

OPEN ACCESS

Searching for axions and ALPs from string theory

To cite this article: Andreas Ringwald 2014 *J. Phys.: Conf. Ser.* **485** 012013

View the [article online](#) for updates and enhancements.

You may also like

- [A next-generation liquid xenon observatory for dark matter and neutrino physics](#)

J Aalbers, S S AbdusSalam, K Abe et al.

- [Physics Beyond the Standard Model with Future X-Ray Observatories: Projected Constraints on Very-light Axion-like Particles with Athena and AXIS](#)

Júlia Sisk-Reynés, Christopher S. Reynolds, Michael L. Parker et al.

- [Physics beyond colliders at CERN: beyond the Standard Model working group report](#)

J Beacham, C Burrage, D Curtin et al.



ECS
The
Electrochemical
Society
Advancing solid state &
electrochemical science & technology

DISCOVER
how sustainability
intersects with
electrochemistry & solid
state science research

Searching for axions and ALPs from string theory

Andreas Ringwald

Deutsches Elektronen-Synchrotron DESY, Notkestr. 85, D-22607 Hamburg, Germany

E-mail: andreas.ringwald@desy.de

Abstract. We review searches for closed string axions and axion-like particles (ALPs) in IIB string flux compactifications. For natural values of the background fluxes and TeV scale gravitino mass, the moduli stabilisation mechanism of the LARGE Volume Scenario predicts the existence of a QCD axion candidate with intermediate scale decay constant, $f_a \sim 10^{9 \div 12}$ GeV, associated with the small cycles wrapped by the branes hosting the visible sector, plus a nearly massless and nearly decoupled ALP associated with the LARGE cycle. In setups where the visible sector branes are wrapping more than the minimum number of two intersecting cycles, there are more ALPs which have approximately the same decay constant and coupling to the photon as the QCD axion candidate, but which are exponentially lighter. There are exciting phenomenological opportunities to search for these axions and ALPs in the near future. For $f_a \sim 10^{11 \div 12}$ GeV, the QCD axion can be the dominant part of dark matter and be detected in haloscopes exploiting microwave cavities. For $f_a \sim 10^{9 \div 10}$ GeV, the additional ALPs could explain astrophysical anomalies and be searched for in the upcoming generation of helioscopes and light-shining-through-a-wall experiments.

1. Introduction

The QCD axion arises [1, 2] in the course of the arguably most plausible solution of the strong CP puzzle [3], that is the non-observation of a θ -angle term in QCD. In this context, the axion field a is introduced as a dynamical θ -angle term, enjoying a shift symmetry, $a \rightarrow a + \text{constant}$, broken only by anomalous CP-violating couplings to gauge fields. Correspondingly, its most general low-energy effective Lagrangian below the weak scale has the form [4],

$$\begin{aligned} \mathcal{L} = & \frac{1}{2} \partial_\mu a \partial^\mu a - \frac{g_3^2}{32\pi^2} \left(\bar{\theta} + \frac{a}{f_a} \right) F_{3,\mu\nu}^b \tilde{F}_3^{b,\mu\nu} - \frac{e^2}{32\pi^2} C_a \gamma \frac{a}{f_a} F_{\mu\nu}^{\text{em}} \tilde{F}_{\text{em}}^{\mu\nu} \\ & + \sum_\Psi \left[\frac{1}{2} (\tilde{X}_{\psi_R} + \tilde{X}_{\psi_L}) \bar{\Psi} \gamma^\mu \gamma_5 \Psi + \frac{1}{2} (\tilde{X}_{\psi_R} - \tilde{X}_{\psi_L}) \bar{\Psi} \gamma^\mu \Psi \right] \frac{\partial_\mu a}{f_a}, \end{aligned} \quad (1)$$

where $F_{3,\mu\nu}^b$ is the b th component of the gluon field strength tensor, $\tilde{F}_3^{b,\mu\nu} = \frac{1}{2} \epsilon^{\mu\nu\rho\sigma} F_{3,\rho\sigma}^b$ its dual, g_3 is the strong coupling, $F_{\mu\nu}^{\text{em}}$ is the electromagnetic field strength, Ψ denotes standard model matter fields, and f_a is the axion decay constant. The θ -term in the QCD Lagrangian can then be eliminated by absorbing it into the axion field, $a = \bar{a} - \bar{\theta} f_a$. Moreover, the topological charge density $\propto \langle F_{3,\mu\nu}^b \tilde{F}_3^{b,\mu\nu} \rangle \neq 0$, induced by topological fluctuations of the gluon fields such as QCD instantons, provides a nontrivial potential for the axion field \bar{a} which is minimized at zero expectation value, $\langle \bar{a} \rangle = 0$, thus wiping out strong CP violation, and giving the fluctuating field around this minimum, the QCD axion, a mass

$$m_a = \frac{m_\pi f_\pi}{f_a} \frac{\sqrt{m_u m_d}}{m_u + m_d} \simeq 0.6 \text{ meV} \times \left(\frac{10^{10} \text{ GeV}}{f_a} \right), \quad (2)$$



in terms of the light (u, d) quark masses, the pion mass m_π and the pion decay constant f_π [1], For large axion decay constant f_a , we see that the axion is a very weakly interacting (cf. Eq. (1)) slim particle [5, 6, 7, 8]. In particular, its coupling to photons [9, 10, 11] and electrons,

$$\mathcal{L} \supset -\frac{\alpha}{2\pi f_a} \left(C_{a\gamma} - \frac{2}{3} \frac{m_u + 4m_d}{m_u + m_d} \right) \frac{a}{4} F_{\mu\nu}^{\text{em}} \tilde{F}_{\text{em}}^{\mu\nu} + \frac{C_{ae}}{2f_a} \bar{e} \gamma^\mu \gamma_5 e \partial_\mu a, \quad (3)$$

are very much suppressed, e.g.

$$g_{a\gamma} \equiv \frac{\alpha}{2\pi f_a} \underbrace{\left(C_{a\gamma} - \frac{2}{3} \frac{m_u + 4m_d}{m_u + m_d} \right)}_{C_{a\gamma}} \sim 10^{-13} \text{ GeV}^{-1} \left(\frac{10^{10} \text{ GeV}}{f_a} \right) C_{a\gamma}. \quad (4)$$

