

SYNMAG PHOTOMETRY: A FAST TOOL FOR CATALOG-LEVEL MATCHED COLORS OF EXTENDED SOURCES

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ABSTRACT

Obtaining reliable, matched photometry for galaxies imaged by different observatories represents a key challenge in the era of wide-field surveys spanning more than several hundred square degrees. Methods such as flux fitting, profile fitting, and PSF homogenization followed by matched-aperture photometry are all computationally expensive. We present an alternative solution called “synthetic aperture photometry” that exploits galaxy profile fits in one band to efficiently model the observed, point-spread-function-convolved light profile in other bands and predict the flux in arbitrarily sized apertures. Because aperture magnitudes are the most widely tabulated flux measurements in survey catalogs, producing synthetic aperture magnitudes (SYNMAGs) enables very fast matched photometry at the catalog level, without reprocessing imaging data. We make our code public and apply it to obtain matched photometry between Sloan Digital Sky Survey *ugriz* and UKIDSS *YJHK* imaging, recovering red-sequence colors and photometric redshifts with a scatter and accuracy as good as if not better than FWHM-homogenized photometry from the GAMA Survey. Finally, we list some specific measurements that upcoming surveys could make available to facilitate and ease the use of SYNMAGs.

Key words: galaxies; photometry – methods; data analysis – techniques; photometric

Online-only material: color figures

1. INTRODUCTION

Astronomy is entering an era of truly panoramic deep imaging surveys. Building on the legacy from the Sloan Digital Sky Survey (SDSS; York et al. 2000; Abazajian et al. 2009), the coming years will see a rapidly expanding patchwork of overlapping imaging campaigns, with surveys like DES,⁶ HSC,⁷ KIDS,⁸ VIKINGS,⁹ and eventually Euclid and LSST¹⁰ providing hundreds of square degrees of new data on rapid timescales. Fast and reliable methods for combining information from these disparate surveys are clearly needed.

A key challenge is measuring reliable matched photometry that samples, as closely as possible, the flux emitted from the same regions of a galaxy as observed across the many wavebands sampled by these data sets. For galaxy studies, the goal is to obtain accurate colors and spectral energy distributions (SEDs) over a large wavelength range in order to derive photometric redshifts (photo-zs) and physical properties like stellar mass (M_*) from template fitting. But as extended sources, galaxies present particular challenges for matched photometry because their range in size, shape, and surface brightness is not easily accounted for by a single model or choice of aperture.

Three basic methods have been developed for obtaining matched photometry for extended sources from a set of multi-band images with different seeing, scales, and other properties. The first and perhaps most common is point-spread function

(PSF) homogenization followed by matched-aperture photometry. Here, all of the overlapping images in a set of observations are convolved to the PSF of the band with the worst seeing. The convolved images are then registered (often with interpolation to the same pixel grid) and the flux through a common aperture is measured on the pixel data in each band at the location of every source. This method has been studied in detail by Hildebrandt et al. (2012) who stress the value of spatially dependent convolution kernels that are designed to make the PSF uniform across every image in all bands. The method has been successfully applied to a variety of recent data sets including COSMOS (Capak et al. 2007) and GAMA (Hill et al. 2011). The disadvantages include the loss of information resulting from the PSF degradation, the introduction of correlated noise resulting from regridding and interpolation, and the fact that aperture flux measurements (especially for faint sources) are not ideal flux estimators.

The second method for obtaining matched photometry is “flux fitting.” This approach is often adopted when the PSF sizes of the imaging data are strongly mismatched, for example when comparing photometry from the *Hubble Space Telescope* to ground-based imaging or much lower resolution near-IR data. In this case, the high-resolution image is first convolved to the lower resolution seeing. The pixel-by-pixel flux of the “blurred” objects identified in the high-resolution image are then compared directly to the pixel flux in the low-resolution image. A minimization scheme determines the flux ratio between the two bands for every high-resolution source, fitting all neighbors simultaneously to account for overlapping profiles in the low-resolution image. This technique forms the basis of the TFIT package (Laidler et al. 2007) and the CONVPHOT software (de Santis et al. 2007) and has been

⁶ Dark Energy Survey.

⁷ Hyper-Suprime Cam Survey.

⁸ Kilo-Degree Survey.

⁹ VISTA Kilo-degree Infrared Galaxy Survey.

¹⁰ Large Synoptic Sky Telescope.

used widely (e.g., Papovich et al. 2001; Labbé et al. 2005; Grazian et al. 2006). It has the advantage of making use of the available spatial information, but does assume—as all methods do to some extent—that the shape of the flux profiles is independent of wavelength. In this case, the shape is determined by the pixel-to-pixel flux in the single, high-resolution band. It can thus be subject to noise spikes and other anomalies.

The final method is “profile fitting” which generalizes the “flux fitting” defined above. The most widely used example is the photometry from SDSS, specifically the ModelMag photometry which is described in Strauss et al. (2002), with more information available on the internet.¹¹ Here a smooth intrinsic profile is determined for every source, usually in the deepest band. In each of the other bands, this profile is convolved with the local PSF and scaled to the pixel flux in that band. Assuming the profile is wavelength independent and correct, this method maximizes the signal-to-noise on the flux measurement and obtained colors. But the adopted profile need not be a perfect match to the observed profile. In this case it can be thought of as a spatial weighting function applied to the total flux. This is the approach taken by Kuijken (2008) for measuring the “Gaussian-aperture-and-PSF” (GaaP) flux, a seeing-independent flux estimator. Here the profile fitting is done to the shapelet basis function with fits obtained in every band. The shapelets can then be convolved analytically to a common PSF under which the GaaP flux and resulting colors are defined. Since the shapelet fitting is performed in every band, however, the photometric scatter should be larger than in methods that fit all fluxes to a single intrinsic profile.

Each of these methods has strengths and weaknesses, but all of them suffer from a major problem when considering the size and pace at which new imaging data are becoming available. Multiband imaging data over thousands of square degrees can quickly sum to ~ 10 terabytes or more. If new data are continuously being added, keeping the matched photometry up to date can be extremely expensive computationally. In addition to the raw computing time, the human cost of mastering and manipulating so much raw data is significant. Instead, it would be valuable to have the option of measuring the photometric properties of sources in a given survey and filter one time only, but in a way that allows for a straightforward comparison with other filters and surveys without revisiting the pixel data.

