

A CANDIDATE FOR THE MOST LUMINOUS OB ASSOCIATION IN THE GALAXY

MUBDI RAHMAN, CHRISTOPHER MATZNER, AND DAE-SIK MOON

Department of Astronomy & Astrophysics, University of Toronto, 50 St. George Street, Toronto, Ontario M5S 3H4, Canada; rahman@astro.utoronto.ca
 Received 2010 December 2; accepted 2011 January 18; published 2011 January 31

ABSTRACT

The Milky Way harbors giant H II regions, which may be powered by star complexes more luminous than any known Galactic OB association. Being across the disk of the Galaxy, however, these brightest associations are severely extinguished and confused. We present a search for one such association toward the most luminous H II region in the recent catalog by Murray & Rahman, which, at ~ 9.7 kpc, has a recombination rate of $\sim 7 \times 10^{51} \text{ s}^{-1}$. Prior searches have identified only small-scale clustering around the rim of this shell-like region, but the primary association has not previously been identified. We apply a near-infrared color selection and find an overdensity of point sources toward its southern central part. The colors and magnitudes of these excess sources are consistent with O- and early B-type stars at extinctions $0.96 < A_K < 1.2$, and they are sufficiently numerous (406 ± 102 after subtraction of field sources) to ionize the surrounding H II region, making this a candidate for the most luminous OB association in the Galaxy. We reject an alternate theory, in which the apparent excess is caused by localized extinction, as inconsistent with source demographics.

Key words: infrared: stars – open clusters and associations: individual (Dragonfish Association) – stars: formation – stars: massive

Online-only material: color figures

1. INTRODUCTION

Very massive OB associations are objects of intense interest. Given that the association birthrate (\dot{N}_{cl}) is observed to vary with association mass (M_{cl}) as $d\dot{N}_{\text{cl}}/d \ln M_{\text{cl}} \propto 1/M_{\text{cl}}$ (McKee & Williams 1997), or possibly slightly flatter, a large portion of each galaxy’s star formation occurs within its most massive OB associations. Giant OB associations are most capable of disrupting their gaseous environment, from their natal molecular clouds to the entire galactic neighborhood, and they inflate superbubbles which erupt from the disk, feeding a galactic halo or fountain. They sculpt and illuminate the giant H II regions seen in distant galaxies and are a notable feature of starburst evolution. Being the most extreme examples of intense star formation in the current universe, they are laboratories for the physics of star cluster formation such as environmental influences on the initial mass function (IMF), formation of very massive stars, and dynamical evolution with rapid stellar evolution.

Although several of the brightest known OB associations within the Milky Way approach “super star cluster” status ($M > 10^4 M_{\odot}$; Portegies Zwart et al. 2010), it is possible that even more luminous associations have escaped detection because of severe extinction and confusion in the Galactic plane. For instance, the Galactic OB associations with the largest known ionizing output, S , are those powering the NGC 3603 H II, Arches, Quintuplet, and Galactic center regions ($\log S/\text{s}^{-1} \simeq 51.5, 51.0, 50.9$, and 50.5 , respectively; Figer 2008; Conti & Crowther 2004); however, the upper limit of the Milky Way association distribution is estimated to be even brighter ($\log S_{\text{u}}/\text{s}^{-1} \simeq 51.7$; McKee & Williams 1997).

Recently, Rahman & Murray (2010) identified 40 star-forming complexes within the 13 most luminous ($S > 3 \times 10^{51} \text{ s}^{-1}$) free-free emission regions observed by the *Wilkinson Microwave Anisotropy Probe* (WMAP) and cataloged by Murray & Rahman (2010). Many of these complexes are potential hosts of embedded super star clusters. Of particular interest is the most luminous unconfused star-forming complex located at $(l, b) =$