Laboratory experiments as well as astrophysics, in particular the non-observation of solar axions by the helioscope CAST and the non-observation of drastic energy losses in horizontal branch stars or white dwarfs, constrain the axion decay constant, divided by the appropriate dimensionless coupling constants, to a scale much above the weak scale (see also Fig. 1),

$$\frac{f_a}{C_{a\gamma}} > 10^7 \text{ GeV} \Leftrightarrow g_{a\gamma} < 10^{-10} \text{ GeV}^{-1}, \quad \frac{f_a}{C_{ae}} > 10^9 \text{ GeV}. \quad (5)$$

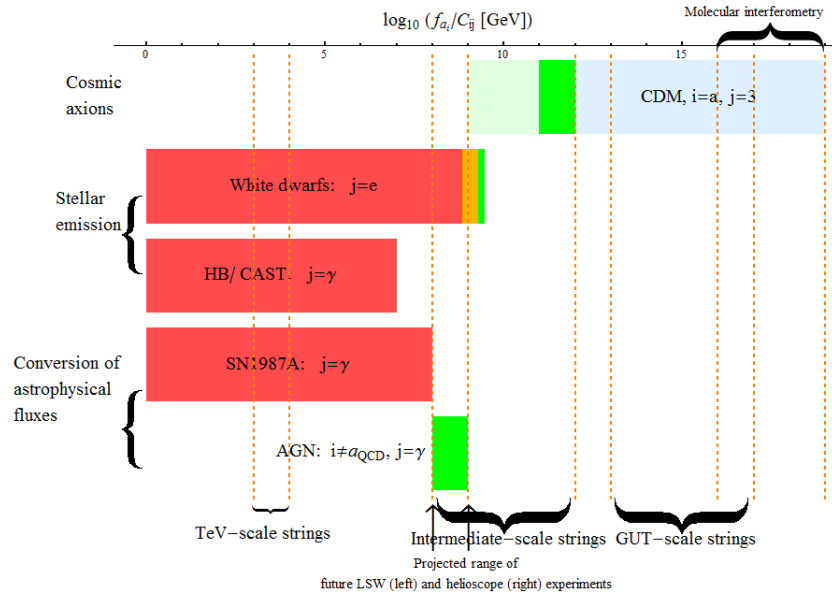


Figure 1. A summary of constraints on and hints for the couplings f_{a_i}/C_{ij} of axion-like particles a_i to standard model particles j [22]. Red regions are excluded, and the orange region would be excluded by red giants but is compatible with the hints from white dwarfs. The green regions from top to bottom correspond respectively to the classic ‘axion dark matter window’, hints of an axion from white dwarf cooling and transparency of the Universe to very high energy gamma rays. The blue region would be excluded by dark matter overproduction in the absence of a dilution mechanism or tuning of the initial misalignment angle.

Intriguingly, for an even higher decay constant, $f_a \gtrsim 10^{11} \text{ GeV}$, the QCD axion can contribute significantly to cold dark matter (CDM), being non-thermally produced in the early universe

via initial misalignment of the axion field, resulting in coherent field oscillations corresponding to a condensate of non-relativistic axions [12, 13, 14]. In fact, assuming that the reheating temperature after inflation is below f_a and that there is no dilution by, e.g., late decays of particles beyond the standard model, the expected cosmic mass fraction in QCD axion CDM is

$$\Omega_a h^2 \approx 0.71 \times \left(\frac{f_a}{10^{12} \text{ GeV}} \right)^{7/6} \left(\frac{\Theta_a}{\pi} \right)^2, \quad (6)$$

where Θ_a is the initial misalignment angle.

Therefore, the QCD axion is necessarily associated with a very high energy scale and so it is natural to search for it in ultra-violet completions of the standard model such as string theory. Indeed, it has long been known that the low-energy effective field theory of string compactifications predicts natural candidates for the QCD axion [15, 16, 17, 18, 19], often even an ‘axiverse’ [20], containing many additional light axion-like particles (ALPs) whose masses are logarithmically hierarchical. But only very recently explicit moduli-stabilised string theoretic examples with a viable QCD axion candidate and possibly additional light ALPs have been constructed (cf. Sec. 2), with decay constants ranging from the GUT scale, $f_a \sim 10^{16}$ GeV [21] down to the intermediate scale, $f_a \sim 10^{9 \div 12}$ GeV [22] – the latter offering exciting opportunities to detect effects of axions and ALPs in astrophysics and in the upcoming generation of axion experiments (cf. Sec. 3).

2. Axions and ALPs in IIB string flux compactifications

String theory requires the existence of six extra space dimensions, which appear to be unobservable because they are supposed to be compact and of very small size. String phenomenology is then the attempt to make contact between the perturbative ten dimensional (10D) effective field theories (EFTs) describing the massless degrees of freedom of string theory at very high energies – say, the heterotic or type II (with D -branes) EFT – and the low energy physics in our 4D real world. In fact, different 4D low energy EFTs emerge depending on the EFT to start with in 10D. One of the fundamental tasks of string phenomenology is to find a compactification whose low energy EFT reproduces (a suitable extension of) the standard model. Along this way it is also mandatory to understand moduli stabilisation: expectation values of the moduli fields, which parameterise the shape and size of the extra dimensions, determine many parameters of the low energy EFT, such as gauge and Yukawa couplings. Often, one needs to incorporate quantum corrections in order to fix many of these expectation values and to give their associated particle excitations a non-zero mass.

Moduli stabilisation is best understood in type IIB string flux compactifications, on which we concentrate in the following. They provide explicit and well motivated realisations of brane world scenarios: the standard model is supposed to live on a stack of space-time filling branes wrapping cycles in the compact dimensions, while gravity propagates in the bulk, leading to a string scale $M_s \sim M_P/\sqrt{\mathcal{V}}$ possibly much smaller than the Planck scale M_P , at the expense of a large compactification volume $\mathcal{V} \gg 1$ (in units of the string length). Importantly, axion-like fields emerge in those compactifications inevitably as Kaluza-Klein zero modes of ten dimensional form fields, as we will review in the following subsection.