In this work, we present an approach with this goal in mind that delivers “synthetic aperture magnitudes” (SYNMAGs), enabling very fast matched photometry with reasonable precision. Higher precision results can be obtained with a full analysis of all available pixel data (for example, using profile-fitting methods), but SYNMAGs provide a fast alternative until such an analysis can be undertaken. Our method works entirely at the catalog level. We require recorded parameters for profile fitting in at least one band, with measured aperture magnitudes (of arbitrary size) and PSFs recorded for the other bands. As a test case, we derive matched photometry from the combination of SDSS and the UKIRT Infrared Deep Sky Survey’s Large Area Survey (UKIDSS/LAS) and show that its performance is as good, if not better, than PSF-homogenized methods. The SYNMAG technique is easily applied to any data set that overlaps with SDSS, but is applicable generally to surveys employing profile-fitting photometry in at least one band, such as CS82 (T. Erben et al., in preparation), HSC, and DES. We make the

SYNMAG code publicly available. In a follow-up paper (T. D. Higgs et al., in preparation), we present a second methodology specifically tuned to the SDSS–UKIDSS comparison that offers additional flexibility.

An overview of the SYNMAG method is presented in Section 2, with specific notes relevant to SDSS given in Section 2.2. The use of multi-Gaussian expansions (MGEs) is described in Section 2.4, while Section 3 presents various tests of SYNMAGs applied to real data. In Section 4 we conclude with some discussion of “application-oriented photometry” and describe how future surveys can provide catalog information to facilitate the use of SYNMAGs. The Appendix provides recipes for applying the SYNMAG software to real data.

2. METHOD

2.1. Overview

We begin with an overview of our approach. While our specific motivation is to apply this method to obtain matched photometry between SDSS and UKIDSS imaging data sets, the methodology will be appropriate for many other applications. As in nearly all treatments of multiband photometry including the use of matched-aperture photometry, we will assume there are no color gradients. This is obviously a poor assumption because it ignores how color depends on aperture size and obscures the meaning of an integrated color (see Section 4), but is standard practice and simplifies the problem.

Consider two imaging data sets that we will refer to as D_A and D_B . They account for two sets of filter bands, obtained in different conditions with different PSFs, perhaps on different telescopes. In the specific case we describe later, D_A will refer to the SDSS DR7 data set and D_B to UKIDSS. Our method requires that profile fitting has been performed on the imaging data from at least one of the filter bands in one of the data sets. We will assume model profile fit parameters are available in the D_A catalog and will choose a definitive band for these profiles, p , which is typically the deepest band in D_A . In SDSS, p is defined as the r band. Over all filter bands, i , available from both data sets, our goal is to obtain a reliable set of matched colors, C_{pi} , where $C_{pi} = m_p - m_i$.

We assume profile fits are not available in D_B , otherwise matched colors could be obtained by comparing the profile fits in D_A to D_B . Instead, we assume that only aperture photometry has been measured and recorded in the object catalogs for D_B . This is often the case because aperture photometry is far easier to perform than profile fitting.

For a given band, i , we require a catalog measurement from D_B of the spatially dependent PSF, or PSF_i . We then convolve the fitted 2D profile Σ_p from D_A with PSF_i and integrate the result from $R = 0$ to $R = R_{\text{aper}}$, where R_{aper} is the radius used to define the aperture photometry recorded in D_B . Note that we have *not* assumed circular symmetry. While the apertures used are circular,¹² our method works with 2D profile fits that describe the varied, projected axis ratios of galaxies on the sky. We do assume that the true flux is well described by smoothly declining profiles, however. In this way, we have “synthesized” the aperture flux that would have been measured in the p band (from D_A) under the same seeing conditions as the i band (from D_B) and within the same aperture as used to measure aperture photometry in the D_B catalog. We use $m_{p,\text{SYN},i}$ to

¹¹ See <http://www.sdss.org/dr7/algorithms/photometry.html> and <http://www.astro.Princeton.EDU/~rhl/photo-lite.pdf>.

¹² In practice, our method could be extended to work with non-circular apertures.

refer to the resulting synthetic magnitude or “SYNMAG.” We can then construct the full set of colors across D_A and D_B as $C_{pi} = m_{p,\text{SYN},i} - m_{i,\text{aper}}$.

In some cases, as in SDSS, colors internal to D_A or D_B may already be available and more reliable than the SYNMAG colors. These can be substituted for the SYNMAG colors when available because the assumption of no color gradients means that, given two filter bands, a galaxy by definition has only one color. In practice, of course, SYNMAG colors will differ from other techniques (e.g., ModelMag colors), but we emphasize that our goal is not to obtain the *best* set of matched colors, but the optimal one given constraints on both computational time and the difficulty of downloading, understanding, and organizing large imaging data sets. Taking advantage of pre-computed colors internal to one of the data sets can lead to an overall improvement in the characterization of a galaxy’s true colors. It is also possible to obtain colors from profile-fitting magnitudes derived in different bands and even with different profile shapes by comparing the flux integral of the fit profiles at a radius that encloses most of the light. If the assumed profile shapes are significantly different, however, biases may be introduced¹³ that depend on galaxy morphology and concentration.

We can also use the set of colors, C_{pi} , to generate a set of PSF-matched magnitudes. Let us assume that well-defined “total” magnitudes, $m_{p,\text{tot}}$, are available for the p band in D_A , perhaps again via profile fitting. In SDSS, these could be defined as r_{ModelMag} or $r_{\text{CMModelMag}}$. We can then define a new set of total-magnitude, PSF-matched photometry for each band, i , as $m_{i,\text{tot}} = m_{p,\text{tot}} - C_{pi}$.

Finally, we note that in the typical D_B data set, aperture photometry is performed in several apertures of varying size. It is possible to compute SYNMAGs and resulting colors for each of these apertures. A weighted average could produce a better color measurement, and in the limit of very many apertures, this is equivalent to fitting the intrinsic profile, Σ_p , to the circularly averaged radial flux profile in the band i . This moves closer to the construction of ModelMags in SDSS, in which Σ_p is fitted to the 2D pixel data in band i . We do not pursue this extension here, however, in part because aperture photometry spanning only a few pixels near the centers of objects requires a careful treatment of the flux variation *within* pixels. Even with well-resolved sources, these central aperture fluxes may not be computed reliably by many survey pipelines. Because they represent regions of high surface brightness, this may bias attempts to fit Σ_p directly to the tabulated aperture flux measurements.