$(298.4, -0.4)$ with a kinematic distance of 9.7 kpc. With an output of $\log S/\text{s}^{-1} \simeq 51.8$, the stars powering this region may represent the Milky Way’s most luminous OB association. (We refer to this region and its cluster as G298 or, inspired by its $8 \mu\text{m}$ appearance, as the “Dragonfish Nebula”; Figure 1.) The region appears in $8 \mu\text{m}$ as a closed bubble surrounded by a prominent shell. Previous searches using near-infrared (NIR) point-source catalogs from the Two Micron All Sky Survey (2MASS; Dutra et al. 2003) and *Spitzer* GLIMPSE (Mercer et al. 2005) have concentrated on the brightest regions of free-free emission around the shell and relied on integrated source counts. Although some clustering was discovered in these studies (Figure 2), the central ionizing association is yet undiscovered.

In this Letter, we identify a candidate for the central OB association powering the Dragonfish, primarily by applying a color selection to the 2MASS catalog. Our candidate is potentially the most luminous (and thus massive) OB association in the present Milky Way.

2. FEASIBILITY AND METHODS

Only very luminous stars of deeply embedded OB associations are visible due to distance and extinction. Whether the association can be identified using the 2MASS point-source catalog (Skrutskie et al. 2006) depends on the OB association radius and field star density.

Radii in the range 10–30 pc are likely since most associations, irrespective of mass, are disrupted immediately after birth (Roberts 1957). OB association birthsites in the Milky Way have a limited range of mass column densities, $\Sigma \simeq 0.3 \text{ g cm}^{-2} \pm 0.5 \text{ dex}$ (Fall et al. 2010). Turbulent line widths in these birthsites are $\sigma = 5\alpha^{1/2}[(0.5/\varepsilon)(\Sigma/.3 \text{ g cm}^{-2})M_5]^{1/4} \text{ km s}^{-1}$, for associations of mass $10^5 M_5 M_{\odot}$ which form with efficiency ε from a region with turbulent virial parameter α . Crossing and formation timescales are both shorter than the ionizing lifetime and the terminal expansion speed of an OB association is likely to be $\sim \sigma$. For an association of $\sim 10^5 M_{\odot}$, an age between 2

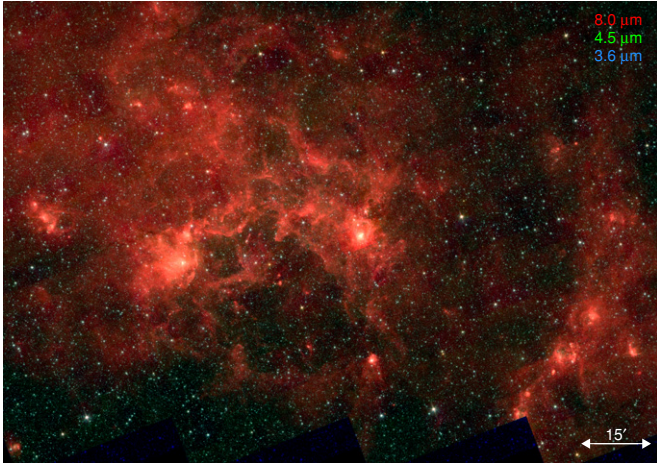


Figure 1. G298 or the Dragonfish region seen in the *Spitzer* GLIMPSE 3.6 (blue), 4.5 (green), and 8.0 (red) μm bands (Benjamin et al. 2003). (A color version of this figure is available in the online journal.)

and 4 Myr, and an expansion speed between 4 and 6 km s^{-1} , we therefore expect an expanding population currently 8–24 pc in radius, possibly surrounding a dense core which survived the disruption event. For a physical radius $R = 10(R/10 \text{ pc})$ pc with the ionizing luminosity of the Dragonfish as measured from *WMAP*, the association source density is $10(R/10 \text{ pc})^{-2} \text{ arcmin}^{-2}$. We measure the 2MASS source population density in the Galactic plane to be 14–17 arcmin^{-2} . For $R < 8$ pc the association dominates raw 2MASS source counts, whereas for $R > 10$ pc it is a minor variation of the 2MASS population which is difficult to detect using raw source counts.

