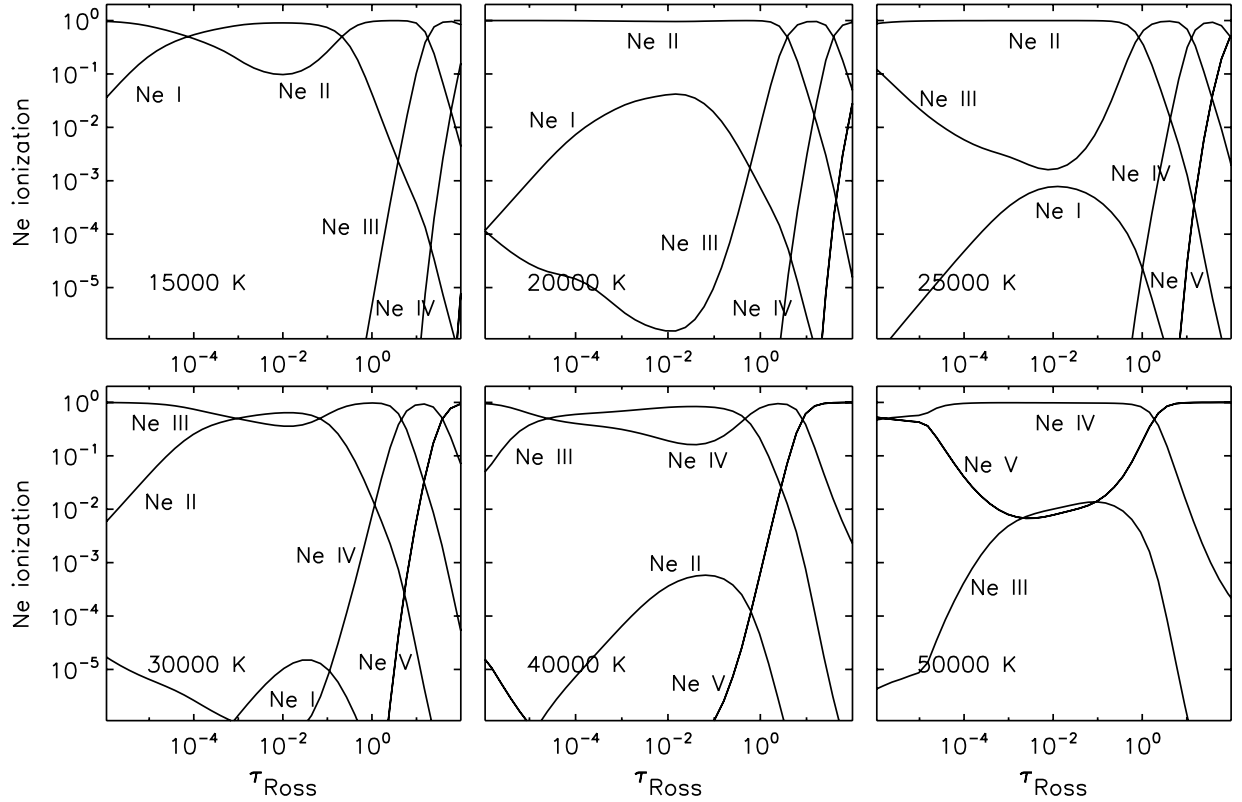


FIG. 12.—Ionization fractions of neon as function of depth in six NLTE model atmospheres, $T_{\text{eff}} = 15,000, 20,000, 25,000, 30,000, 40,000,$ and $50,000$ K, $\log g = 4.0$, and solar composition. [See the electronic edition of the Supplement for a color version of this figure.]

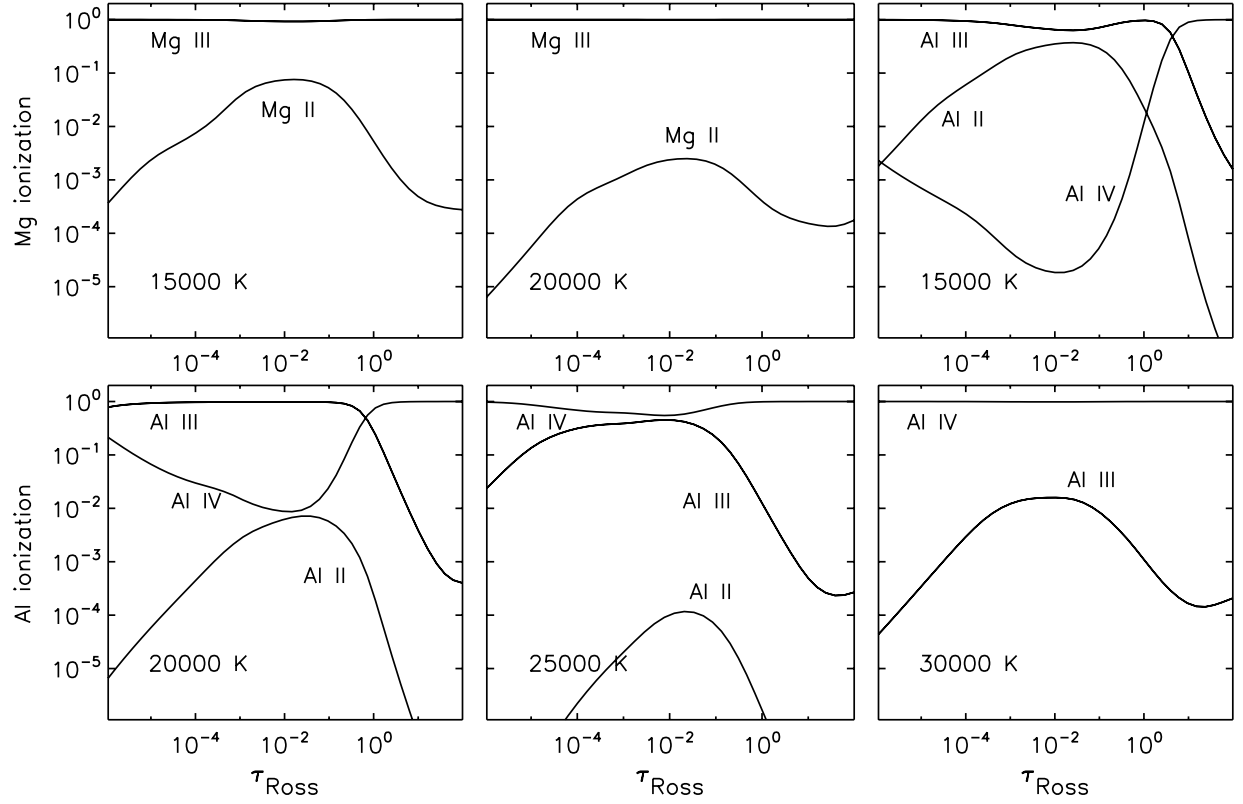


FIG. 13.—Ionization fractions of magnesium and aluminum as function of depth in four NLTE model atmospheres, $T_{\text{eff}} = 15,000, 20,000, 25,000$, and $30,000$ K, $\log g = 4.0$, and solar composition. [See the electronic edition of the Supplement for a color version of this figure.]