2.1. Tree-level candidates for axions and ALPs

We will consider here type IIB flux compactifications on Calabi-Yau orientifolds X in the presence of space-time filling $D7$ branes and $O7$ planes. Below the Kaluza-Klein scale M_{KK} , they lead to a low-energy $\mathcal{N} = 1$ (supergravity) EFT in 4D. Their 4D closed string moduli, comprised by the axio-dilaton $S = e^{-\phi} + iC_0$, the complex structure moduli U_α , $\alpha = 1, \dots, h_{-1}^{2,1}(X)$, and the

Kähler moduli,

$$T_i = \tau_i + i c_i, \quad \tau_i = \text{Vol}(D_i), \quad c_i = \int_{D_i} C_4, \quad i = 1, \dots, h^{1,1}, \quad (7)$$

are obtained via the Kaluza-Klein reduction of the massless bosonic fields of the original 10D theory – the latter including, in the Ramond-Ramond sector, the forms C_0 and C_4 , and, in the Neveu Schwarz-Neveu Schwarz sector, the dilaton ϕ . Importantly, the number of Kähler moduli T_i is determined by the topology of X , namely the number of inequivalent four-cycles D_i of X (we have specialised for convenience to orientifold projections such that $h_-^{1,1} = 0 \Rightarrow h_+^{1,1} = h^{1,1}$). Their imaginary parts c_i have all the properties of axion-like fields, as we will see next.

As already mentioned, (a suitable extension of) the standard model is realised in such compactifications through (stacks of) space-time filling $D7$ -branes wrapping some of the four-cycles D_i . At low energies, the dynamics of a $D7$ -brane reduces to a $U(1)$ gauge theory that lives on its eight-dimensional world-volume. Moreover, $D7$ -branes can be magnetised by turning on internal magnetic fluxes \mathcal{F} . Taking all these effects into account, the 4D low energy effective Lagrangian of the axion-like fields c_i , including the brane-localised $U(1)$ gauge bosons A_i , is obtained from the Kaluza-Klein reduction of the $D7$ -brane action as follows [23, 24]:

$$\begin{aligned} \mathcal{L} \supset & - \left(dc_\alpha + \frac{M_P}{\pi} A_i q_{i\alpha} \right) \frac{\mathcal{K}_{\alpha\beta}}{8} \wedge \star \left(dc_\beta + \frac{M_P}{\pi} A_j q_{j\beta} \right) + \frac{1}{4\pi M_P} r^{i\alpha} c_\alpha \text{tr}(F \wedge F) \\ & + \frac{M_P^2}{2(2\pi)^2} A_i A_j q_{i\alpha} \mathcal{K}_{\alpha\beta} q_{j\beta} - \frac{r^{i\alpha} \tau_\alpha}{4\pi M_P} \text{tr}(F_i \wedge \star F_i), \end{aligned} \quad (8)$$

where $M_P = (8\pi G_N)^{-1/2} \simeq 2.4 \cdot 10^{18}$ GeV is the reduced Planck mass, while the various other quantities in Eq. (8) are defined as follows:

- The Kähler metric $K_{\alpha\beta}$, describing the kinetic mixing of the axion-like fields, is obtained via $\mathcal{K}_{\alpha\beta} \equiv \frac{\partial^2 K}{\partial \tau_\alpha \partial \tau_\beta}$ from the Kähler potential, the latter taking, at tree-level, the following form:

$$K_{\text{tree}} = -2 \ln \mathcal{V} - \ln(S + \bar{S}) - \ln \left(-i \int_X \Omega \wedge \bar{\Omega} \right). \quad (9)$$

It depends implicitly on the complex structure moduli via the holomorphic $(3,0)$ -form Ω and on the Kähler moduli via the Calabi-Yau volume \mathcal{V} , measured by an Einstein frame metric $g_{\mu\nu}^E = e^{-\phi/2} g_{\mu\nu}^s$, and expressed in units of the string length $\ell_s = 2\pi\sqrt{\alpha'}$, in terms of its tension α' ,

$$\mathcal{V} = \frac{1}{6} \int_X J \wedge J \wedge J = \frac{1}{6} k_{\alpha\beta\gamma} t^\alpha t^\beta t^\gamma. \quad (10)$$

Here, the Kähler form J has been expanded, in a basis $\{\hat{D}_\alpha\}_{\alpha=1}^{h^{1,1}}$ of $H^{1,1}(X, \mathbb{Z})$ of two-forms which are Poincaré dual to D_α , as $J = t^\alpha \hat{D}_\alpha$, and we denoted the triple intersection numbers of X by $k_{\alpha\beta\gamma}$. The volume, the Kähler form and ultimately the Kähler metric can then be obtained as a function of the τ_α by inverting the following relations:

$$\tau_\alpha = \frac{1}{2} \int_X \hat{D}_\alpha \wedge J \wedge J = \frac{\partial \mathcal{V}}{\partial t^\alpha} = \frac{1}{2} k_{\alpha\beta\gamma} t^\beta t^\gamma. \quad (11)$$

The string scale $M_s = 1/\ell_s$ is obtained from the dimensional reduction of the IIB supergravity action,

$$M_s = M_P / \sqrt{4\pi \mathcal{V}}. \quad (12)$$

- The couplings $r^{i\alpha}$ appearing in the gauge kinetic term and in the axionic couplings to gauge bosons, the respective last terms in both lines of Eq. (8), are the expansion coefficients of the two form $\hat{D}_i = r^{i\alpha}\omega_\alpha$,

$$r^{i\alpha} \equiv \ell_s^{-4} \int_{D_i} \tilde{\omega}^\alpha = \ell_s^{-4} \int \hat{D}_i \wedge \tilde{\omega}^\alpha, \quad \alpha = 1, \dots, h^{1,1}, \quad (13)$$

where the basic forms satisfy $\ell_s^{-4} \int \omega_\beta \wedge \tilde{\omega}^\alpha = \delta_\beta^\alpha$. The gauge coupling can be inferred from the gauge kinetic term as

$$\frac{1}{g_i^2} = \frac{r^{i\alpha}\tau_\alpha}{2\pi M_P} \times \begin{cases} 1 & U(1) \\ 1/2 & SU(N) \end{cases} \quad (14)$$

- The couplings $q_{i\alpha}$ appearing in the (Stückelberg) mass terms for the $U(1)$ gauge fields A_i in Eq. (8) are given by

$$q_{i\alpha} \equiv \ell_s^{-2} \int_{D_i} \omega_\alpha \wedge \frac{\mathcal{F}}{2\pi} = \ell_s^{-4} \int_{D_i} \omega_\alpha \wedge \ell_s^2 \mathcal{F}, \quad (15)$$

where \mathcal{F} is the gauge flux on D_i . Axions c_α experiencing such a coupling disappear from the low energy EFT because they are eaten by the corresponding $U(1)$ gauge boson [25, 26].