2.2. Application to SDSS and UKIDSS

Synthetic aperture photometry is valuable in a variety of contexts, especially as an effective way to derive a first set of matched photometry ahead of a more detailed analysis involving the pixel data. Our original motivation, however, came from the desire to obtain optical through near-IR colors for galaxies targeted by the Baryon Oscillation Spectroscopic Survey (BOSS). We desired matched colors between the SDSS imaging and UKIDSS. In this section we provide relevant details about these specific data sets, especially SDSS. These are not only important for understanding the ways we have tested our approach, they will be useful for applying synthetic aperture photometry to the overlap between SDSS and other data sets.

The SDSS imaging is reduced and processed by the PHOTO pipeline, which has been packaged in `sdss3tools`. The source detection, flagging, profile fitting, and photometry performed by PHOTO is described on the SDSS Web site¹⁴ under “Algorithms.” The SDSS camera is described in Gunn et al. (1998), the telescope system in Gunn et al. (2006), and the filter system in Fukugita et al. (1996). A technical summary is provided in York et al. (2000). Most relevant for our purposes is the profile fitting and definition of ModelMags, which are the default recommendation for obtaining extended-source colors within SDSS.

A PSF-convolved, two-dimensional de Vaucouleurs (deV), exponential (Exp), and PSF profile is fit to every detected source in every band, *ugriz*. For the deV and Exp profiles, the best-fit effective radius (along the major axis), axis ratio and position angle, and normalization in the form of a total magnitude—as well as the estimated errors on these quantities—are recorded in the PHOTOBJ tables in the SDSS database. The standard form for the exponential profile, with r measured along the major axis, is given by

$$I_{\text{Exp}}(r) = I_{e,\text{Exp}} \exp \left[-\kappa_E \left(\frac{r}{R_e} - 1 \right) \right], \quad (1)$$

where $\kappa_E = 1.67835$, as needed to set the definition of Σ_e to the surface brightness at the half-light or effective radius, R_e . In PHOTO v5_6 the Exp profile is truncated at large radii in the following manner:

$$I_{\text{Exp}}^{\text{SDSS}}(r) = \begin{cases} I_{\text{Exp}} & r/R_e < 3 \\ I_{\text{Exp}}[1 - (r/R_e - 3)^2]^2 & 3 < r/R_e < 4. \\ 0 & r/R_e > 4 \end{cases} \quad (2)$$

For the deV profile, “softening” is applied in the center so that the profile is given as

$$I_{\text{deV}}(r) = I_{e,\text{deV}} \exp \left[-\kappa_d \left(\left[\left(\frac{r}{R_e} \right)^2 + 0.0004 \right]^{\frac{1}{8}} - 1 \right) \right], \quad (3)$$

with $\kappa_d = 7.66925$. The truncation for the deV profile is accomplished as follows:

$$I_{\text{deV}}^{\text{SDSS}}(r) = \begin{cases} I_{\text{deV}} & r/R_e < 7 \\ I_{\text{deV}}[1 - (r/R_e - 7)^2]^2 & 7 < r/R_e < 8. \\ 0 & r/R_e > 8 \end{cases} \quad (4)$$

These definitions can be found in `makeprof.c` in the PHOTO package. We note that some examples of specific SDSS profiles are contained in `phFitobj.h`, but the values given there for the total flux after accounting for the imposed truncation can only be reconciled if the truncation follows the form $[1 - (r/R_e - 3)^2]^{1.0}$ for Exp and $[1 - (r/R_e - 7)^2]^{1.0}$ for deV (that is, the terms are not squared). We instead take as correct the forms hard coded in `makeprof.c` and described by Equations (2) and (4). In terms of total magnitudes, the truncations lead to differences of 0.01–0.03 mag compared to the standard forms. Meanwhile, the applied profile softening at the center of the deV profile is more significant compared to the standard profile and can lead to differences of 0.05 mag.

The deepest band of imaging in SDSS is the r band which is used to determine the profile-matched photometry known as

¹³ A sense of the magnitude of such biases is given by the difference between Exponential and de Vaucouleurs magnitudes in SDSS, roughly $0.2 \text{ mag} \pm 0.1$ for galaxies near $r = 20$.

¹⁴ For DR7, see <http://www.sdss.org/dr7/algorithms/index.html>.

ModelMags. These are defined as follows. For each object, the better of the r -band Exp or deV fits as measured in all the filters is taken as the default profile for that object. This profile is then scaled in normalization (but in no other parameters) to the pixel data in the other bands at the source location. The integral of the total flux of the resulting best fit in each band defines the ModelMag in that band. ModelMag photometry provides galaxy colors with less scatter than other estimators, a strong motivation for the SYNMA approach which enables the use of ModelMags in conjunction with photometry matched to other imaging data sets.

2.3. Mock Image Implementation

In the next two sections we outline two methods for implementing synthetic aperture photometry. The first approach is more “brute force” in nature but provides a simple way of testing the more sophisticated MGE implementation that follows.

For each source, we take the profile fit parameters describing the intrinsic profile, Σ_p , from D_A catalog. These parameters should include a characterization of the profile shape (e.g., exponential, de Vaucouleurs, or more generally a Sérsic- n value), the effective radius, R_e , the profile normalization (e.g., the surface brightness, Σ_e , at R_e) and ideally the axis ratio. We then use these parameters to create a mock postage stamp image of Σ_p with zero noise. The mock image is convolved with the spatially dependent PSF of the target band in D_B and the flux in the now convolved profile is summed over the relevant aperture(s) defined in D_B , returning the SYNMAgs in the p band for that source, $m_{p,\text{SYN},i}$.

There are a number of obvious limitations to this implementation. First, our goal has been to achieve a very fast way of producing matched photometry. The image convolution here is the slowest step, and while we have avoided returning to the raw pixel data, the mock image implementation is not that different than working with the real data since both require convolutions performed on 2D images. It is not clear that the ease of this approach justifies giving up the advantages of a full-fledged re-analysis of the imaging in D_A and D_B .