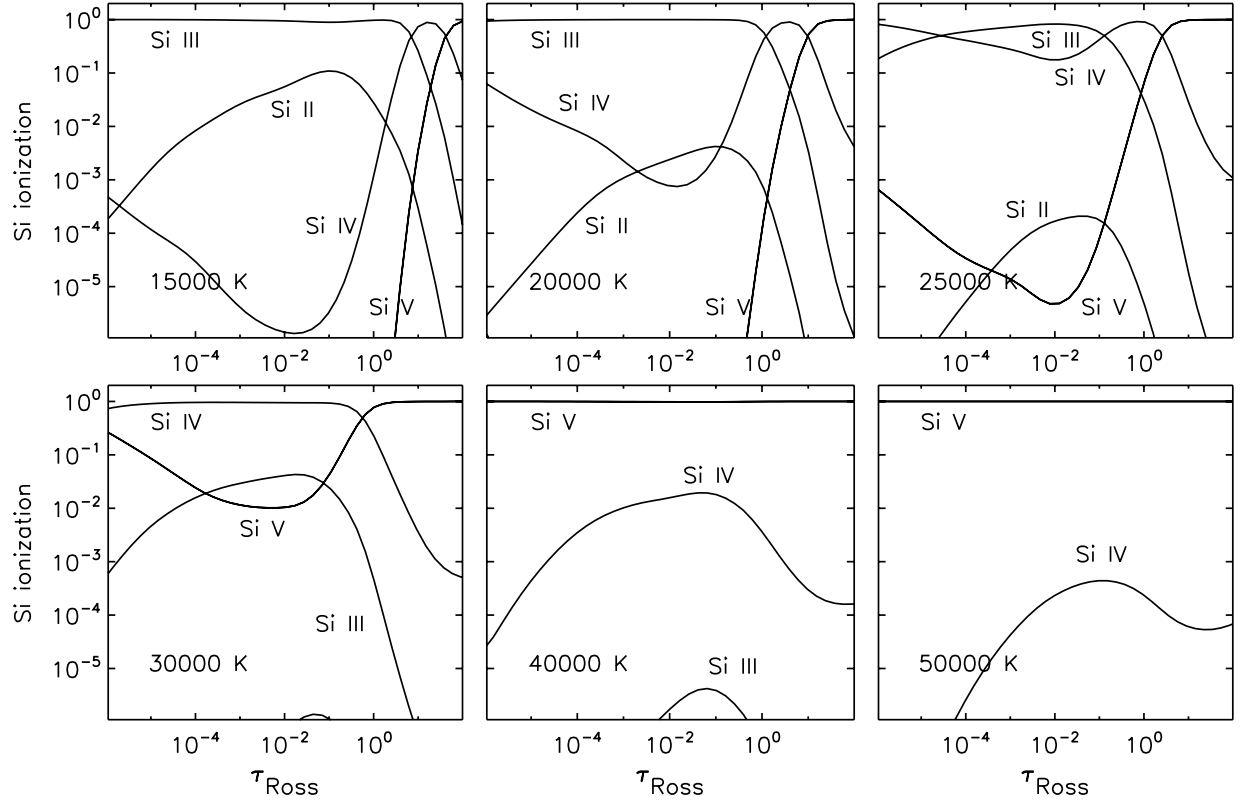


FIG. 14.—Ionization fractions of silicon as function of depth in six NLTE model atmospheres, $T_{\text{eff}} = 15,000, 20,000, 25,000, 30,000, 40,000$, and $50,000$ K, $\log g = 4.0$, and solar composition. [See the electronic edition of the Supplement for a color version of this figure.]

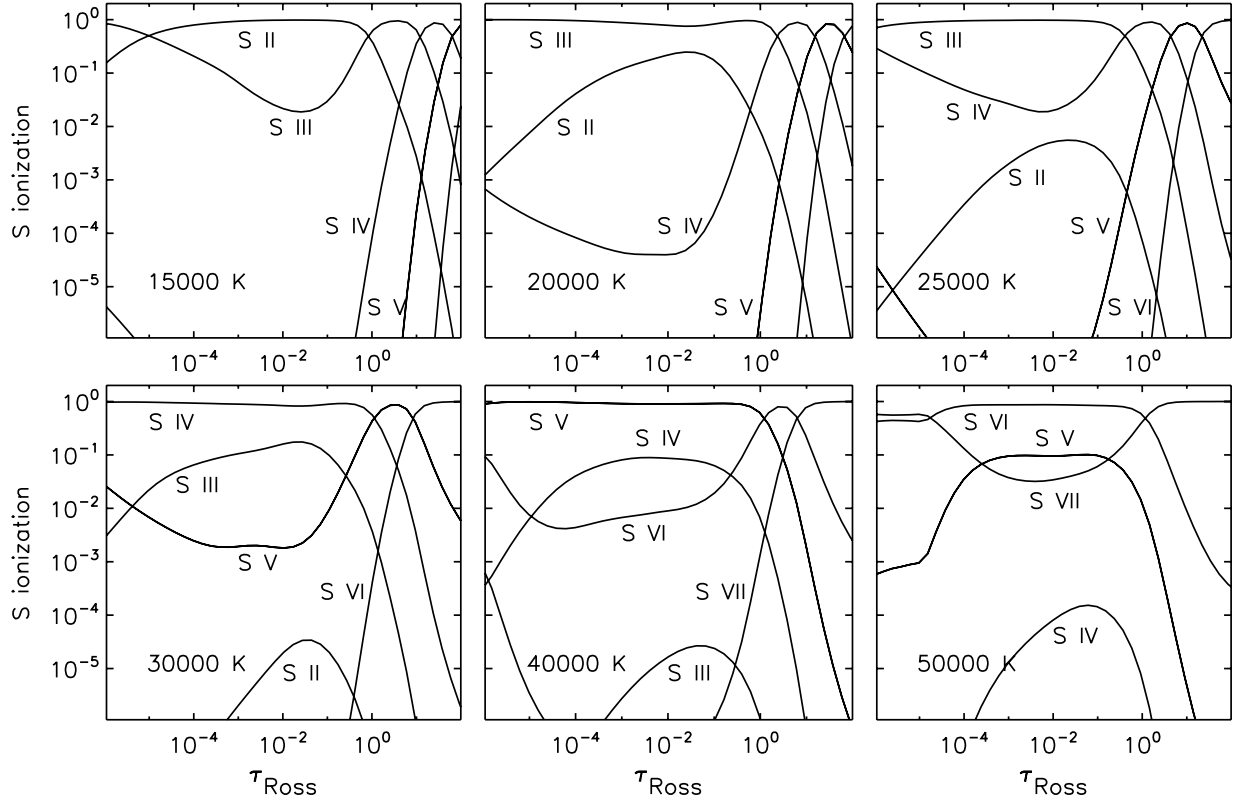


FIG. 15.—Ionization fractions of sulfur as function of depth in six NLTE model atmospheres, $T_{\text{eff}} = 15,000, 20,000, 25,000, 30,000, 40,000,$ and $50,000$ K, $\log g = 4.0$, and solar composition. [See the electronic edition of the Supplement for a color version of this figure.]