Thus, it appears from Eq. (8), that IIB string flux compactifications have potentially many, $h^{1,1} - d$, axion candidates c_i , where d is the number of $U(1)$ bosons A_i which get a Stückelberg mass by eating the associated axions. However, before reaching this conclusion one has to consider perturbative and non-perturbative corrections to this tree-level result. In fact, such corrections are necessarily to be taken into account in the course of the stabilisation of the associated Kähler moduli τ_i . Importantly, the mechanisms to fix the τ_i may also generate large masses for the corresponding axions c_i [27, 28], as we will see next.

2.2. Axions and ALPs in moduli stabilised IIB string flux compactifications

In IIB string flux compactifications, the tree-level superpotential [29]

$$W_{\text{tree}} = \int_X G_3 \wedge \Omega. \quad (16)$$

which is generated by turning on background fluxes of the form $G_3 = F_3 + iSH_3$, where $F_3 = dC_2$ and $H_3 = dB_2$, does not depend on the Kähler moduli, but on the dilaton C and the complex structure moduli U_α . This implies that the dilaton and the complex structure moduli can be fixed at tree-level by imposing vanishing F-term conditions [30]. By appropriate tuning of the internal fluxes, one can always fix the dilaton such that the string coupling, $g_s = 1/\text{Re}(S)$, is in the perturbative regime. Further effects from fixing S and U are then parametrised by the flux-dependent constant $W_0 = \langle W_{\text{tree}} \rangle$ and by an overall factor in the F-term scalar potential arising from the S and U dependent part in the tree-level Kähler potential (9).

The Kähler moduli τ_i , however, remain precisely massless at leading semiclassical order, because of the no-scale structure of K_{tree} . Their stabilisation requires taking into account perturbative (p) and nonperturbative (np) α' and g_s corrections to the tree-level result,

$$W = W_{\text{tree}} + \delta W_{\text{np}}, \quad K = K_{\text{tree}} + \delta K_{\text{p}} + \delta K_{\text{np}}, \quad (17)$$

eventually leading also to non-trivial potentials, i.e. masses, for the axions c_i associated with the scalars τ_i . These masses can arise, however, only via the non-perturbative corrections to

the superpotential, δW_{np} , and to the Kähler potential, δK_{np} : only those can break the shift symmetry of the axions c_i .

The respective order of magnitude of the perturbative versus the nonperturbative corrections to the scalar potential is set by W_0 . The main mechanisms proposed for τ moduli stabilisation and their consequences for the physics of their associated c axions can be characterised as follows:

- The first scenario of Kähler moduli stabilisation in IIB string compactifications neglected the corrections to the Kähler potential and considered only the non-perturbative corrections to the superpotential [31],

$$\delta W_{\text{np}} = \sum_{i=1}^{h^{1,1}} A_i(S, U) e^{-a_i T_i}, \quad (18)$$

which arises by $ED3$ instantons (in which case $a_i = 2\pi$) or by stacks of $D7$ branes supporting a condensing gauge theory (for which $a_i = 6\pi/b_0$ with b_0 being the coefficient of the one-loop beta function), wrapping the four-cycles D_i . The threshold effects A_i can be considered as $\mathcal{O}(1)$ constants since they depend on the complex structure moduli which are flux-stabilised at tree-level. The four-cycle volumina are then fixed at

$$\tau_i \sim \frac{1}{a_i} \ln \left(\frac{W_0}{A_i} \right). \quad (19)$$

Thus, W_0 has to be fine-tuned to extremely small values, $W_0 \ll 1$, in order that the volume is large, $\mathcal{V} \sim \tau_i^{3/2} \gg 1$, the latter being a prerequisite of the underlying supergravity approximation. Therefore, in these scenarios it is extremely hard to lower the string scale (12) much below the Planck scale. Moreover, since the axionic shift symmetry $c_i \rightarrow c_i + \text{constant}$ of all axions is broken non-perturbatively by Eq. (18), all the axion candidates get a large mass of the order of the mass of the particle excitations of the associated Kähler moduli (the so-called “saxions”),

$$m_{c_i} \sim m_{\tau_i} \sim a_i W_0 M_P / \mathcal{V},$$

leaving no candidate for a QCD axion [16], let alone a light ALP.

- This can be avoided if the non-perturbative effects arise from wrapping an *ample* four-cycle,

$$D_{\text{am}} = \sum_{i=1}^{h^{1,1}} \lambda_i D_i, \text{ with } \lambda_i > 0 \ \forall i = 1, \dots, h^{1,1}, \quad (20)$$

in terms of a basis $\{D_i\}$ of $H_4(X, \mathbb{Z})$, resulting in a contribution to the superpotential of the form

$$\delta W_{\text{np}} = A e^{-a T_{\text{am}}} = A e^{-a \sum_{i=1}^{h^{1,1}} \lambda_i T_i}. \quad (21)$$

In this case, by fine-tuning $W_0 \sim A e^{-a T_{\text{am}}} \ll 1$, this single non-perturbative effect can generate a minimum for all the Kähler moduli τ_i within the regime of validity of the EFT [32]. However it can lift only one axion corresponding to the imaginary part of the ample divisor modulus: $c_{\text{am}} = \text{Im}(T_{\text{am}})$. All the remaining $h^{1,1} - 1$ axions are massless at leading order and develop a potential only via tiny higher order instanton effects of the form [21]:

$$\delta W_{\text{np}} = A e^{-a T_{\text{am}}} + \sum_{i=1}^{h^{1,1}-1} A_i e^{-n_i a_i T_i}, \quad (22)$$

where T_i is a combination of moduli orthogonal to $T_{\text{am}} \ \forall i = 1, \dots, h^{1,1} - 1$.

Thus, this moduli stabilisation scenario gives rise to an axiverse [20]: possibly many, $h^{1,1} - 1$, axions which acquire a mass spectrum which is logarithmically hierarchical. This number may still be diminished in case that some of the axions are eaten by the Stückelberg mechanism to provide masses to brane-localised $U(1)$ gauge bosons.

Two main concerns regarding the microscopic realisation of this scenario are related to the difficulty to find an ample divisor which is rigid (and so definitely receives non-perturbative effects), and the possibility to choose gauge fluxes that avoid chiral intersections between the instanton and the visible sector. Moreover, it should be noted that in this case again very large volumina are not possible: the string scale (12) is expected not much below the Planck scale, of order the GUT scale.