The NIR color information provides an additional avenue to search for the embedded associations. We make use of this in the search for the embedded association within the Dragonfish. Ionization from a coeval stellar population drops sharply after 4 Myr as its most massive and powerful stars turn off the main sequence, so it is unlikely, though not impossible, that G298’s ionizing population is any older. The ionizing population is sufficiently rich ($S > 10^{50} \text{ s}^{-1}$; Kennicutt et al. 1989) to fully sample the stellar IMF; its very massive stars should still be present and visible.

The association is likely highly extinguished, even if it has cleared its natal cloud. At a distance of 9.7 kpc the fiducial distance–extinction relationship within the Galactic disk, $A_V/D \simeq 1.6 \text{ kpc}^{-1}$ (Binney & Merrifield 1998, p. 137), predicts $A_V \simeq 16$; this corresponds to a K_s -band extinction $A_{K_s} \simeq 1.0$ (Nishiyama et al. 2008). We adopt the Nishiyama et al. (2009) extinction law, which is steeper than that of Cardelli et al. (1989). We rely primarily on the 2MASS point-source catalog for our analysis, but use three additional data sets to check for consistency: the USNO-B catalog of visible magnitudes, the *Spitzer* GLIMPSE catalog at 3.6 and 4.5 μm , and deeper J - and H -band photometry of a small section of the candidate.

For our search, we target main-sequence, massive stars at a distance of ~ 9.7 kpc and an extinction of around $A_K \simeq 1$. Further, association members visible in all three bands have spectral types of roughly B1 or earlier, so their intrinsic emission in the NIR bands ($\lambda \sim 1.0\text{--}2.4 \mu\text{m}$) is entirely in the Rayleigh–Jeans limit ($(J - H, H - K) = (-0.11, -0.10)$; Martins & Plez 2006). With the predicted extinction, NIR colors should be $(J - H, H - K) \simeq (1.1, 0.5)$. Association sources should lie in a tight group along the reddening vector

from this common initial color. If the upper IMF follows the Salpeter (1955) slope to an upper limit of $120 M_\odot$, we expect $\sim 400 \times 10^{-0.34(A_K - 1.0)}(D/9.7 \text{ kpc})^{1.16}$ members with confident NIR colors ($[J, H, K_s] < [15.9, 15.0, 14.3]$, for which the signal-to-noise ratio (S/N) exceeds five in all bands). Note that if the association is older than 4 Myr or is not coeval, it must be even more numerous because more stars are required to make up for the missing ionization from the top of the main sequence.

The first steps of our analysis are to select sources with the color of an early-type star behind a range of extinctions bracketing the expected value, to test for a statistically significant overdensity within the confines of the 8 μm bubble, and then to vary the adopted extinction range in order to optimize the statistical significance of the overdensity. In practice we determine source density using the angular offset θ_N between a source and its N th neighbor, as $\Sigma_* = N/(\pi\theta_N^2)$, and use $N \sim 20\text{--}100$ to reduce Poisson noise at some expense in resolution. Once the angular scale of the putative association is determined, we assess its significance by comparing source density within this region to the mean $\bar{\Sigma}_*$ and standard deviation σ_Σ of those within independent, identically sized regions in a two-degree field centered on, but excluding, the 8 μm bubble. Statistical significance is measured by $(\Sigma_* - \bar{\Sigma}_*)/\sigma_\Sigma$.

Once a candidate is identified, we check that the same procedure applied to regions outside the 8 μm bubble yields no candidates of similar significance. We then examine the color–color and color–magnitude diagrams (CCDs and CMDs, respectively) of similarly sized regions to check that an apparent overdensity is not caused by features of the distribution of stars and extinction.