A second problem compounds the issue. Because our starting point is the intrinsic profile, Σ_p , the pixel scale “resolution” of the mock image has to be high enough to accommodate the steeply rising interiors of profiles with small R_e , especially deV profiles or those with high Sérsic- n values. In our brute force approach, we found the needed resolution simply by decreasing the pixel scale of the mock images until there were no differences in the SYNMAgs measured at different resolutions. In tests of massive galaxy profiles at $z \sim 0.5$ (see Section 3.1) we found convergence at a mock image resolution of roughly $0.01 \text{ arcsec pixel}^{-1}$. This resolution results in prohibitively large PSF convolution times that were the whole motivation of synthetic photometry. A more clever approach with an adaptive resolution grid would address this problem, but given the other limitations, we adopt a completely different and more elegant approach in the following section.

2.4. Multi-Gaussian Expansion Implementation

Synthetic aperture photometry becomes attractive compared to a re-analysis of imaging pixel data once it is significantly faster. Because the slowest step in methods that work with pixel data is often PSF convolution, we develop a method here which decomposes intrinsic profiles and PSFs into MGEs. The fact that convolutions between Gaussians can be calculated analytically

makes synthetic photometry with MGEs extremely fast, as long as the MGE decomposition of the profile and PSF need only be performed once for the full data set. We will show this is commonly the case.

The power of MGEs in the context of smoothly decreasing profiles common in astronomy was recognized early on (e.g., Trumpler & Weaver 1953) and applied rigorously to flux profiles such as the King model and Moffat PSF by Bendinelli (1991). Their utility and a modern implementation was presented in Emsellem et al. (1994) and also discussed in Connolly et al. (2000).

2.4.1. MGE Decomposition

Our first step is to describe the intrinsic profiles, Σ_p , that are fit to the profile-band data in the data set, D_A , by an MGE, that is, a sum of j Gaussian profiles:

$$\Sigma_p(r) = \sum_j \Sigma_{G,j}(r) = \sum_j A_j \exp\left(-\frac{r^2}{2\sigma_j^2}\right), \quad (5)$$

where A_j is the normalization (central flux density) and the width of each Gaussian is given by σ_j . The summed flux of each Gaussian is $L_j = 2\pi A_j \sigma_j^2$. As we describe below, it is possible to determine an MGE for a profile of arbitrary size and shape, a process that takes less than a minute. However, as long as the chosen intrinsic form of the profile scales with an effective radius (R_e) and normalization (Σ_e), the σ_j and A_j of a single MGE determined for that profile form can also be scaled in the same way. For a defined profile shape, this means a single MGE can be scaled to match profiles with arbitrary size and flux. Determining the MGE for the intrinsic profile of every source in D_A is not required, although a separate MGE is needed for each profile *shape* that was fit for. In SDSS, this means an MGE must be determined for both the Exp and deV profiles. For Sérsic profiles generally, an MGE is required for each Sérsic- n value. From the point of view of SYNMAgs, it is helpful if the number of profile shapes fitted for in D_A (i.e., the number Sérsic- n values fitted for) is restricted to a modest number. The MGE descriptions of these shapes can be determined once and then used for any subsequent measurement of SYNMAgs involving the data sets in question.

So far we have considered axisymmetric profiles defined in one dimension. Two-dimensional profiles are easily described by a sum of 2D Gaussians. We let the profile’s major axis be the x -axis and define the axis ratio, $q = b/a$, where a and b are characteristic sizes along the major and minor axes, respectively. The MGE becomes

$$\Sigma_p(x, y) = \sum_j A_j \exp\left(-\frac{x^2}{2\sigma_{x,j}^2} - \frac{y^2}{2q^2\sigma_{x,j}^2}\right), \quad (6)$$

after substituting $\sigma_y = q\sigma_x$. For a given profile shape, any 2D representation of arbitrary size, flux, or axis ratio can thus be described by a single 1D MGE.

The MGE must now be carefully determined with an eye toward minimizing the number of Gaussian terms while preserving precision in the MGE approximation of the intrinsic profile. We find that a least-squares fit to the full (circular) 2D profile provides the best approximation to the true surface brightness profile because the fit is then weighted by the intensity. MGE approximations for standard Exp and deV profiles as well as the modified profiles used in SDSS are given in Table 1 and their 1D

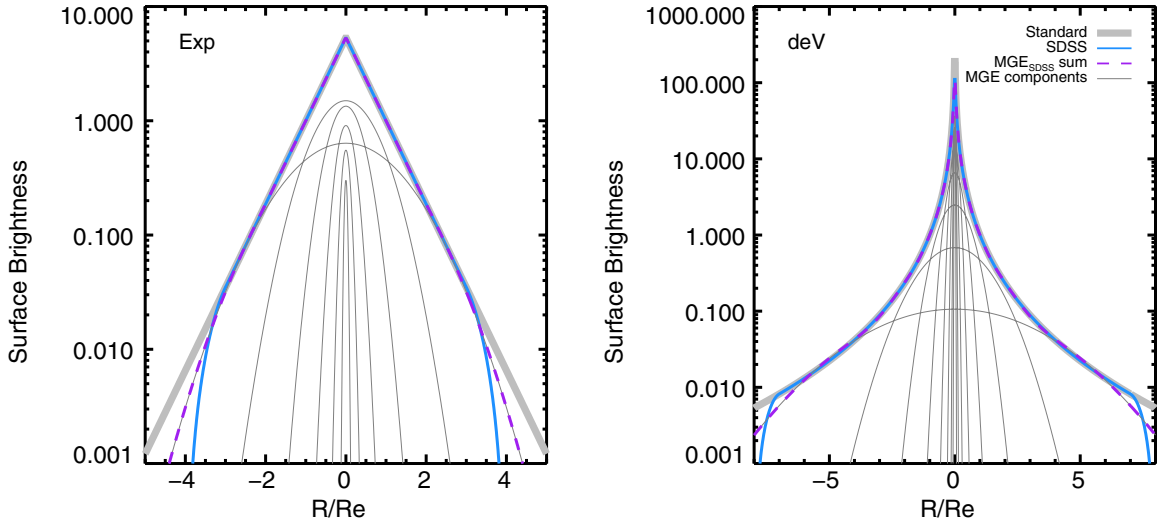


Figure 1. Sample Exponential (left) and de Vaucouleurs (right) profile, illustrating the MGE decomposition. The standard forms are shown as solid purple lines, while the SDSS-modified forms are shown in dark gray. Surface brightness units are arbitrary and scaled so that $\Sigma(R_e) \equiv \Sigma_e = 1.0$. The MGE components for the modified profiles are overplotted as thin, light gray lines. They sum to give the dashed blue curves.