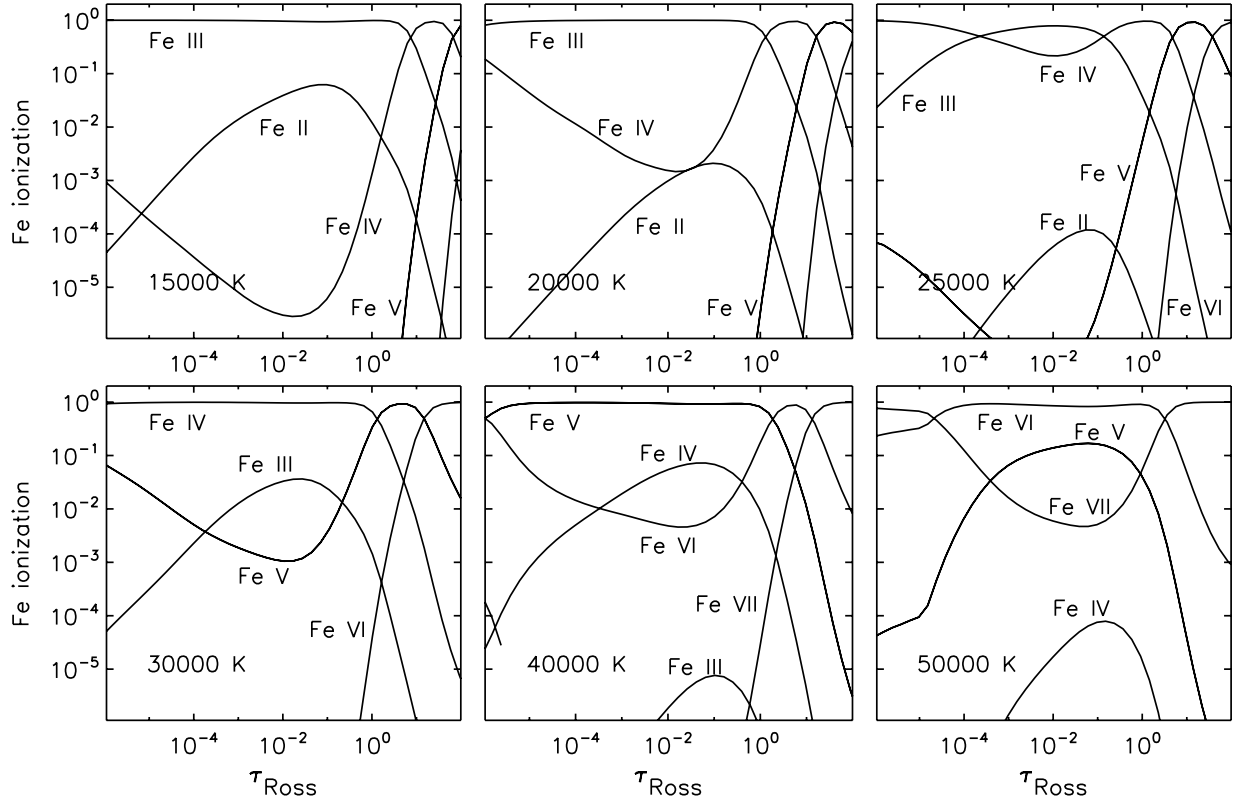


FIG. 16.—Ionization fractions of iron as function of depth in six NLTE model atmospheres, $T_{\text{eff}} = 15,000, 20,000, 25,000, 30,000, 40,000,$ and $50,000$ K, $\log g = 4.0$, and solar composition. [See the electronic edition of the Supplement for a color version of this figure.]

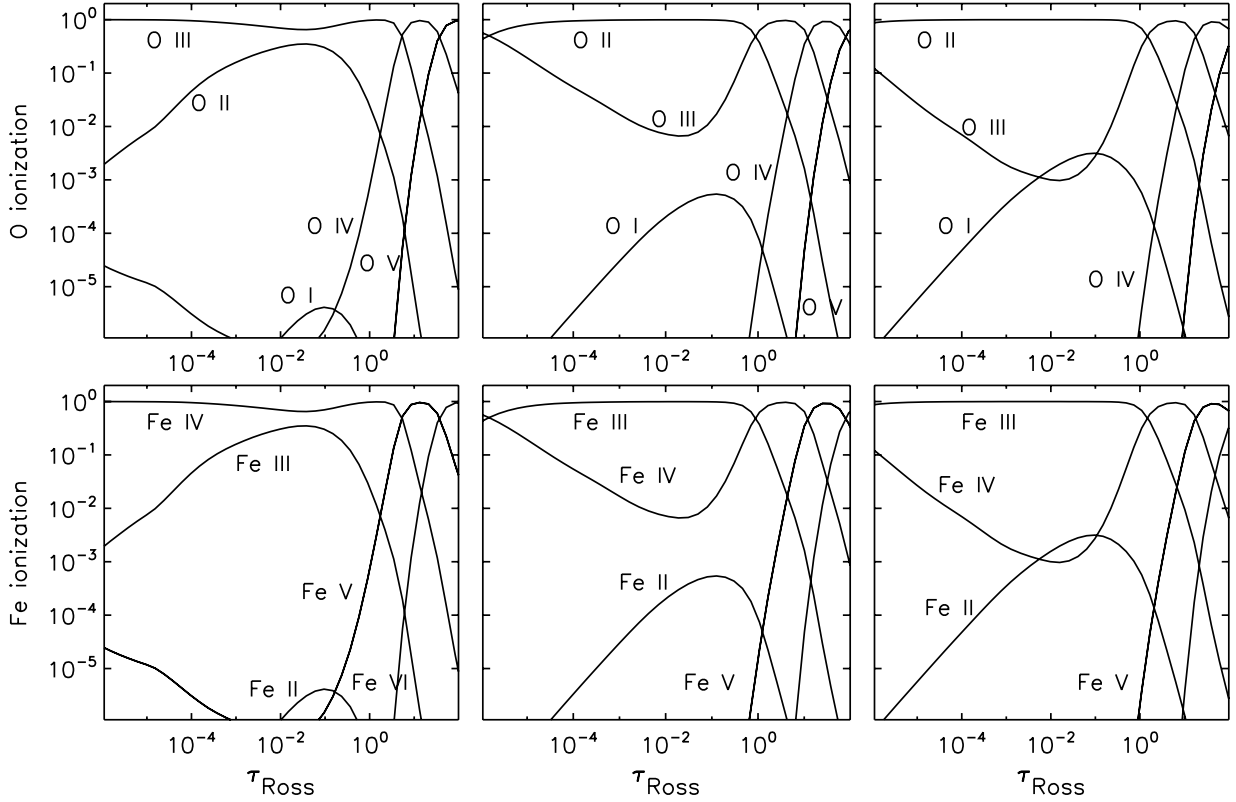


FIG. 17.—Ionization fractions of oxygen and iron as functions of depth in 3 NLTE model atmospheres with different surface gravities ($\log g = 2.25, 3.25, 4.25$; from left to right). Solar composition and $T_{\text{eff}} = 20,000$ K is assumed for the three models. [See the electronic edition of the Supplement for a color version of this figure.]

5.3. Bolometric Corrections and Ionizing Fluxes

Bolometric corrections (Table 3) are calculated using the following expression:

$$\text{BC}_V = m_{\text{bol}} - V = (-2.5 \log F_{\text{bol}} - 11.487) - (-2.5 \log F_V - 21.100), \quad (1)$$

where F_{bol} is the bolometric flux and F_V is the flux through the Johnson V filter, computed from the model atmospheres' flux distribution. The bolometric flux is computed by trapezoidal integration over the complete frequency range, while we have used an IDL version of the program *ubvuser* (Kurucz 1993) kindly made available by Wayne Landsman to obtain the V magnitudes. The first constant is defined by assuming a solar value, $\text{BC}_V^{\odot} = -0.07$, while the second constant defines the zero point of the V magnitude scale.