3. RESULTS

3.1. Candidate OB Association

Based on the expected source colors from Section 2, we apply an NIR color cut of $1.0 < J - H < 1.4$ and $0.44 < H - K_s < 0.62$, corresponding to colors of hot stars with $0.9 < A_K < 1.2$. Within the color cut, the two-degree diameter field encompassing the Dragonfish has a mean stellar density $\bar{\Sigma}_* = 1.4 \text{ arcmin}^{-2}$, which we refer to as the field star density. The standard deviation of Σ_* averaged over association-sized regions is $\sigma_\Sigma = 0.30 \text{ arcmin}^{-2}$. We find a significant overdensity of stars in the Dragonfish Nebula at $(l, b) = (298^\circ 55', -0^\circ 72')$, located inside the shell of the star-forming complex (Figure 2). The 1σ stellar overdensity contour ($\Sigma_* = \bar{\Sigma}_* + \sigma_\Sigma$) has semimajor and semiminor axes ($11', 10'$), corresponding to (31, 28) pc; however, the association may extend beyond the observational limit. The peak of the overdensity is $\Sigma_* = 3.7 \text{ arcmin}^{-2}$, 8σ above the background level on cluster-sized regions, adopting the 100th nearest neighbor. The region contains 897 sources within the adopted boundaries. Subtracting the contribution due to background sources, the candidate association is composed of 406 ± 102 sources, consistent with the value determined based on the chosen IMF and measured luminosity (Section 2). Using the 1σ boundaries, the mean density of the association is $\Sigma_* + 4.0\sigma_\Sigma$. We refer to the candidate as the “Dragonfish Association.”

The association is strongly concentrated within the central $4.1'$, but contains an extended asymmetric envelope. The central, dense, symmetrical region has a stellar density which depends as $\theta^{-\alpha}$ with $\alpha = 0.44 \pm 0.03$ for $2' < \theta < 3'$, where the value of α depends on the assumed density of field sources. At larger radii the azimuthally averaged density drops more rapidly, but is dominated by the asymmetries visible in Figure 1. Half