(A color version of this figure is available in the online journal.)

Table 1
MGE Parameters

N_{terms}	Σ_{Exp}		Σ_{deV}		$\Sigma_{\text{Exp}}^{\text{SDSS}}$		$\Sigma_{\text{deV}}^{\text{SDSS}}$	
	σ_j	A_j	σ_j	A_j	σ_j	A_j	σ_j	A_j
1	5.1	0.45	0.1	328.94	3.5	0.31	1.5	30.33
2	13.8	0.80	0.5	176.20	9.4	0.55	3.2	38.13
3	28.8	1.22	1.5	99.93	19.8	0.91	6.5	26.07
4	53.2	1.46	3.9	52.88	37.4	1.34	13.0	14.27
5	91.2	1.04	9.9	24.55	67.9	1.50	26.2	6.59
6	150.2	0.22	24.7	9.29	122.5	0.64	53.6	2.47
7	63.9	2.51	115.5	0.68
8	192.6	0.35	289.9	0.11

Notes. The components are normalized so that the sum yields $\Sigma(R_e) \equiv \Sigma_e = 1.0$ in arbitrary units of surface flux density. The units of σ_j are $0.01R_e$.

radial profiles are shown in Figure 1. In terms of enclosed flux, the number of terms used in these MGEs provide accuracies at the 0.01 mag level out to $10R_e$ and so we deem them adequate approximations for our purposes. More details on the optimal determination of these fits as well as more accurate representations employing additional Gaussian terms are provided in Hogg & Lang (2012).

It is generally difficult to fit the imposed truncation in the modified SDSS profiles with Gaussians, although the central smoothing in the modified de Vaucouleurs profiles makes an MGE solution easier in the central regions. The MGEs begin to deviate from the true SDSS profiles at radii larger than roughly $3.5R_e$ for the Exp profile and $6R_e$ for the Dev profile (see Figure 1). This introduces aperture flux discrepancies of roughly 0.01–0.02 mag, and so we consider these MGEs adequate for our purposes.

2.4.2. PSF Convolution with MGEs

The *observed* galaxy profile, I_p is the convolution of Σ_p with the PSF profile, R_{PSF} . The Fourier transform, \mathcal{F} , of this convolution can be expressed by the product of Fourier transforms:

$$\mathcal{F}(\Sigma_p \otimes R_{\text{PSF}}) = \mathcal{F}(\Sigma_p)\mathcal{F}(R_{\text{PSF}}). \quad (7)$$

The power of the MGE approach lies in the fact that the Fourier transform of a Gaussian is also a Gaussian and easily determined analytically. Fourier transforms are also linear, so that for each Gaussian term, j ,

$$\mathcal{F}(\Sigma_p(r)) = \sum_j \int A_j \exp\left(-\frac{r^2}{2\sigma_j^2} - 2\pi i \vec{k}_r \cdot \vec{r}\right) dr. \quad (8)$$

Referring to Equation (6), $\mathcal{F}(\Sigma_p(x, y))$ can be expressed as

$$\mathcal{F}(\Sigma_p(x, y)) = 2\pi \sum_j A_j \sigma_j^2 q \exp\left(-\frac{2\pi^2 \sigma_{x,j}^2}{2} (k_x^2 + q^2 k_y^2)\right). \quad (9)$$

For the derivations here, we assume that R_{PSF} can also be decomposed into a series of circular Gaussian terms, k , with coefficients R_k and standard deviations, $\sigma_{R,k}$. A two-component Gaussian PSF, often used to account for power in the wings of the PSF, is an obvious example. Taking the inverse Fourier transform of Equation (7), we can express the PSF-convolved “observed” profile, I_p , as

$$I_p(x, y) = \sum_j \sum_k C_{j,k} \exp\left(-\frac{x^2}{2(\sigma_{x,j}^2 + \sigma_{R,k}^2)} - \frac{y^2}{2(q^2 \sigma_{x,j}^2 + \sigma_{R,k}^2)}\right), \quad (10)$$

where the coefficients $C_{j,k}$ are given by

$$C_{j,k} = \frac{2\pi A_j R_k q \sigma_{x,j}^2 \sigma_{R,k}^2}{\sqrt{(\sigma_{x,j}^2 + \sigma_{R,k}^2)(q^2 \sigma_{x,j}^2 + \sigma_{R,k}^2)}}. \quad (11)$$

While we have assumed circular PSF components to this point, the SYNMAJ approach and associated software handles non-circular, spatially offset, and even negative components, providing a flexible way to describe more complicated PSF shapes via R_{PSF} .

2.4.3. MGE Aperture Flux and Profile

Equation (10) provides an analytic approximation, I_p , of a galaxy light profile with an intrinsic shape measured in the p band, but observed with the PSF of another band. In our case, the second band comes from a different data set, D_B for which only aperture photometry is available. To compare the photometry recorded in D_B to the aperture flux, F_p , which would have been measured in the p band with the same PSF, we must integrate Equation (10) over the aperture(s) used in D_B . Because I_p is generally elliptical while the apertures are circular, a numerical integration is required. Working in Cartesian coordinates, the 2D integral of Equation (10) over a circular aperture of radius R can be expressed as

$$F_p(R) = \sum_j \sum_k \int_0^R \int_0^{c(x)} 4C_{j,k} \exp\left(-\frac{x^2}{2\hat{\sigma}_x^2}\right) \exp\left(-\frac{y^2}{2\hat{\sigma}_y^2}\right) dx dy, \quad (12)$$

where $c(x) = \sqrt{R^2 - x^2}$, and $\hat{\sigma}$ represents the effective standard deviation of a PSF-convolved MGE term, such that $\hat{\sigma}_x^2 = (\sigma_{x,j}^2 + \sigma_{R,k}^2)$, and $\hat{\sigma}_y^2 = (q^2\sigma_{x,j}^2 + \sigma_{R,k}^2)$. The second integral can be expressed in terms of the error function, so that a fast 1D implementation of the numerical integral can be written as

$$F_p(R) = 4C_{j,k}\hat{\sigma}_y \sqrt{\frac{\pi}{2}} \int_0^R \exp\left(-\frac{x^2}{2\hat{\sigma}_x^2}\right) \operatorname{erf}\left(\sqrt{\frac{R^2 - x^2}{2\hat{\sigma}_y^2}}\right) dx. \quad (13)$$