- The latter difficulties are avoided in the LARGE volume scenario (LVS), which realizes the possibility of exponentially large volumina for generic values of $W_0 \sim \mathcal{O}(1)$ [33] and allows for the construction of explicit globally and locally consistent Calabi-Yau examples with magnetised D7-branes, realising MSSM or GUT like chiral extensions of the standard model [34, 35].

The LVS requires the existence of a single *del Pezzo* four-cycle τ_{dP} which guarantees the emergence of a non-perturbative contribution to the superpotential

$$\delta W_{\text{np}} = A e^{-a T_{\text{dP}}} \quad (23)$$

via an $ED3$ instanton or gaugino condensation. This effect fixes τ_{dP} at a small size and gives the corresponding saxion and axion a large mass [36],

$$m_{\tau_{\text{dP}}} \sim m_{c_{\text{dP}}} \sim \frac{W_0 \sqrt{\ln \mathcal{V}}}{\mathcal{V}} M_P. \quad (24)$$

All the other τ moduli are stabilised perturbatively by α' or g_s effects or by D-terms arising from magnetised branes. The exponential large volume emerges from an interplay between the non-perturbative contribution associated with del Pezzo cycle and the leading α' correction [37]:

$$\delta K_P \simeq -\frac{\zeta}{g_s^{3/2} \mathcal{V}}, \text{ with } \zeta \propto (h^{1,2} - h^{1,1}), \quad (25)$$

which yields, for $h^{1,2} > h^{1,1} > 1$ (i.e. negative Euler number), a supersymmetry-breaking anti de Sitter (AdS) minimum at exponentially large volume [33]:

$$\mathcal{V} \sim W_0 e^{a \tau_{\text{dP}}}. \quad (26)$$

Hence only one axion, $c_{\text{dP}} = \text{Im}(T_{\text{dP}})$, becomes heavy whereas all the other ones (except those eaten up by anomalous $U(1)$ s, whose scalar partners are fixed by the above mentioned D-terms) remain light and develop a potential via subleading higher order instanton effects. The simplest version of a standard LVS is build upon a Swiss-cheese Calabi-Yau three-fold with volume given by [38]:

$$\mathcal{V} = \alpha \left(\tau_b^{3/2} - \sum_{i=1}^{h^{1,1}-1} \gamma_i \tau_i^{3/2} \right). \quad (27)$$

It is dominated by the exponentially large volume τ_b of one of the four-cycles, D_b , while all the other, $h^{1,1} - 1$, four-cycles are small.

Importantly, the realisation of an LVS requires $h^{1,1} \geq 4$ (cf. Fig. 2): a single del Pezzo divisor D_{dP} to support the non-perturbative effects, one big four-cycle τ_b to parametrise

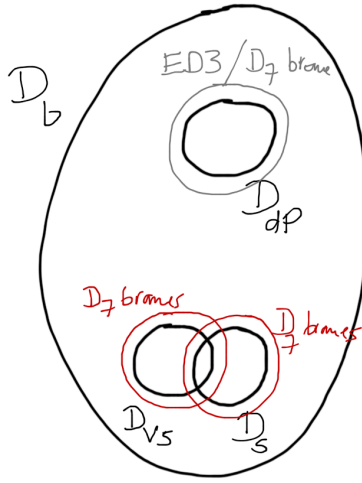


Figure 2. Schematic picture of the simplest Swiss-cheese setup of a compactification in the LARGE volume scenario: a big four-cycle (=divisor) D_b with exponentially large volume $\tau_b \sim \mathcal{V}^{2/3}$; a del Pezzo divisor D_{dp} supporting the leading non-perturbative effect ($ED3$ instanton or gaugino condensation of a pure Yang-Mills theory described by a stack of $D7$ branes); a rigid divisor D_{vs} supporting the stack of $D7$ branes describing the visible sector; a rigid divisor D_s intersecting with D_{vs} and supporting a stack of magnetised $D7$ branes for D-term stabilisation.

the large volume, $\mathcal{V} \sim \tau_b^{3/2}$, one small rigid four-cycle D_{vs} to support the stack of space-time filling magnetised $D7$ -branes corresponding to (a suitable extension of) the MSSM and intersecting with another small cycle D_s whose brane setup is there to provide D-terms which stabilise τ_{vs} and thus the visible sector gauge coupling $g_{vs}^{-2} \sim \tau_{vs}$. Two of the $h^{1,1} \geq 4$ axion-like fields disappear from the low energy spectrum, however: c_{dp} , as explained above, and c_s , which is eaten by the Stückelberg mechanism. Thus, realisations of the LVS axiverse involve at least two light axions: one QCD axion candidate plus one ALP [22]. More generic models for $h^{1,1}$ very large will include an arbitrarily large number of ALPs.

The main scales in the model are:

$$\begin{aligned} M_s &= \frac{M_P}{\sqrt{4\pi\mathcal{V}}} \sim 10^{10} \text{ GeV}, m_{\tau_s} \sim \frac{M_P}{\mathcal{V}^{1/2}} \sim 10^{10} \text{ GeV}, M_{KK} \sim \frac{M_P}{\mathcal{V}^{2/3}} \sim 10^9 \text{ GeV}, \\ m_{\tau_{dp}} &\sim \sqrt{g_s} W_0 \frac{M_P}{\mathcal{V}} \ln \mathcal{V} \sim 30 \text{ TeV}, m_{\text{soft}} \sim m_{3/2} \sim \sqrt{g_s} W_0 \frac{M_P}{\mathcal{V}} \sim 1 \text{ TeV}, \\ m_{\tau_{vs}} &\sim \alpha_{vs} m_{3/2} \sim 40 \text{ GeV}, m_{\tau_b} \sim \frac{M_P}{\mathcal{V}^{3/2}} \sim 0.1 \text{ MeV}. \end{aligned} \quad (28)$$

The numerical values have been given for generic values of the underlying parameters, $g_s \sim 0.1, W_0 \sim 1$, and for a volume $\mathcal{V} \sim 10^{14}$, demonstrating that, for an intermediate string scale, the LVS naturally realises TeV-scale SUSY.