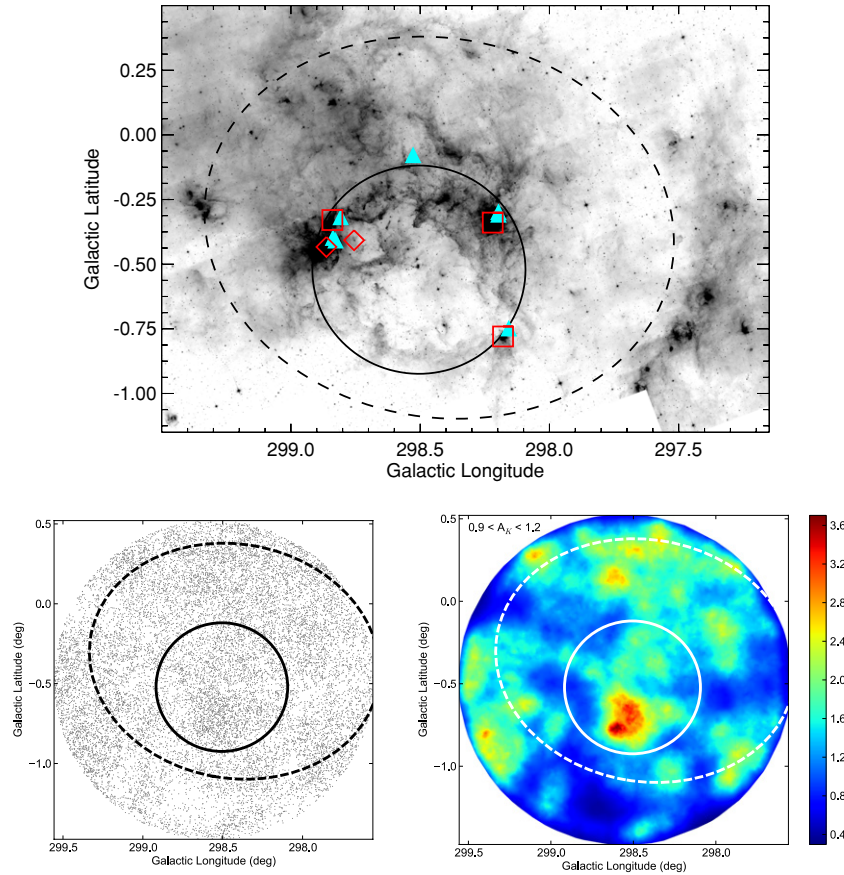


Figure 2. Top: the $8\,\mu\text{m}$ GLIMPSE image of the Dragonfish Nebula. The dotted ellipse indicates the location of the *WMAP* source, and the solid ellipse indicates the location of the star-forming complex. The cyan triangles indicate the location of known H II regions. Clockwise from top: G298.56-0.11, G298.2-0.3, G298.19-0.78, G298.9-0.4, and G298.837-0.347. The red squares and diamonds indicate the regions of stellar clustering identified by Dutra et al. (2003) and Mercer et al. (2005). Bottom left: the individual 2MASS sources within the field surrounding the Dragonfish Nebula passing the color cut. Bottom right: the 2MASS stellar densities of the field surrounding the Dragonfish Nebula at the color cut corresponding to O stars with extinctions between $0.9 < A_K < 1.2$. The color bar indicates the stellar density in stars per square arcminute.

(A color version of this figure is available in the online journal.)

the excess sources are found within $4''.7$, or $13(D/9.7\text{ kpc})$ pc, similar to the half-light radii of other Galactic OB associations (Portegies Zwart et al. 2010).

Assuming half the projected mass falls within this radius, and assigning a total mass of $10^5 M_\odot$ we infer that the association's central regions are older than their virial crossing times, whereas its outskirts are not ($T_{cr} = 22$ Myr). It therefore seems likely to be unbound on the whole (Gieles & Portegies Zwart 2011) but to contain a bound core; however this bears further investigation.

Using the USNO and *Spitzer* GLIMPSE catalogs, we are able to confirm that the stars identified in the Dragonfish Association have magnitudes and colors consistent with OB stars with the stated extinction, but not further able to distinguish the association population from the background. In the GLIMPSE wavelengths ($>3\,\mu\text{m}$), this is because nearly all stars have identical, Rayleigh-Jeans colors, whereas with the USNO wavelengths ($<1\,\mu\text{m}$), the spectral energy distribution is strongly dominated by the extinction. Thus, template colors of most stellar types can be fit to the additional photometry by varying the assumed distance or extinction of the star.

3.2. Color-Color and Color-Magnitude Diagrams

The NIR color information provides an additional avenue to determine important parameters of the candidate association,

including the statistical significance of the overdensity, the possible background contributors, and the color variation of all point sources within the association boundaries. Figure 3 presents the CCD and CMD of the point sources within the two-degree field centered on the Dragonfish Nebula, showing the distribution of all 2MASS point sources, as well as the relative location of the color cut. We indicate in red the point sources that fall within the association boundaries. We indicate in blue all stars within the field that meet the association's color cut. The position of the main sequence is indicated at a distance of 9.7 kpc and extinguished by $A_K = 1.0$. The reference magnitudes for the O spectral type are taken from Martins & Plez (2006) and for the remaining spectral types from Pickles (1998). The path taken by an O5V star if placed at varying distances, assuming the average extinction to distance ratio, is indicated by the red line. The path taken by any given reference star is just a translation of this path on the diagram. The color cut is sufficiently large to permit spectral classes of A and earlier at the given reddening. However, the sensitivity of 2MASS limits stars at this distance and extinction to be no later than B1V.