Figure 18 illustrates the major dependence of the bolometric correction with the effective temperature (solar metallicity, $\log g = 4.0$). We have extended the plot to encompass both the BSTAR2006 and the OSTAR2002 models. Dependences in terms of the other stellar parameters (gravity, metallicity, and microturbulent velocity) are much weaker than the temperature dependence. While we find differences up to 0.1 mag over the whole range in metallicity (twice solar to metal-free), the changes remain limited to a few hundredths of a magnitude for different gravities and microturbulences. Larger differences are found for models very close to the Eddington limit, but these models are in a domain where our basic assumption of static atmospheres starts to break up. The small step (0.05 dex) in the overlapping range between the two grids is therefore most likely a consequence of the differences in the treatment of opacities, that is, it likely results

from the different ions, levels, and lines explicitly included in the model atmospheres. The bolometric corrections of “CN” models do not differ from models with unaltered compositions, and they are thus not listed in Table 3.

Ionizing photons are emitted essentially by hot, massive stars, and the extreme ultraviolet flux drops precipitously in early B stars with diminishing effective temperatures. Therefore, we only provide in Table 4 a summarized extension of the OSTAR2002 Lyman continuum fluxes to the early B star domain.

5.4. Radiative Acceleration and Effective Gravities

The comprehensive treatment of opacities in the BSTAR2006 model atmospheres allows for an accurate estimate of radiative pressure on the atmospheric structure. TLUSTY thus provides a run of radiative acceleration with depth, see Figure 2 in Paper I for an illustration. In every model atmosphere, the radiative acceleration goes through a local maximum in the continuum forming region ($\tau_{\text{Ross}} \approx 1$), and then shows another strong increase in very superficial layers ($\tau_{\text{Ross}} \lesssim 10^{-6}$). In low-gravity models, the radiative acceleration may exceed gravity at these depths indicating that the top layers are unstable. The assumption of hydrostatic equilibrium then break down there. We have numerically limited the radiative acceleration in these superficial layers to ensure convergence following the prescription used in Paper I. The model spectra are not affected by this approximation, because these superficial layers only influence the strong resonance lines that form in a stellar wind. The photospheric stability is more appropriately defined by the maximum reached by the radiative acceleration around optical depth unity (see Fig. 2, Paper I). Figure 19 displays isocontours of this maximum relative to the gravitational acceleration, $\Gamma_{\text{rad}} = \max(g_{\text{rad}})/g$, as a function of effective temperature and

TABLE 3
BOLOMETRIC CORRECTIONS AS FUNCTION OF EFFECTIVE TEMPERATURE, GRAVITY, MICROTURBULENT VELOCITY, AND METALLICITY

BOLOMETRIC CORRECTION (mag)													
T_{eff} (K)	$\log g$	$V_{\text{turb}} = 2 \text{ km s}^{-1}$						$V_{\text{turb}} = 10 \text{ km s}^{-1}$					
		$Z/Z_{\odot} = 2.0$	$Z/Z_{\odot} = 1.0$	$Z/Z_{\odot} = 0.5$	$Z/Z_{\odot} = 0.2$	$Z/Z_{\odot} = 0.1$	$Z/Z_{\odot} = 0.0$	$Z/Z_{\odot} = 2.0$	$Z/Z_{\odot} = 1.0$	$Z/Z_{\odot} = 0.5$	$Z/Z_{\odot} = 0.2$	$Z/Z_{\odot} = 0.1$	$Z/Z_{\odot} = 0.0$
15,000.....	1.75	-1.29	-1.32	-1.33	-1.35	-1.36	-1.39	-1.26	-1.29	-1.32	-1.34	-1.36	-1.40
	2.00	-1.23	-1.26	-1.28	-1.30	-1.31	-1.33	-1.19	-1.23	-1.26	-1.29	-1.30	-1.34
	2.25	-1.21	-1.24	-1.26	-1.28	-1.29	-1.31	-1.17	-1.21	-1.24	-1.27	-1.28	-1.32
	2.50	-1.20	-1.23	-1.26	-1.28	-1.29	-1.30	-1.16	-1.20	-1.23	-1.26	-1.28	-1.31
	2.75	-1.20	-1.23	-1.25	-1.27	-1.28	-1.30	-1.16	-1.20	-1.23	-1.26	-1.27	-1.30
	3.00	-1.20	-1.23	-1.25	-1.27	-1.28	-1.30	-1.15	-1.20	-1.23	-1.26	-1.27	-1.30
	3.25	-1.20	-1.23	-1.25	-1.27	-1.28	-1.30
	3.50	-1.19	-1.23	-1.25	-1.27	-1.28	-1.30
	3.75	-1.19	-1.22	-1.25	-1.27	-1.28	-1.29
	4.00	-1.19	-1.22	-1.25	-1.27	-1.28	-1.29
	4.25	-1.19	-1.22	-1.24	-1.26	-1.27	-1.29
	4.50	-1.19	-1.22	-1.24	-1.26	-1.27	-1.29
	4.75	-1.19	-1.22	-1.24	-1.26	-1.27	-1.29
	2.00	-1.39	-1.42	-1.44	-1.46	-1.47	-1.50	-1.34	-1.38	-1.41	-1.44	-1.46	-1.50
	2.25	-1.37	-1.40	-1.42	-1.44	-1.45	-1.47	-1.32	-1.36	-1.39	-1.42	-1.44	-1.48
16,000.....	2.50	-1.36	-1.39	-1.41	-1.43	-1.44	-1.46	-1.31	-1.35	-1.39	-1.42	-1.43	-1.47
	2.75	-1.36	-1.39	-1.41	-1.43	-1.44	-1.46	-1.31	-1.35	-1.39	-1.42	-1.43	-1.46
	3.00	-1.36	-1.39	-1.41	-1.43	-1.44	-1.46	-1.31	-1.36	-1.39	-1.42	-1.43	-1.46
	3.25	-1.36	-1.39	-1.41	-1.43	-1.44	-1.46
	3.50	-1.36	-1.39	-1.41	-1.43	-1.44	-1.46
	3.75	-1.36	-1.39	-1.41	-1.43	-1.44	-1.46
	4.00	-1.36	-1.39	-1.41	-1.43	-1.44	-1.46
	4.25	-1.36	-1.39	-1.41	-1.43	-1.44	-1.46
	4.50	-1.36	-1.39	-1.41	-1.43	-1.44	-1.45
	4.75	-1.35	-1.38	-1.41	-1.43	-1.44	-1.45
	2.00	-1.54	-1.57	-1.59	-1.61	-1.62	-1.66	-1.51	-1.53	-1.56	-1.59	-1.61	-1.67
	2.25	-1.51	-1.54	-1.56	-1.58	-1.59	-1.62	-1.45	-1.49	-1.53	-1.56	-1.58	-1.63
	2.50	-1.50	-1.53	-1.56	-1.58	-1.59	-1.61	-1.45	-1.49	-1.52	-1.56	-1.57	-1.61
	2.75	-1.50	-1.53	-1.56	-1.58	-1.59	-1.61	-1.45	-1.50	-1.53	-1.56	-1.57	-1.61
	3.00	-1.51	-1.54	-1.56	-1.58	-1.59	-1.61	-1.46	-1.50	-1.53	-1.56	-1.58	-1.61
17,000.....	3.25	-1.51	-1.54	-1.56	-1.58	-1.59	-1.61
	3.50	-1.51	-1.54	-1.57	-1.58	-1.59	-1.61
	3.75	-1.51	-1.54	-1.57	-1.59	-1.59	-1.61
	4.00	-1.52	-1.55	-1.57	-1.59	-1.60	-1.61
	4.25	-1.52	-1.55	-1.57	-1.59	-1.60	-1.61
	4.50	-1.52	-1.55	-1.57	-1.59	-1.59	-1.61
	4.75	-1.51	-1.54	-1.57	-1.58	-1.59	-1.61