- Finally, the contributions from both δK_p and non-perturbative terms in the superpotential δW_{np} from gaugino condensation on stacks of 4-cycle wrapping $D7$ branes allow for a 2nd branch of supersymmetry breaking “Kähler uplifted” vacua distinct from the LVS branch. These vacua can be either AdS, Minkowski or dS by themselves without any need for further ‘external’ sources of uplifting. Their possible existence was first pointed out in [39], while they were shown to be viable in producing controlled large-volume vacua for all moduli in a supergravity analysis [40, 41]. Recently, [42] provided the first explicit global and consistent F-theory constructions of such Kähler uplifted dS vacua. On Swiss-cheese Calabi-Yau threefolds the Kähler uplifting branch generates a class of minima for the volume moduli, where $\tau_b \sim N_b$ with N_b the rank of the condensing gauge group of the $D7$ brane stack wrapping the large divisor D_b . The blow-up Kähler moduli $\tau_i \sim N_i$ are stabilised at smaller volume dictated by the rank of the corresponding condensing gauge groups. For ranks $N_b \sim 30 \div 100$ this leads to stabilisation of all Kähler moduli at an overall volume $\mathcal{V} \sim N_b^{3/2} \sim 10^{2 \div 3}$, with volume moduli masses suppressed by a factor $1/\mathcal{V}$

compared the masses of the complex structure moduli and the axio-dilaton from fluxes. So far, stabilisation of the $h^{1,1}$ τ_i moduli utilises an identical number of non-perturbative contributions to the superpotential from gaugino condensation. Hence, all associated c_i axions are rendered massive with mass scales tied to their scalar partners. Kähler uplifting with an (partially) ample divisor utilising less than $h^{1,1}$ instanton effects for stabilisation and in turn realising an axiverse represents an open question.

2.3. Axion and ALP decay constants and their couplings to visible sector particles in the LVS

Let us now for the remainder of this review concentrate on the LVS, because, as we have seen, it predicts the existence of a QCD axion candidate plus at least one ALP. Their decay constants f_{a_i} and their couplings C_{ij} to gauge bosons (or matter fields) j can be read off from the matching of the prediction (8) with the generalisation of the generic low energy effective Lagrangian (1) to the case of many axion-like fields, by transforming in the former from the original basis $\{c_i\}$ to a basis of fields $\{a_i\}$ with canonically normalised kinetic terms [22]. For the simple Swiss-cheese Calabi-Yau setup discussed above and illustrated in Fig. 2, the original axion fields c_b and c_{vs} can be written in terms of the canonically normalised fields a_b and a_{vs} as:

$$\frac{c_b}{\tau_b} \simeq \mathcal{O}(1) a_b + \mathcal{O}\left(\tau_{vs}^{3/4} \mathcal{V}^{-1/2}\right) a_{vs}, \quad \frac{c_{vs}}{\tau_{vs}} \simeq \mathcal{O}(1) a_b + \mathcal{O}\left(\tau_{vs}^{-3/4} \mathcal{V}^{1/2}\right) a_{vs}, \quad (29)$$

leading to

$$f_{a_b} \simeq \frac{M_{KK}}{4\pi}, \quad f_{a_{vs}} \simeq \frac{M_s}{\sqrt{4\pi\tau_{vs}}^{1/4}}, \quad (30)$$

$$C_{bb} \simeq \mathcal{O}(1), \quad C_{vsb} \simeq \mathcal{O}\left(\mathcal{V}^{-1/3}\right), \quad C_{bvs} \simeq \mathcal{O}\left(\mathcal{V}^{-2/3}\right), \quad C_{vs vs} \simeq \mathcal{O}(1). \quad (31)$$

The axion a_{vs} can be identified with the QCD axion. A range of values for the proper QCD axion decay constant, spanning the classic QCD axion window,

$$f_a \equiv \frac{f_{a_{vs}}}{C_{vs vs}} \sim M_s \sim \sqrt{\frac{M_P m_{3/2}}{g_s^{1/2} W_0}} \sim 10^{9 \div 12} \text{ GeV}, \quad (32)$$

are then possible depending upon the exact numerical coefficients, the value of the gravitino mass $m_{3/2} \gtrsim \text{TeV}$, and the tuning of $g_s \lesssim 1$ and W_0 .

The large cycle ALP a_b has a smaller decay constant $f_{a_b} \sim M_{KK}$, but its coupling to the standard Model gauge bosons is completely negligible, $C_{bvs} \simeq \mathcal{O}(\mathcal{V}^{-2/3})$.

More ALPs, but with decay constant similar to the one of the QCD axion,

$$f_{\text{ALP}_i} \sim f_{a_{vs}} \sim M_s, \quad (33)$$

and coupling to gauge bosons other than the gluon similar to the one of the QCD axion,

$$C_{\text{ALP}_i vs} \sim C_{vs vs} \sim \mathcal{O}(1) \Rightarrow g_{i\gamma} \equiv \frac{\alpha}{2\pi f_{a_i}} C_{i\gamma} \sim 10^{-15} \div 10^{-11} \text{ GeV}^{-1}, \quad (34)$$

are obtained, if there are more small cycles D_{vs_i} intersecting the visible branes, but without introducing additional D-term conditions. Their masses are expected to be smaller than the mass of the QCD axion and to be distributed logarithmically hierarchical [22],

$$m_{\text{ALP}_i} \sim e^{-n\pi\tau_{\text{ALP}_i}} \times \begin{cases} M_P, & \text{for } \delta W_{\text{np}} \text{ terms or QCD-like masses,} \\ m_{3/2}, & \text{for } \delta K_{\text{np}} \text{ terms.} \end{cases} \quad (35)$$

These findings have been reproduced also in explicit constructions of LVS examples, exploiting concrete Calabi-Yau orientifolds and semi-realistic $D7$ -brane and flux setups [22, 35]. That paper considers also LVS variants with asymmetric fibred Calabi-Yaus and sequestered SUSY breaking, allowing for more freedom in $m_{3/2}$ and m_{soft} , at the expense of more fine tuning in W_0 , and discusses briefly the cosmological evolution of the different LVS incarnations.