In Figure 4, we bin the CCD and CMD to investigate the overabundance in the region within the Dragonfish Association as compared to the surrounding regions. These figures indicate the location and prominence of the cluster overdensity as a function of both color and magnitude. For the CCD, we bin the

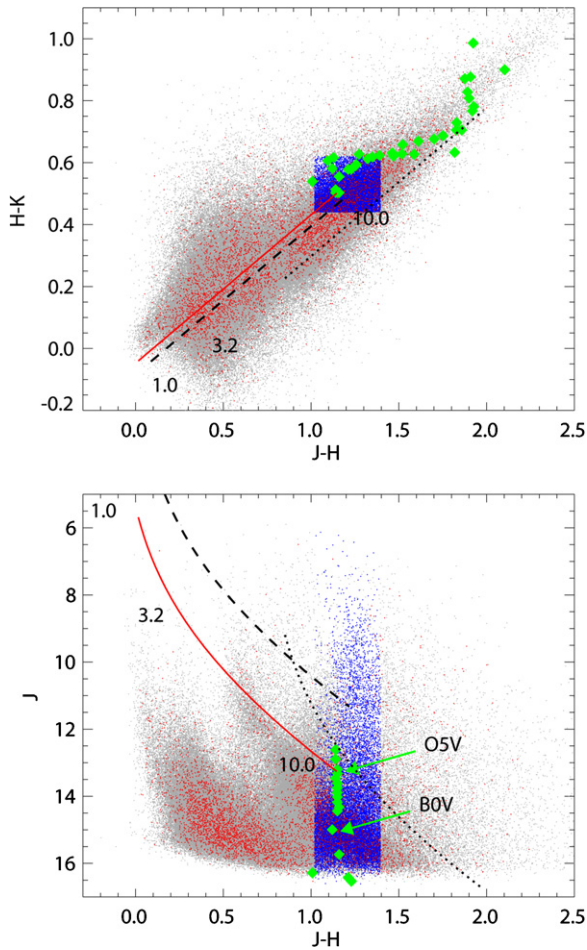


Figure 3. Color-color (top) and color-magnitude (bottom) diagrams of the 2MASS point sources. The gray points are all points within the two-degree field, while the red points are the sources within the association boundaries. The blue points indicate all sources within the color cut. The green diamonds indicate the main sequence at a distance of 9.7 kpc with an extinction of $A_K = 1.0$. The locations of an O5V and an B0V star are indicated for reference. The red line indicates the track of an O5V star at various distances, with the numbers indicating the distance in kpc. Similar paths are shown for K3II (black dotted line) and B0I (black dashed line) stars.

(A color version of this figure is available in the online journal.)

sources in a 35^2 bin grid spanning $-0.3 < J - H < 2.5$ and $-0.2 < H - K < 1.1$. For the CMD, we similarly use a 35^2 bin grid with $5 < J - H < 17.0$, and identical $J - H$ limits. We grid the point sources that fall within the association boundaries and subtract the average bin value from the surrounding field. In both the CCD and CMD, an overabundance of point sources appears in the range $1.0 < J - H < 1.5$. We note an underdensity in the CCD at $(J - H, H - K) = (0.8, 0.3)$. The total number of sources “missing” in the underdensity is approximately 200. The underdensity is associated with a galactic structure feature in the surrounding field which we discuss in depth in a forthcoming paper.

Given that the colors of these underdense sources are separated from those of the putative OB association by a vector that is nearly parallel to the extinction vector, could the association be only an illusion caused by intervening extinction with $A_K \sim 0.25$? If so we would expect an overall deficit of stars in this direction (due to the extinction of dim sources), but there is no such trend. In fact the missing sources are about half as numerous as the surplus ones. Corroborating this point, we note