TABLE 3—*Continued*

BOLOMETRIC CORRECTION (mag)													
T_{eff} (K)	$\log g$	$V_{\text{turb}} = 2 \text{ km s}^{-1}$						$V_{\text{turb}} = 10 \text{ km s}^{-1}$					
		$Z/Z_{\odot} = 2.0$	$Z/Z_{\odot} = 1.0$	$Z/Z_{\odot} = 0.5$	$Z/Z_{\odot} = 0.2$	$Z/Z_{\odot} = 0.1$	$Z/Z_{\odot} = 0.0$	$Z/Z_{\odot} = 2.0$	$Z/Z_{\odot} = 1.0$	$Z/Z_{\odot} = 0.5$	$Z/Z_{\odot} = 0.2$	$Z/Z_{\odot} = 0.1$	$Z/Z_{\odot} = 0.0$
18,000.....	2.00	-1.72	-1.74	-1.75	-1.77	-1.79	-1.83	-1.76	-1.78	-1.80	-1.85
	2.25	-1.64	-1.67	-1.69	-1.72	-1.73	-1.77	-1.58	-1.62	-1.65	-1.69	-1.71	-1.77
	2.50	-1.63	-1.67	-1.69	-1.71	-1.72	-1.75	-1.57	-1.61	-1.65	-1.69	-1.70	-1.76
	2.75	-1.64	-1.67	-1.69	-1.71	-1.72	-1.75	-1.58	-1.62	-1.66	-1.69	-1.71	-1.75
	3.00	-1.64	-1.68	-1.70	-1.72	-1.73	-1.75	-1.59	-1.63	-1.67	-1.70	-1.71	-1.75
	3.25	-1.65	-1.68	-1.70	-1.72	-1.73	-1.75
	3.50	-1.65	-1.69	-1.71	-1.73	-1.74	-1.75
	3.75	-1.66	-1.69	-1.71	-1.73	-1.74	-1.76
	4.00	-1.66	-1.69	-1.71	-1.73	-1.74	-1.76
	4.25	-1.66	-1.69	-1.72	-1.73	-1.74	-1.76
	4.50	-1.66	-1.69	-1.72	-1.73	-1.74	-1.76
	4.75	-1.66	-1.69	-1.72	-1.73	-1.74	-1.75
19,000.....	2.25	-1.77	-1.79	-1.82	-1.84	-1.86	-1.90	-1.73	-1.76	-1.79	-1.82	-1.84	-1.91
	2.50	-1.75	-1.79	-1.81	-1.84	-1.85	-1.89	-1.69	-1.73	-1.77	-1.81	-1.83	-1.89
	2.75	-1.76	-1.79	-1.82	-1.84	-1.85	-1.88	-1.69	-1.74	-1.78	-1.81	-1.83	-1.88
	3.00	-1.77	-1.80	-1.83	-1.85	-1.86	-1.88	-1.71	-1.76	-1.79	-1.82	-1.84	-1.88
	3.25	-1.78	-1.81	-1.83	-1.85	-1.86	-1.89
	3.50	-1.79	-1.82	-1.84	-1.86	-1.87	-1.89
	3.75	-1.79	-1.82	-1.85	-1.86	-1.87	-1.89
	4.00	-1.80	-1.83	-1.85	-1.87	-1.88	-1.89
	4.25	-1.80	-1.83	-1.85	-1.87	-1.88	-1.89
	4.50	-1.80	-1.83	-1.85	-1.87	-1.88	-1.89
	4.75	-1.81	-1.83	-1.86	-1.87	-1.88	-1.89
	20,000.....	2.25	-1.91	-1.93	-1.95	-1.98	-1.99	-2.04	-1.91	-1.92	-1.94	-1.97	-1.99
2.50		-1.87	-1.90	-1.93	-1.95	-1.97	-2.01	-1.81	-1.85	-1.88	-1.92	-1.95	-2.01
2.75		-1.87	-1.91	-1.93	-1.96	-1.97	-2.01	-1.80	-1.85	-1.89	-1.93	-1.95	-2.01
3.00		-1.89	-1.92	-1.95	-1.97	-1.98	-2.01	-1.82	-1.87	-1.90	-1.94	-1.96	-2.01
3.25		-1.90	-1.93	-1.96	-1.98	-1.99	-2.01
3.50		-1.91	-1.94	-1.96	-1.98	-1.99	-2.02
3.75		-1.92	-1.95	-1.97	-1.99	-2.00	-2.02
4.00		-1.92	-1.96	-1.97	-2.00	-2.00	-2.02
4.25		-1.93	-1.96	-1.98	-2.00	-2.01	-2.02
4.50		-1.93	-1.96	-1.99	-2.00	-2.01	-2.03
4.75		-1.94	-1.97	-1.99	-2.00	-2.01	-2.02
21,000.....		2.25	-2.16	-2.17	-2.22	-2.16	-2.18
	2.50	-1.98	-2.01	-2.04	-2.06	-2.08	-2.12	-1.94	-1.97	-2.01	-2.04	-2.07	-2.13
	2.75	-1.98	-2.01	-2.04	-2.07	-2.09	-2.13	-1.92	-1.96	-2.00	-2.04	-2.06	-2.13
	3.00	-1.99	-2.03	-2.06	-2.08	-2.10	-2.13	-1.92	-1.97	-2.01	-2.05	-2.07	-2.13
	3.25	-2.01	-2.04	-2.07	-2.09	-2.10	-2.14
	3.50	-2.02	-2.06	-2.08	-2.10	-2.11	-2.14
	3.75	-2.03	-2.07	-2.09	-2.11	-2.12	-2.14
	4.00	-2.04	-2.08	-2.09	-2.12	-2.13	-2.15
	4.25	-2.05	-2.08	-2.10	-2.12	-2.13	-2.15
	4.50	-2.06	-2.09	-2.11	-2.13	-2.13	-2.15
	4.75	-2.06	-2.09	-2.11	-2.13	-2.13	-2.15