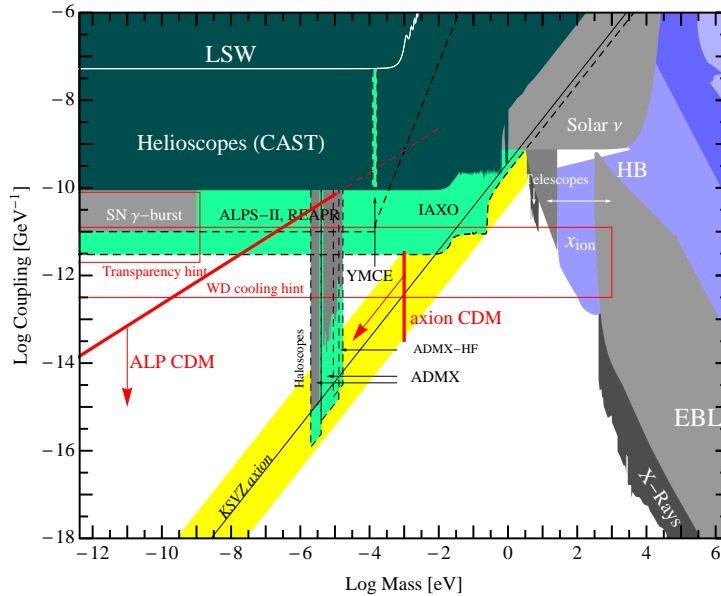


Figure 3. Axion and ALP coupling to photons, $g_{i\gamma} \equiv \alpha C_{i\gamma}/(2\pi f_{a_i})$, vs. its mass (adapted by Javier Redondo [49] from Refs. [50, 51]). The yellow band is the generic prediction for the QCD axion, exploiting Eqs. (2) and (4), which relate its mass with its coupling to photons.

3. Opportunities to probe the intermediate string scale LVS

3.1. Haloscope searches

We have seen, that the LVS predicts – for the least fine-tuning of fluxes, such that $g_s \sim 0.1$ and $W_0 \sim 1$, and a TeVish gravitino mass – an intermediate string scale and thus a QCD axion in the classic window, cf. Eq. (32). For decay constants in the upper part of this window, $f_a \gtrsim 10^{11 \div 12}$ GeV, the QCD axion is expected to contribute substantially to the cold dark matter in the universe, see Eq. (6). Therefore, the intermediate string scale LVS can be probed by haloscope searches for axion cold dark matter [43] such as ADMX [44, 45, 46, 47]. These experiments exploit the coupling (4) by searching for the signal of dark matter axion to photon conversions in a narrow bandwidth microwave cavity sitting in a strong magnetic field. As can be seen from the light green area in Fig. 3 labelled as “Haloscopes”, a substantial range of the QCD axion dark matter parameter range will be probed by ADMX and other haloscopes [48] in the next decade.

3.2. Helioscope searches

A complementary search for the QCD axion in the lower part of the classic window, $f_a \gtrsim 10^{9 \div 10}$ GeV, can be conducted with the next generation of axion helioscopes [43], in which one tries to detect solar axions by their conversion into photons inside of a strong magnet pointed towards the sun. Indeed, the projected sensitivity of the proposed International Axion Observatory IAXO [52] covers nicely a part of QCD axion parameter space which will not be covered by the haloscope searches, as can be seen in Fig. 3.

A very welcome feature of helioscopes is that they do not lose sensitivity towards low masses: their projected sensitivity are best and stay constant at small masses, see Fig. 3. That means, with IAXO one may also probe the LVS axiverse, in particular the possible existence of more ALPs with approximate the same coupling to photons as the QCD axion.

This is very important in view of recent tantalising astrophysical hints, such as the anomalous transparency of the Universe for TeV photons [53] and the anomalous cooling of white

dwarfs [54, 55], which could be explained by the existence of an ALP with decay constants, couplings and mass [56, 57, 58, 59, 60, 61, 62]:

$$\frac{f_{a_i}}{C_{ie}} \simeq (0.7 \div 2.6) \times 10^9 \text{ GeV}, \quad \frac{f_{a_i}}{C_{i\gamma}} \sim 10^8 \text{ GeV}, \quad m_{a_i} \lesssim 10^{-9} \div 10^{-10} \text{ eV}, \quad (36)$$

compatible with the prediction of an intermediate string scale LVS axiverse with $C_{i\gamma}/C_{ie} \sim 10$. The projected sensitivity of IAXO nicely overlaps with the ALP parameter region required to explain these hints, see Fig. 3.

3.3. Light-shining-through-walls searches

Intriguingly, this parameter region for ALPs can also be partially probed by purely laboratory based light-shining-through-walls experiments [63], where laser photons are sent along a strong magnetic field, allowing for their conversion into ALPs, which may then reconvert in the strong magnetic field behind a blocking wall into photons, apparently shining through the wall and susceptible to detection. The projected sensitivities of the proposed experiments ALPS-II at DESY and REAPR at Fermilab partially cover the expectations (34) from an intermediate string scale LVS axiverse and from the hints (36) from astrophysics, see Figs. 1 and 3.

4. Conclusions

String phenomenology holds the promise of an axiverse – the QCD axion plus a (possibly large) number of further ultralight axion-like particles, possibly populating each decade of mass down to the Hubble scale, 10^{-33} eV. However, although a plenitude of axion-like fields is a generic prediction of string theory, there may be few or no light axions remaining once constraints such as tadpole cancellation and moduli stabilisation are taken into account.

Interestingly, the promise of an axiverse seems to be fulfilled in the LARGE Volume Scenario (LVS) of IIB string flux compactifications. In fact, the simplest globally consistent LVS constructions with magnetised $D7$ -branes and chirality have at least two light axions: a QCD axion candidate, with a decay constant of order the string scale, which is intermediate, $f_a \sim M_s \sim M_P/\sqrt{V} \sim (M_P m_{3/2}/W_0)^{1/2} \sim 10^{9 \div 12}$ GeV, for a TeV scale gravitino mass $m_{3/2}$ and an expectation value of the flux induced tree-level expectation value W_0 of the superpotential of order one, plus a nearly decoupled superlight axion-like particle. In setups where the small size branes describing the visible sector are wrapping more than the minimally required two intersecting four-cycles, there are more ultralight axion-like particles which have approximately the same decay constant and coupling to the photon as the QCD axion candidate.

At both ends of the above range of the decay constant there are exciting phenomenological opportunities. For $f_a \sim 10^{11 \div 12}$ GeV, the QCD axion can be the dominant part of dark matter and be detected in haloscopes exploiting microwave cavities. For $f_a \sim 10^{9 \div 10}$ GeV, the additional ALPs could explain astrophysical anomalies and be searched for in the upcoming generation of helioscopes or light-shining-through-a-wall experiments.

Acknowledgments

I would like to thank Michele Cicoli and Mark Goodsell for the great collaboration on the topics in this review and Alexander Westphal for many enlightning discussions.