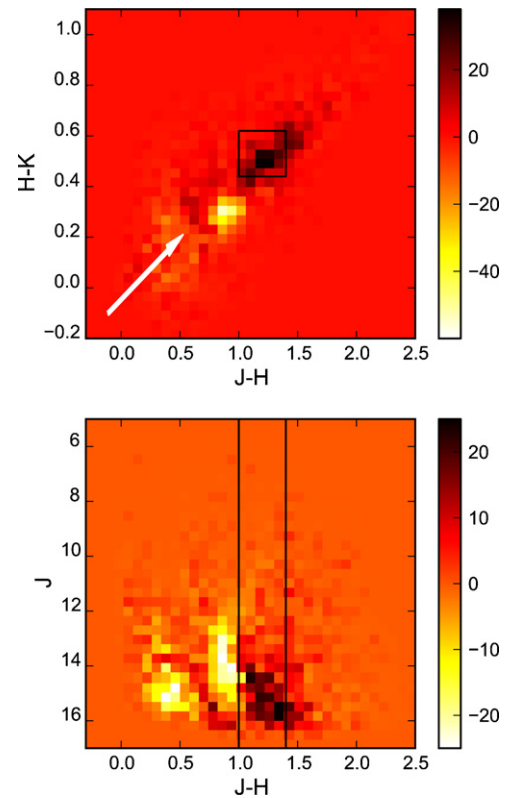


Figure 4. Binned color-color (top) and color-magnitude (bottom) diagrams of the 2MASS point sources within the candidate association boundaries with the binned background removed from the diagrams. The cluster color cut is indicated with the black lines. The color bars indicate the number of stars in each bin. The white arrow on the color-color diagram indicates the extinction vector corresponding to an $A_K = 0.5$ from the unreddened location of O stars.

(A color version of this figure is available in the online journal.)

that the vector separating the two populations is slightly steeper than the Nishiyama et al. (2008) extinction law (whereas the overpopulation is perfectly consistent with OB stars reddened by that law). Moreover, sources which appear underdense on this plot vary on larger angular scales than that of the association; their population is mostly a function of galactic latitude, although there are somewhat fewer of them toward the association. We therefore conclude that the association does in fact exist, but there may also exist localized extinction which affects the color distribution of sources.

To further investigate the role of patchy extinction, we have obtained additional J - and H -band photometry toward a central $3.9'$ field in the candidate association using the Wide-Field IR Camera on the 2.5 m du Pont Telescope, with limiting magnitudes of $J = 17.1$ and $H = 16.3$ at an S/N of 3 (J. Radigan 2010, private communication). We extract 906 stars in this field. With this additional photometry, we find stars in the range of $0.8 < J - H < 0.96$, with a continuous density of stars throughout this color range. We also find no evidence for any shadowing features which would indicate a strong extinction feature. This supports our conclusion that the apparent association is not an extinction feature.

4. CONFUSING SOURCES

In order to identify possible contaminating stars within the color cut, we model the NIR colors and magnitudes of template stars of different spectral types and luminosity classes. We vary the distances of the model stars to examine which stars meet the

cut above the 2MASS limiting magnitudes. We use reference magnitudes, extinction relationship, and the average extinction-to-distance ratio as in Section 3.2. All main-sequence OB stars down to B1V can fall within the color and magnitude limits. With the 2MASS limiting magnitudes, assuming the extinction-to-distance ratio, all subtypes of O stars are visible to 13 kpc. All main-sequence stellar types later than B1V fall below the magnitude limits. OB giants (down to B9III) and K1III to M0III giants can fall within the color and magnitude limits. All supergiant stars can fall within the association color cut and magnitude limits. We show examples of stars that may fall within the color and magnitude limits in Figure 3. The red-clump giant feature, composed of early K giant stars (López-Corredoira et al. 2002; Indebetouw et al. 2005) partially overlaps the association color cut toward the faint end, but is not the primary contributor to the background population.

Using the Robin et al. (2004) model of the stellar population we produce a synthetic catalog of stars. We simulate the field in the direction and with the size of the Dragonfish Association, using the 2MASS limiting magnitudes from Section 2, and a diffuse extinction law of $1.6 A_V/\text{kpc}$. The model background population underestimates the total number of sources. However, we find that the majority of stars within the color cut are early K giant stars with masses between 1 and $3 M_\odot$ and distances between 3.0 and 4.5 kpc; we assume that these represent our contaminating sources.