We can also use the MGE description of I_p to estimate the effective radius and axis ratio of the *observed* flux profile. Working along the major axis—or equivalently assuming the axis ratio is 1.0—we can define an effective half-light radius, $R_{e,\text{maj}}$, analogous to the R_e defined for the intrinsic profile, Σ_p . The integral of the 2D flux is easily calculated:

$$L(x) = \sum_{j,k} 2\pi C_{j,k} \hat{\sigma}_x^2 \left[1 - \exp\left(-\frac{x^2}{2\hat{\sigma}_x^2}\right) \right]. \quad (14)$$

Setting this equal to the (major-axis) “half-flux” given by $L_{1/2,\text{maj}} = \sum \pi C_{j,k} \hat{\sigma}_x^2$ and solving for x , one obtains $R_{e,\text{maj}}$. In practice, this definition of half-light radius is not as intuitive as in the case of Σ_p because the axis ratio depends on radius (as we show below). Thus, the minor-axis half-light radius is not simply related to $R_{e,\text{maj}}$. We can alternatively define a second measure of the half-light radius, $R_{e,\text{circ}}$, that corresponds to the size of a circular aperture which contains half of the total flux. Here, we integrate Equation (13) until the enclosed 2D flux is equal to the true half-flux as given by $L_{1/2} = \sum \pi C_{j,k} \hat{\sigma}_x \hat{\sigma}_y$.

Turning to the effective axis ratio, while the intrinsic axis ratio is fixed, the observed value will be a function of radius. Seeing will round out profiles at a scale corresponding to the PSF FWHM, but will have a smaller effect on elongated (or inclined) profiles at larger radii. The *convolved* axis ratio of each Gaussian term is given by

$$\hat{q}_{j,k} = \frac{\hat{\sigma}_y^2}{\hat{\sigma}_x^2} = \frac{q^2 \left(\frac{\sigma_{x,j}^2}{\sigma_{R,k}^2} \right) + 1}{\left(\frac{\sigma_{x,j}^2}{\sigma_{R,k}^2} \right) + 1}. \quad (15)$$

To derive an estimate of the observed profile’s axis ratio, $\langle q \rangle$, we can weight each term by approximately the half-flux:

$$\langle q \rangle = \frac{\sum w_{j,k} \hat{q}_{j,k}}{\sum w_{j,k}}, \quad (16)$$

where the weights are given by $w_{j,k} = C_{j,k} \hat{\sigma}_x \hat{\sigma}_y$.

3. TESTS WITH GALAXY PHOTOMETRY

The motivation for SYNMAgS is fast matched photometry for overlapping data sets that span large areas of the sky and thus make a re-analysis of pixel data difficult and time consuming. Since we cannot use the stellar locus—a common way of testing photometry methods—we perform tests of the SYNMAgS approach using galaxy magnitudes and colors. We first show that synthetic magnitudes are consistent with SDSS FiberMags. We then compare galaxy colors from synthetic photometry to “FWHM-homogenized” matched-aperture colors from the GAMA Survey. We consider the derived scatter in the color of the red sequence and in the photometric redshifts determined from these two sets of photometry, finding that SYNMAgS perform as well or better than the standard but more laborious FWHM-homogenization method. The latter two tests require redshift information which we take from the BOSS.

3.1. SDSS FiberMags

In all five bands, and for every detected source, the SDSS PHOTO pipeline provides a “FiberMag” estimate of the 3” diameter aperture magnitude, assuming seeing of 2” FWHM. This is meant to approximate the flux measured by a spectroscopic fiber in SDSS-I/II. Since SYNMAgS can provide synthetic aperture photometry for an aperture of arbitrary size and with arbitrary seeing, a natural first test is to compare SYNMAgS approximations of FiberMags to their actual values.

With an eye toward future applications of SYNMAgS photometry for BOSS galaxies, we perform this exercise using a sub-region of the SDSS Stripe 82 Coadd catalog (Annis et al. 2011) with $355^\circ < \text{R.A.} < 360^\circ$. We select extended sources with non-stellar “psfmag” colors and detections in the UKIDSS LAS catalog (DR4). This effectively selects galaxies with $M_* > 10^{11} M_\odot$ and $0.2 < z < 0.7$. We apply cuts on both SDSS and UKIDSS flags to ensure a clean sample and remove blended objects by excluding SDSS sources with the BLEND flag set. The exact nature of the flag cuts does not influence our results.

We have found that a precise understanding of the PSF is important for resolving discrepancies at the 0.05 mag level. In the SDSS imaging, the PSF can be approximated by a two-component Gaussian. The second component has a width that is twice that of the first component, but accounts for only 10% of the PSF power. In the PHOTO pipeline (see `measureObj.c`), the FiberMag is determined by convolving a sub-image by a single Gaussian kernel with $\sigma_k^2 = \sigma_{2.0}^2 - \sigma_{\text{mom}}^2$, where $\sigma_{2.0}$ is the value of σ for a Gaussian with FWHM = 2.0 (i.e., $2.0/2.3548$) and σ_{mom} is the adaptive second moment size of the local PSF (referred to as `M_rr_cc`). Thus, the FiberMags are *not* measured on images with an FWHM 2.0 Gaussian PSF but with a more complicated PSF profile with significant power in the wings.