TABLE 3—*Continued*

BOLOMETRIC CORRECTION (mag)													
T_{eff} (K)	$\log g$	$V_{\text{turb}} = 2 \text{ km s}^{-1}$						$V_{\text{turb}} = 10 \text{ km s}^{-1}$					
		$Z/Z_{\odot} = 2.0$	$Z/Z_{\odot} = 1.0$	$Z/Z_{\odot} = 0.5$	$Z/Z_{\odot} = 0.2$	$Z/Z_{\odot} = 0.1$	$Z/Z_{\odot} = 0.0$	$Z/Z_{\odot} = 2.0$	$Z/Z_{\odot} = 1.0$	$Z/Z_{\odot} = 0.5$	$Z/Z_{\odot} = 0.2$	$Z/Z_{\odot} = 0.1$	$Z/Z_{\odot} = 0.0$
22,000.....	2.50	-2.10	-2.12	-2.15	-2.18	-2.20	-2.24	-2.07	-2.10	-2.13	-2.17	-2.19	-2.24
	2.75	-2.08	-2.12	-2.15	-2.17	-2.19	-2.24	-2.03	-2.07	-2.11	-2.15	-2.17	-2.24
	3.00	-2.10	-2.13	-2.16	-2.19	-2.20	-2.24	-2.03	-2.07	-2.11	-2.15	-2.18	-2.24
	3.25	-2.11	-2.15	-2.18	-2.20	-2.21	-2.25
	3.50	-2.13	-2.16	-2.19	-2.21	-2.22	-2.26
	3.75	-2.14	-2.18	-2.20	-2.22	-2.23	-2.26
	4.00	-2.15	-2.19	-2.21	-2.23	-2.24	-2.26
	4.25	-2.16	-2.19	-2.22	-2.24	-2.25	-2.27
	4.50	-2.17	-2.20	-2.22	-2.24	-2.25	-2.27
	4.75	-2.18	-2.21	-2.23	-2.24	-2.25	-2.27
23,000.....	2.50	-2.24	-2.26	-2.28	-2.31	-2.32	-2.37	-2.25	-2.25	-2.27	-2.30	-2.32	-2.38
	2.75	-2.19	-2.22	-2.25	-2.28	-2.29	-2.33	-2.14	-2.18	-2.22	-2.26	-2.28	-2.33
	3.00	-2.20	-2.23	-2.26	-2.29	-2.30	-2.35	-2.14	-2.18	-2.22	-2.26	-2.28	-2.35
	3.25	-2.21	-2.25	-2.27	-2.30	-2.32	-2.36
	3.50	-2.23	-2.26	-2.29	-2.32	-2.33	-2.36
	3.75	-2.24	-2.28	-2.30	-2.33	-2.34	-2.37
	4.00	-2.26	-2.29	-2.32	-2.34	-2.35	-2.38
	4.25	-2.27	-2.30	-2.32	-2.34	-2.35	-2.38
	4.50	-2.28	-2.31	-2.33	-2.35	-2.36	-2.38
	4.75	-2.28	-2.32	-2.34	-2.35	-2.36	-2.38
24,000.....	2.50	-2.43	-2.43	-2.43	-2.45	-2.47	-2.52	-2.49	-2.46	-2.45	-2.46	-2.47	-2.53
	2.75	-2.29	-2.32	-2.35	-2.37	-2.39	-2.43	-2.25	-2.29	-2.32	-2.36	-2.38	-2.44
	3.00	-2.29	-2.33	-2.35	-2.38	-2.40	-2.44	-2.24	-2.28	-2.32	-2.36	-2.39	-2.44
	3.25	-2.31	-2.34	-2.37	-2.40	-2.41	-2.46
	3.50	-2.32	-2.36	-2.39	-2.41	-2.43	-2.47
	3.75	-2.34	-2.37	-2.40	-2.43	-2.44	-2.47
	4.00	-2.35	-2.39	-2.41	-2.44	-2.45	-2.48
	4.25	-2.37	-2.40	-2.42	-2.45	-2.46	-2.48
	4.50	-2.38	-2.41	-2.43	-2.45	-2.46	-2.49
	4.75	-2.39	-2.42	-2.44	-2.46	-2.47	-2.49
25,000.....	2.75	-2.40	-2.42	-2.45	-2.48	-2.50	-2.54	-2.38	-2.40	-2.43	-2.46	-2.49	-2.55
	3.00	-2.38	-2.42	-2.45	-2.47	-2.49	-2.53	-2.34	-2.38	-2.42	-2.46	-2.48	-2.53
	3.25	-2.40	-2.43	-2.46	-2.49	-2.51	-2.55
	3.50	-2.42	-2.45	-2.48	-2.51	-2.52	-2.56
	3.75	-2.43	-2.47	-2.49	-2.52	-2.53	-2.57
	4.00	-2.45	-2.48	-2.51	-2.53	-2.55	-2.58
	4.25	-2.46	-2.49	-2.52	-2.54	-2.55	-2.58
	4.50	-2.47	-2.50	-2.53	-2.55	-2.56	-2.59
	4.75	-2.48	-2.51	-2.54	-2.56	-2.57	-2.59

TABLE 3—*Continued*

BOLOMETRIC CORRECTION (mag)													
T_{eff} (K)	$\log g$	$V_{\text{turb}} = 2 \text{ km s}^{-1}$						$V_{\text{turb}} = 10 \text{ km s}^{-1}$					
		$Z/Z_{\odot} = 2.0$	$Z/Z_{\odot} = 1.0$	$Z/Z_{\odot} = 0.5$	$Z/Z_{\odot} = 0.2$	$Z/Z_{\odot} = 0.1$	$Z/Z_{\odot} = 0.0$	$Z/Z_{\odot} = 2.0$	$Z/Z_{\odot} = 1.0$	$Z/Z_{\odot} = 0.5$	$Z/Z_{\odot} = 0.2$	$Z/Z_{\odot} = 0.1$	$Z/Z_{\odot} = 0.0$
26,000.....	2.75	−2.52	−2.54	−2.56	−2.58	−2.60	−2.65	−2.52	−2.53	−2.55	−2.57	−2.59	−2.66
	3.00	−2.48	−2.51	−2.53	−2.56	−2.58	−2.62	−2.43	−2.47	−2.51	−2.55	−2.57	−2.62
	3.25	−2.49	−2.52	−2.55	−2.58	−2.59	−2.63
	3.50	−2.51	−2.54	−2.57	−2.60	−2.61	−2.65
	3.75	−2.52	−2.56	−2.58	−2.61	−2.62	−2.66
	4.00	−2.54	−2.57	−2.60	−2.62	−2.64	−2.67
	4.25	−2.55	−2.58	−2.61	−2.63	−2.65	−2.68
	4.50	−2.56	−2.59	−2.62	−2.64	−2.65	−2.68
27,000.....	4.75	−2.57	−2.60	−2.63	−2.65	−2.66	−2.69
	2.75	−2.68	−2.68	−2.69	−2.71	−2.72	−2.77	−2.68	−2.70	−2.70	−2.71	−2.72	−2.77
	3.00	−2.57	−2.60	−2.62	−2.65	−2.67	−2.71	−2.54	−2.57	−2.60	−2.64	−2.66	−2.72
	3.25	−2.57	−2.61	−2.63	−2.66	−2.68	−2.71
	3.50	−2.59	−2.62	−2.65	−2.68	−2.69	−2.73
	3.75	−2.61	−2.64	−2.67	−2.70	−2.71	−2.74
	4.00	−2.62	−2.66	−2.68	−2.71	−2.72	−2.76
	4.25	−2.63	−2.67	−2.69	−2.72	−2.73	−2.77
28,000.....	4.50	−2.65	−2.68	−2.70	−2.73	−2.74	−2.77
	4.75	−2.66	−2.69	−2.71	−2.74	−2.75	−2.78
	2.75	−2.86	−2.85	−2.85	−2.86	−2.86	−2.89	−2.87	...	−2.90
	3.00	−2.67	−2.69	−2.72	−2.74	−2.76	−2.81	−2.65	−2.67	−2.70	−2.73	−2.75	−2.81
	3.25	−2.66	−2.69	−2.72	−2.74	−2.76	−2.79
	3.50	−2.67	−2.71	−2.73	−2.76	−2.77	−2.80
	3.75	−2.69	−2.72	−2.75	−2.78	−2.79	−2.82
	4.00	−2.71	−2.74	−2.76	−2.79	−2.80	−2.83
29,000.....	4.25	−2.72	−2.75	−2.78	−2.80	−2.81	−2.85
	4.50	−2.73	−2.76	−2.79	−2.81	−2.82	−2.85
	4.75	−2.74	−2.77	−2.80	−2.82	−2.83	−2.86
	3.00	−2.78	−2.80	−2.82	−2.84	−2.85	−2.90	−2.78	−2.79	−2.80	−2.83	−2.85	−2.90
	3.25	−2.74	−2.77	−2.80	−2.82	−2.84	−2.88
	3.50	−2.75	−2.78	−2.81	−2.84	−2.85	−2.88
	3.75	−2.77	−2.80	−2.83	−2.85	−2.86	−2.89
	4.00	−2.78	−2.82	−2.84	−2.87	−2.88	−2.91
30,000.....	4.25	−2.80	−2.83	−2.86	−2.88	−2.89	−2.92
	4.50	−2.81	−2.84	−2.87	−2.89	−2.90	−2.93
	4.75	−2.82	−2.85	−2.87	−2.90	−2.91	−2.94
	3.00	−2.91	−2.92	−2.93	−2.95	−2.96	−2.99	−2.91	−2.92	−2.92	−2.94	−2.95	−3.00
	3.25	−2.84	−2.86	−2.88	−2.91	−2.92	−2.96
	3.50	−2.83	−2.86	−2.89	−2.91	−2.93	−2.96
	3.75	−2.85	−2.88	−2.90	−2.93	−2.94	−2.97
	4.00	−2.86	−2.89	−2.92	−2.94	−2.95	−2.98
	4.25	−2.87	−2.91	−2.93	−2.95	−2.96	−2.99
	4.50	−2.89	−2.92	−2.94	−2.96	−2.97	−3.00
	4.75	−2.90	−2.93	−2.95	−2.97	−2.98	−3.01