5. CONCLUSIONS

We have used the 2MASS point-source catalog to search for the central OB association within the Dragonfish Nebula by analyzing the on-sky stellar density with NIR colors consistent with extinguished O stars. We find an overdensity of point sources within the bubble at an extinction range of $0.96 < A_K < 1.2$. Comparing the CCD and CMD of the candidate association with the surrounding field, we confirm the presence of the overdensity at the chosen color range. We reject an extinction explanation for the overdensity and infer that the contaminating star population predominantly consists of K giants 3–4.5 kpc away. Our candidate association contains 406 ± 102 members visible in 2MASS with magnitudes consistent with O and early-B stars, similar to the ionizing population expected from the region’s free-free luminosity. We infer from its assumed mass, radius, and maximum age that it is likely to be unbound on the whole but may contain a bound core; however further investigation is necessary to determine its

dynamical state. This is a candidate for the most luminous (and consequently the most massive) OB association in the Galaxy. If confirmed by upcoming spectroscopic observations, the Dragonfish Association is an outstanding candidate for multiwavelength examination, as it is a promising laboratory of massive star formation, stellar dynamics, and feedback processes within the Galactic environment.

We thank J. Radigan for obtaining NIR data for us. We thank N. Murray, P. G. Martin, and R. Breton for the many helpful discussions. This publication makes use of data products from the Two Micron All Sky Survey, which is a joint project of the University of Massachusetts and the Infrared Processing and Analysis Center/California Institute of Technology, funded by the National Aeronautics and Space Administration and the National Science Foundation.

REFERENCES

- Benjamin, R. A., et al. 2003, *PASP*, **115**, 953
 Binney, J., & Merrifield, M. (ed.) 1998, *Galactic Astronomy* (Princeton, NJ: Princeton Univ. Press)
 Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, **345**, 245
 Conti, P. S., & Crowther, P. A. 2004, *MNRAS*, **355**, 899
 Dutra, C. M., Bica, E., Soares, J., & Barbay, B. 2003, *A&A*, **400**, 533
 Fall, S. M., Krumholz, M. R., & Matzner, C. D. 2010, *ApJ*, **710**, L142
 Figer, D. F. 2008, in *IAU Symp. 250, Massive Stars as Cosmic Engines*, ed. F. Bresolin, P. A. Crowther, & J. Puls (Cambridge: Cambridge Univ. Press), 247
 Gieles, M., & Portegies Zwart, S. F. 2011, *MNRAS*, **410**, L6
 Indebetouw, R., et al. 2005, *ApJ*, **619**, 931
 Kennicutt, R. C., Jr., Edgar, B. K., & Hodge, P. W. 1989, *ApJ*, **337**, 761
 López-Corredoira, M., Cabrera-Lavers, A., Garzón, F., & Hammersley, P. L. 2002, *A&A*, **394**, 883
 Martins, F., & Plez, B. 2006, *A&A*, **457**, 637
 McKee, C. F., & Williams, J. P. 1997, *ApJ*, **476**, 144
 Mercer, E. P., et al. 2005, *ApJ*, **635**, 560
 Murray, N., & Rahman, M. 2010, *ApJ*, **709**, 424
 Nishiyama, S., Nagata, T., Tamura, M., Kandori, R., Hatano, H., Sato, S., & Sugitani, K. 2008, *ApJ*, **680**, 1174
 Nishiyama, S., Tamura, M., Hatano, H., Kato, D., Tanabé, T., Sugitani, K., & Nagata, T. 2009, *ApJ*, **696**, 1407
 Pickles, A. J. 1998, *PASP*, **110**, 863
 Portegies Zwart, S. F., McMillan, S. L. W., & Gieles, M. 2010, *ARA&A*, **48**, 431
 Rahman, M., & Murray, N. 2010, *ApJ*, **719**, 1104
 Roberts, M. S. 1957, *PASP*, **69**, 59
 Robin, A. C., Reylé, C., Derrière, S., & Picaud, S. 2004, *A&A*, **416**, 157
 Salpeter, E. E. 1955, *ApJ*, **121**, 161
 Skrutskie, M. F., et al. 2006, *AJ*, **131**, 1163