For the consistency check here, we do not attempt to recover the applied kernel or the original (and likely spatially varying) PSF of the underlying image. Instead, we approximate the final, effective PSF for FiberMags as a two-component Gaussian

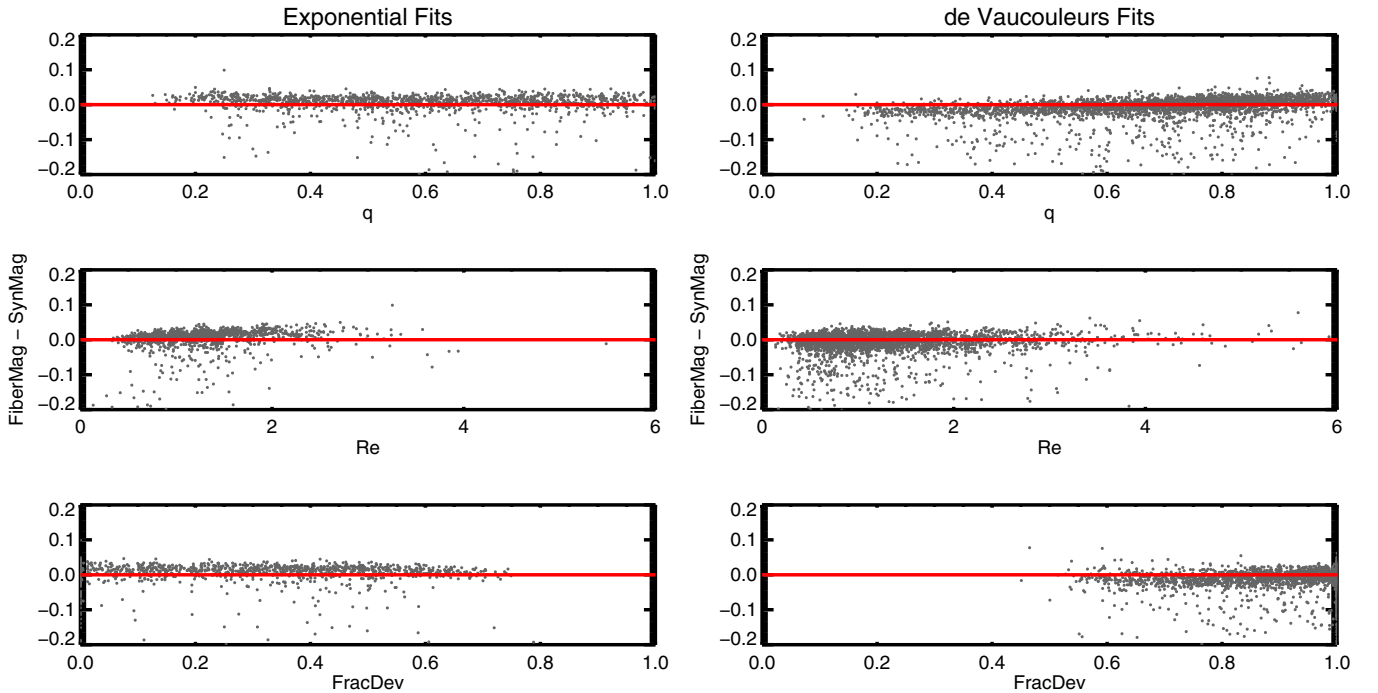


Figure 2. Comparison between r -band FiberMags and SYNMAgS. Sources best fit by exponential profiles are compared in the left-hand column, while de Vaucouleurs fits are shown in the right-hand column. From top to bottom, the panels investigate magnitude differences as a function of intrinsic axis ratio (q), effective radius (R_e), and FracDev, an estimator for how well the profile shape can be fit by a de Vaucouleurs profile. Residuals at the 0.01–0.02 mag level are probably caused by an imperfect treatment of the PSF.

(A color version of this figure is available in the online journal.)

(as described above), with the FWHM of the first term set to $2''.05$. Experimentation revealed that our uncertainty in this choice introduces systematic offsets at the 0.01–0.02 mag level. We then divide the sample into sources that are best fit by exponential profiles versus de Vaucouleurs profiles (this is easily determined by whether the ModelMag matches the ExpMag or deVMag) to see how the SYNMAg MGEs and PSF convolutions perform in either case.

The results are shown in Figure 2 which compares r -band magnitude differences between SYNMAgS and FiberMags as a function of axis ratio, size, and FracDev, a quantity that expresses how well a source can be described by a de Vaucouleurs profile. The SYNMAgS perform well with small systematic residuals at the 0.01–0.02 mag level that probably result from our imperfect treatment of the PSF. We expect additional systematics from the discrepancy between the MGE profile approximations and the truncated SDSS profiles and from the fact that a single-component profile is not a good description of most galaxies. For our purposes, these systematics are less than the statistical uncertainties in most applications of galaxy photometry as well as the photometric calibration uncertainties inherent in comparisons of SDSS to other data sets. Finally, we note a small degree of scatter toward brighter FiberMags. A large fraction ($\sim 80\%$) of these objects were removed by restricting to non-blended sources. We therefore surmise that the remainder are objects with nearby neighbors that contribute flux to the FiberMag aperture or galaxies with bright sub-components that are not accounted for in the profile fitting.

3.2. Red Sequence

Having verified the ability of SYNMAgS to reproduce aperture photometry measured in the SDSS r band, we now investigate how well SYNMAgS provide matched photometry across

differing data sets. Our focus will be on regions of overlap between SDSS imaging and the UKIDSS/LAS. These data sets have been matched previously using standard FWHM homogenization with re-measured aperture photometry in two recent papers. Matsuoka & Kawara (2010) carry out this exercise for 40 deg^2 in Stripe 82. The GAMA survey (Driver et al. 2011) does so in three equatorial fields, totaling 144 deg^2 (Hill et al. 2011). Because the GAMA DR1 catalogs are larger and publicly available,¹⁵ we will use them to test the synthetic photometry method developed here.

As described in Driver et al. (2011) and further detailed by Hill et al. (2011), the GAMA DR1 photometric catalog was constructed by first downloading the reduced images in all nine available bands— $ugriz$ from SDSS, and $YJHK$ from LAS—and convolving all imaging to a common seeing FWHM of $2''.0$. We refer to this as “FWHM homogenization,” and distinguish it from “PSF homogenization” in which the FWHM and detailed shape of the PSF are made identical in every band (see Hildebrandt et al. 2012). All images are registered to a common astrometry, and source detection is performed in the r band, which defines the Kron aperture that is applied to the flux measurements in the other bands. In the G09 and G15 fields, the catalog is limited to a Petrosian magnitude of $r_{\text{pet}} < 19.4$, while in G12 the limit is $r_{\text{pet}} < 19.8$. Further details and various tests of the performance of GAMA photometry are provided in Hill et al. (2011).

Our first test measures the scatter in the rest-frame ($r - J$) colors of the so-called CMASS galaxies from the BOSS (see Eisenstein et al. 2011), most of which lie on the red sequence. The BOSS target selection is briefly described in Anderson et al. (2012) and will be further detailed in N. Padmanabhan et al.