References

- [1] Weinberg S 1978 *Phys. Rev. Lett.* **40** 223
- [2] Wilczek F 1978 *Phys. Rev. Lett.* **40** 279
- [3] Peccei R D and Quinn H R 1977 *Phys. Rev. Lett.* **38** 1440
- [4] Georgi H, Kaplan D B and Randall L 1986 *Phys. Lett. B* **169** 73
- [5] Kim J E 1979 *Phys. Rev. Lett.* **43** 103

- [6] Dine M, Fischler W and Srednicki M 1981 *Phys. Lett. B* **104** 199
- [7] Shifman M A, Vainshtein A I and Zakharov V I 1980 *Nucl. Phys. B* **166** 493
- [8] Zhitnitsky A R 1980 *Sov. J. Nucl. Phys.* **31** 260 [1980 *Yad. Fiz.* **31** 497]
- [9] Bardeen W A and Tye S-H H 1978 *Phys. Lett. B* **74** 229
- [10] Kaplan D B 1985 *Nucl. Phys. B* **260** 215
- [11] Srednicki M 1985 *Nucl. Phys. B* **260** 689
- [12] Preskill J, Wise M B and Wilczek F 1983 *Phys. Lett. B* **120** 127
- [13] Abbott L F and Sikivie P 1983 *Phys. Lett. B* **120** 133
- [14] Dine M and Fischler W 1983 *Phys. Lett. B* **120** 137
- [15] Witten E 1984 *Phys. Lett. B* **149** 351
- [16] Conlon J P 2006 *JHEP* **0605** 078
- [17] Svrcek P and Witten E 2006 *JHEP* **0606** 051
- [18] Conlon J P 2006 *Phys. Rev. Lett.* **97** 261802
- [19] Choi K-S, Kim I-W and Kim J E 2007 *JHEP* **0703** 116
- [20] Arvanitaki A *et al.*, 2010 *Phys. Rev. D* **81** 123530
- [21] Acharya B S, Bobkov K and Kumar P 2010 *JHEP* **1011** 105
- [22] Cicoli M, Goodsell M and Ringwald A 2012 *Preprint* arXiv:1206.0819 [hep-th]
- [23] Grimm T W and Louis J 2004 *Nucl. Phys. B* **699** 387
- [24] Jockers H and Louis J 2005 *Nucl. Phys. B* **705** 167
- [25] Goodsell M, Jaeckel J, Redondo J and Ringwald A 2009 *JHEP* **0911** 027
- [26] Cicoli M, Goodsell M, Jaeckel J and Ringwald A 2011 *JHEP* **1107** 114
- [27] Banks T, Dine M and Gorbатов E 2004 *JHEP* **0408** 058
- [28] Donoghue J F 2004 *Phys. Rev. D* **69** 106012 [2004 *Erratum-ibid. D* **69** 129901]
- [29] Gukov S, Vafa C and Witten E 2000 *Nucl. Phys. B* **584** 69
- [30] Giddings S B, Kachru S and Polchinski J 2002 *Phys. Rev. D* **66** 106006
- [31] Kachru S, Kallosh R, Linde A and Trivedi S P 2003 *Phys. Rev. D* **68** 046005
- [32] Bobkov K, Braun V, Kumar P and Raby S 2010 *JHEP* **1012** 056
- [33] Balasubramanian V, Berglund P, Conlon J P and Quevedo F 2005 *JHEP* **0503** 007
- [34] Cicoli M, Mayrhofer C and Valandro R 2012 *JHEP* **1202** 062
- [35] Cicoli M 2012 these proceedings
- [36] Conlon J P, Quevedo F and Suruliz K 2005 *JHEP* **0508** 007
- [37] Becker K, Becker M, Haack M and Louis J 2002 *JHEP* **0206** 060
- [38] Collinucci A *et al* 2009 *JHEP* **0907** 074
- [39] Balasubramanian V and Berglund P 2004 *JHEP* **0411** 085
- [40] Westphal A 2007 *JHEP* **0703** 102
- [41] Rummel M and Westphal A 2012 *JHEP* **1201** 020
- [42] Louis J, Rummel M, Valandro R and Westphal A *Preprint* arXiv:1208.3208 [hep-th]
- [43] Sikivie P 1983 *Phys. Rev. Lett.* **51** 1415 [1984 *Erratum-ibid.* **52** 695].
- [44] Asztalos S J *et al* 2001 *Phys. Rev. D* **64** 092003
- [45] Asztalos S J *et al* [The ADMX Collaboration] 2010 *Phys. Rev. Lett.* **104** 041301
- [46] Asztalos S J *et al* 2011 *Nucl. Instrum. Meth. A* **656** 39
- [47] Heilman J *et al* [ADMX Collaboration] 2010 *AIP Conf. Proc.* **1274** 115
- [48] Baker O K *et al* 2012 *Phys. Rev. D* **85** 035018
- [49] Redondo J 2012 private communication
- [50] Arias P, Cadamuro D, Goodsell M, Jaeckel J, Redondo J and Ringwald A 2012 *JCAP* **1206** 013
- [51] Hewett J L *et al.* 2012 *Preprint* arXiv:1205.2671 [hep-ex]
- [52] Irastorza I G *et al* 2011 *JCAP* **1106** 013
- [53] Horns D and Meyer M 2012 *JCAP* **1202** 033
- [54] Isern J, Garcia-Berro E, Torres S and Catalan S 2008 *Astrophys. J.* **682** L109
- [55] Isern J, Catalan S, Garcia-Berro E and Torres S 2009 *J. Phys. Conf. Ser.* **172** 012005
- [56] De Angelis A, Mansutti O, Roncadelli M 2007 *Phys. Rev. D* **76** 121301
- [57] Simet M, Hooper D, Serpico P D 2008 *Phys. Rev. D* **77** 063001
- [58] Sanchez-Conde M A *et al* 2009 *Phys. Rev. D* **79** 123511
- [59] Dominguez A, Sanchez-Conde M A and Prada F 2011 *JCAP* **1111** 020
- [60] Tavecchio F, Roncadelli M, Galanti G and Bonnoli G 2012 *Preprint* arXiv:1202.6529 [astro-ph.HE]
- [61] Isern J *et al.*, 2012 *Preprint* arXiv:1204.3565 [astro-ph.SR]
- [62] Horns D *et al* 2012 *Preprint* arXiv:1207.0776 [astro-ph.HE]
- [63] Redondo J and Ringwald A 2011 *Contemp. Phys.* **52** 211