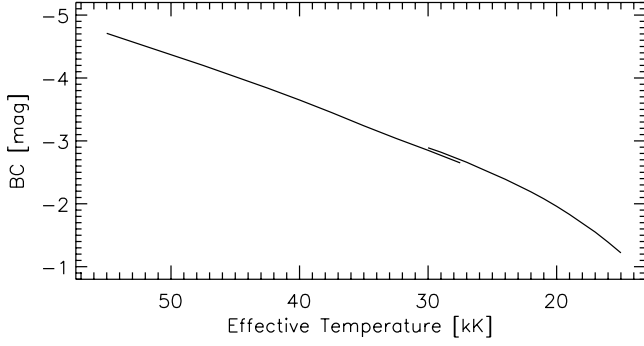


FIG. 18.—Bolometric corrections vs. effective temperature for solar composition, main-sequence OB stars.

gravity. Models tend to become numerically unstable when $\Gamma_{\text{rad}} > 0.9$, setting thus the gravity limit in our grid.

We can extrapolate these radiative acceleration to define the limit of photospheric stability when Γ_{rad} reaches unity, thus providing with an estimate of the Eddington limit in absence of rotation (see Fig. 19). Table 5 gives $\log g_{\text{Edd}}$, the Eddington limit in term of $\log g$, as a function of effective temperature from 55,000 to 15,000 K. We have carried out the same extrapolation with metal-free model atmospheres, and found that the corresponding Eddington limit is reached about 0.05 dex lower than the values derived for solar composition models. Lamers & Fitzpatrick (1988) performed a similar exercise, defining the Eddington limit by extrapolating the radiative accelerations calculated with Kurucz model atmospheres. Their estimates agree very well with our results over the whole range of temperature.

Finally, we can derive effective gravities, $g_{\text{eff}} = g - \max(g_{\text{rad}})$, in absence of rotation (see Table 5). These effective gravities may be used to estimate the characteristic pressure scale height ($h = P_{\text{gas}}/g_{\text{eff}}$) in the photosphere, after possibly correcting g_{eff} further for centrifugal acceleration. These scale heights indicate that the

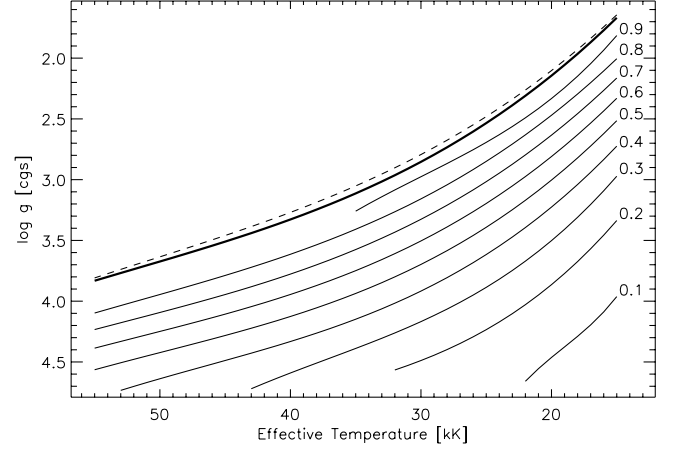


FIG. 19.—Isocurves of radiative acceleration relative to gravitational acceleration, $\Gamma_{\text{rad}} = \max(g_{\text{rad}})/g$, for solar composition model atmospheres, as a function of T_{eff} and $\log g$. Γ_{rad} values are listed right to the isolines. The thick line shows an estimate of the Eddington limit obtained by extrapolation; the dashed line corresponds to the estimate of the Eddington limit for metal-free model atmospheres.

photosphere of main-sequence B stars are compact, with scale heights typically smaller than 1% of the stellar radius. On the other hand, the photosphere of extreme B supergiants ($\log g \leq 2.0$) are quite extended ($h > 1 R_{\odot}$, or several percent of their radius). Caution should be exercised when applying these model atmospheres to analyze the spectra of these supergiants.

6. CONCLUSION

We have constructed a comprehensive grid of metal line-blanketed, NLTE, plane-parallel, hydrostatic model atmospheres for the basic parameters appropriate to early B-type stars. The BSTAR2006 grid considers 16 values of effective temperatures $15,000 \text{ K} \leq T_{\text{eff}} \leq 30,000 \text{ K}$, with 1000 K steps, 13 surface gravities, $1.75 \leq \log g \leq 4.75$ with 0.25 dex steps, and six chemical compositions, from metal-rich relative to the Sun to metal-free. The lower limit of $\log g$ for a given T_{eff} is actually set by an approximate location of the Eddington limit.⁹

⁹ The complete BSTAR2006 grid is available at our Web site at <http://nova.astro.umd.edu>.

TABLE 4

IONIZING FLUXES IN H I LYMAN CONTINUUM AS FUNCTION OF T_{eff} AND $\log g$ FOR SOLAR COMPOSITION MODEL ATMOSPHERES

T_{eff} (K)	$q_0 = \log N_{\text{LyC}} (\text{s}^{-1} \text{cm}^{-2})$		
	$\log g = 4$	$\log g = 3$	$\log g = 2$
55,000.....	25.01
50,000.....	24.82
45,000.....	24.58
40,000.....	24.28
35,000.....	23.84
30,000.....	22.90	23.62	...
29,000.....	22.63	23.45	...
28,000.....	22.35	23.25	...
27,000.....	22.06	22.99	...
26,000.....	21.76	22.66	...
25,000.....	21.44	22.25	...
24,000.....	21.12	21.82	...
23,000.....	20.81	21.40	...
22,000.....	20.51	20.95	...
21,000.....	20.21	20.54	...
20,000.....	19.90	20.16	...
19,000.....	19.59	19.82	...
18,000.....	19.26	19.47	20.74
17,000.....	18.92	19.11	19.63
16,000.....	18.57	18.74	19.11
15,000.....	18.20	18.35	18.65

TABLE 5

EDDINGTON LIMIT AND EFFECTIVE GRAVITIES AS FUNCTION OF EFFECTIVE TEMPERATURE FOR SOLAR COMPOSITION MODEL ATMOSPHERES

T_{eff} (K)	$\log g_{\text{Edd}}$	$\log g_{\text{eff}}$		
		$\log g = 4$	$\log g = 3$	$\log g = 2$
55,000.....	3.83	3.10
50,000.....	3.67	3.37
45,000.....	3.51	3.51
40,000.....	3.33	3.63
35,000.....	3.11	3.73
30,000.....	2.85	3.81	2.02	...
28,000.....	2.73	3.83	2.22	...
26,000.....	2.60	3.86	2.35	...
24,000.....	2.46	3.88	2.48	...
22,000.....	2.31	3.90	2.59	...
20,000.....	2.14	3.92	2.69	...
18,000.....	1.96	3.93	2.76	0.57
15,000.....	1.67	3.96	2.85	1.32

NOTE.—Gravities are in cgs units.

We have intended to provide a more or less definitive grid of model atmospheres in the context of one-dimensional, plane-parallel, homogeneous, hydrostatic models in radiative equilibrium. We have attempted to take into account essentially all important opacity sources (lines and continua) of all astrophysically important ions. Likewise, we have attempted to consider all relevant atomic processes that determine the excitation and ionization balance of all such atoms and ions. The models explicitly include, and allow for departures from LTE, 46 ions of H, He, C, N, O, Ne, Mg, Al, Si, S, and Fe, and about 53,000 individual atomic levels (including about 49 300 iron levels) grouped into 1127 superlevels. Line opacity includes about 39,000 lines from the light elements and 500,000 to 2 million iron lines dynamically selected from a list of about 5.65 million lines.

Although we spent all effort to make sure that the treatment of atomic physics and opacities is as complete and accurate as possible, there are still several points we are aware of that were crudely approximated. Those approximations were necessitated not by shortcomings of our modeling scheme or a lack of adequate computer resources, but by the present lack of sufficient atomic data. These approximations include a crude and approximate treatment of collision rates, a neglect of very high energy

levels of light metals, and ignoring charge exchange reactions. Despite our updating most model atoms to include levels up to $n = 10$, we caution that the analysis of weak lines from highly excited energy levels in the optical and infrared spectrum might require to construct even more detailed model atoms, supplementing the model atoms with higher excitation levels, or possibly also splitting some superlevels into individual levels.

Despite the remaining approximations and uncertainties, we believe that the BSTAR2006 grid represents a definite improvement over previous grids of B star model atmospheres. Combined with our earlier OSTAR2002 grid of O star model atmospheres, we hope that these new model atmospheres will thus be useful for a number of years to come for analyzing the spectrum individual O and B stars, as well as constructing composite model spectra of clusters of young massive stars, OB associations, and starburst galaxies.

This work was supported by NASA grant NAG5-13214 (FUSE D162 program), NASA ADP grants, and by several grants (GO 9116, GO 9848) from the Space Telescope Science Institute, which is operated by the Association of Universities for Research in Astronomy, Inc., under NASA contract NAS5-26555.

REFERENCES

- Asplund, M., Grevesse, N., & Sauval, A. J. 2005, in ASP Conf. Ser. 336, Cosmic Abundances as Records of Stellar Evolution and Nucleosynthesis, ed. T. G. Barnes & F. N. Bash (San Francisco: ASP), 25
- Bautista, M. A. 1996, A&AS, 119, 105
- Bautista, M. A., & Pradhan, A. K. 1997, A&AS, 126, 365
- Bautista, M. A., Romano, P., & Pradhan, A. K. 1998, ApJS, 118, 259
- Bouret, J.-C., Lanz, T., & Hillier, D. J. 2005, A&A, 438, 301
- Bouret, J.-C., Lanz, T., Hillier, D. J., Heap, S. R., Hubeny, I., Lennon, D. J., Smith, L. J., & Evans, C. J. 2003, ApJ, 595, 1182
- Butler, K., Mendoza, C., & Zeippen, C. J. 1993, J. Phys. B, 26, 4409
- Crowther, P. A., Lennon, D. J., & Walborn, N. R. 2006, A&A, 446, 279
- Cunha, K., Hubeny, I., & Lanz, T. 2006, ApJ, 647, L143
- Cunha, K., & Lambert, D. L. 1994, ApJ, 426, 170
- Dufton, P. L., Ryans, R. S. I., Trundle, C., Lennon, D. J., Hubeny, I., Lanz, T., & Allende Prieto, C. 2005, A&A, 434, 1125
- Fernley, J. A., Hibbert, A., Kingston, A. E., & Seaton, M. J. 1999, J. Phys. B, 32, 5507
- Giovanardi, C., Natta, A., & Palla, F. 1987, A&AS, 70, 269
- Grevesse, N., & Sauval, A. 1998, Space Sci. Rev., 85, 161
- Heap, S. R., Lanz, T., & Hubeny, I. 2006, ApJ, 638, 409
- Herrero, A., Puls, J., & Najarro, F. 2002, A&A, 396, 949
- Hibbert, A., & Scott, M. P. 1994, J. Phys. B, 27, 1315
- Hillier, D. J., & Miller, D. L. 1998, ApJ, 496, 407
- Hubeny, I. 1988, Comput. Phys. Commun., 52, 103
- Hubeny, I., & Lanz, T. 1995, ApJ, 439, 875
- Kurucz, R. L. 1993, Kurucz CD-ROM 13, ATLAS 9 Stellar Atmosphere Programs and 2 km/s Grid (Cambridge: SAO)
- Kurucz, R. L. 1994, Kurucz CD-ROM 22, Atomic Data for Fe and Ni (Cambridge: SAO)
- Lamers, H. J. G. L. M., & Fitzpatrick, E. L. 1988, ApJ, 324, 279
- Lanz, T., Artru, M.-C., Le Dourneuf, M., & Hubeny, I. 1996, A&A, 309, 218
- Lanz, T., & Hubeny, I. 2003, ApJS, 146, 417 (Paper I)
- Luo, D., & Pradhan, A. K. 1989, J. Phys. B, 22, 3377
- Martins, F., Schaerer, D., Hillier, D. J., Meynadier, F., Heydari-Malayeri, M., & Walborn, N. R. 2005, A&A, 441, 735
- Mendoza, C., Eissner, W., Le Dourneuf, M., & Zeippen, C. J. 1995, J. Phys. B, 28, 3485
- Mihalas, D., & Auer, L. H. 1970, ApJ, 160, 1161
- Mihalas, D., Heasley, J. N., & Auer, L. H. 1975, NCAR-TN/STR-104 (Boulder: NCAR)
- Nahar, S. N. 1996, Phys. Rev. A, 53, 1545
- . 1997, Phys. Rev. A, 55, 1980
- Nahar, S. N., & Pradhan, A. K. 1993, J. Phys. B, 26, 1109
- Najarro, F., Hillier, D. J., Puls, J., Lanz, T., & Martins, F. 2006, A&A, 456, 659
- The Opacity Project Team. 1995, The Opacity Project, Vol. 1 (Bristol: Inst. Phys.)
- The Opacity Project Team. 1997, The Opacity Project, Vol. 2 (Bristol: Inst. Phys.)
- Peach, G., Saraph, H. E., & Seaton, M. J. 1988, J. Phys. B, 21, 3669
- Santolaya-Rey, A. E., Puls, J., & Herrero, A. 1997, A&A, 323, 488
- Schaller, G., Schaerer, D., Meynet, G., & Maeder, A. 1992, A&AS, 96, 269
- Schmidt-Kaler, T. 1982, in Landolt-Börnstein, New Series, Group VI, Vol. 2b, ed. K. Schaifers & H. H. Voigt (Berlin: Springer), 451
- Sigut, T. A. A. 1999, ApJ, 519, 303
- Tully, J. A., Seaton, M. J., & Berrington, K. A. 1990, J. Phys. B, 23, 3811