¹⁵ <http://www.gama-survey.org/drl/YR1public.php>

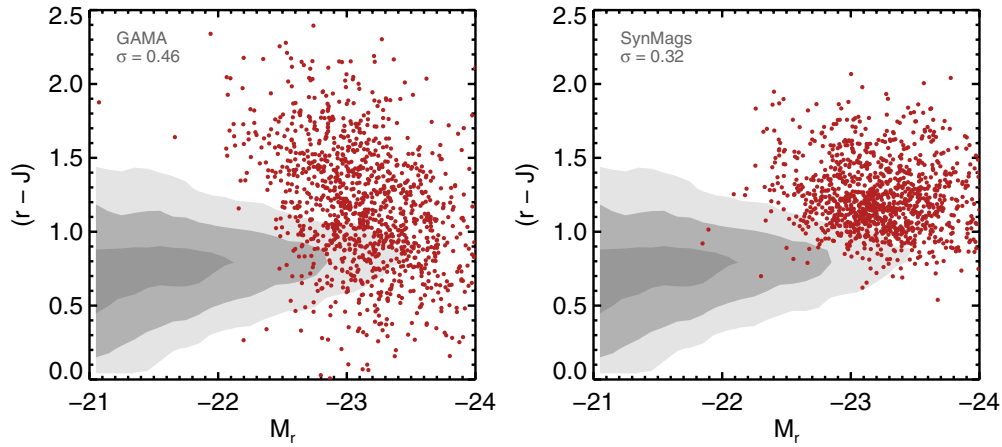


Figure 3. Rest-frame $(r - J)$ color-magnitude diagrams of (mostly) red-sequence CMASS galaxies from the BOSS survey. The observed scatter for the GAMA FWHM-homogenized matched-aperture photometry is larger than for the SYNMAg approach described here. Contours indicate the location of much less luminous COSMOS galaxies in the same redshift range.

(A color version of this figure is available in the online journal.)

(in preparation). The CMASS selection is designed to target very luminous galaxies with $z > 0.4$ and $M_* \gtrsim 10^{11.4} M_\odot$. Most have red colors, although 20%–30% show evidence for disk or irregular components, and 15% have somewhat bluer colors than a standard red-blue color cut (Masters et al. 2011). We select 1203 CMASS galaxies in the three GAMA regions with secure spectroscopic redshifts. We apply K -corrections to infer rest-frame absolute magnitudes using K -correct (Blanton & Roweis 2007) and apply no evolutionary corrections.

SYNMAGs are derived using the method described in previous sections from the SDSS DR7 (Abazajian et al. 2009) and UKIDSS/LAS publicly available photometric catalogs.¹⁶ Detections are required in all bands. We chose to synthesize 2''/8 radius UKIDSS aperture photometry. While we primarily use UKIDSS DR8 catalogs, the GAMA photometry was based on DR4, so we note that we find no differences when using UKIDSS DR4 catalogs instead. The SYNMAgS model the varying PSF in each of the $YJHK$ bands as a Gaussian with the local FWHM taken from the LAS catalogs (note that a more sophisticated PSF treatment is possible). The typical UKIDSS FWHM is 0''.8, which is smaller than the 1''.1 seeing SDSS imaging that defines the intrinsic profiles needed for SYNMAg photometry. Extinction corrections from Schlegel et al. (1998) are applied using values recorded in both the SDSS and UKIDSS catalogs. Note that the Kron-aperture-matched photometry from GAMA DR1 is also extinction corrected.

The rest-frame $(r - J)$ color-magnitude diagrams for both the GAMA photometry (left panel) and SYNMAg photometry (right panel) are shown in Figure 3. The choice of $(r - J)$ provides a test of the photometric matching between the SDSS and UKIDSS data sets. To provide a reference, the contours indicate the location of much less luminous COSMOS galaxies¹⁷ with similar redshifts as the CMASS sample. We note an offset that may arise from the combination of zero-point calibration uncertainties in the COSMOS data set (see Capak et al. 2007) and the fact that the very massive BOSS galaxies are not well represented in the small volume of the COSMOS

field. Because photometric scatter is likely to dominate over intrinsic scatter, the observed scatter in these diagrams gives some indication of the performance of the photometry matching. The SYNMAg photometry appears to do very well, providing a smaller scatter ($\sigma = 0.32$) compared to the GAMA photometry ($\sigma = 0.46$). We note, however, that GAMA DR1 photometry may have been affected by a bug in the PSF convolution (S. Driver 2012, private communication). Our own tests suggest its impact on the GAMA colors may not be significant, but future releases of GAMA photometry may show improvement.

3.3. Photometric Redshifts

The moderate scatter in the SYNMAg SDSS–UKIDSS colors of galaxies on the red sequence suggests that synthetic aperture photometry performs well, but the scatter alone does not tell us whether the colors are “correct.” We therefore perform a second test by comparing the performance of photometric redshifts derived using both the GAMA and SYNMAg approach. Here the goal of accurate photo- z s is more clearly defined, although optimizing photometry based on the performance of photo- z s carries implicit assumptions that should be carefully considered (Section 4).

We again turn to BOSS galaxies with spectroscopic redshifts as a benchmark for evaluating the photo- z performance. We utilize all BOSS targets in this case (including the LOWZ sample) which allows us to probe redshifts from 0.1 to 0.8, thereby testing the ability of multiple SDSS–UKIDSS filter combinations to correctly measure rest-frame spectral features. We apply the v1.99.3 Bayesian photometric redshift (BPZ) code (Benítez 2000; Coe et al. 2006) to both the GAMA matched photometry and the SYNMAg photometry. BPZ includes a procedure for estimating photometric zero-point offsets from the mismatch between the best-fit templates and the observed photometry of galaxies where the redshift is known spectroscopically. The resulting suggested offsets were ~ 0.02 mag in the u through J bands, but 0.1 mag in the H band, and 0.18 mag in the K band, reflecting the poorer quality of the BPZ templates in the near-IR. We apply these offsets to both sets of photometry, although their impact is small.

The comparison of derived photometric redshifts compared to the 2989 BOSS-measured spectroscopic redshifts is shown in Figure 4. We emphasize that the goal is not to obtain the best

¹⁶ Special care is required in querying the UKIDSS database and removing corrections applied to the aperture photometry in the WFCAM Science Archive (WSA). These will be described in detail in K. Bundy et al. (in preparation).

¹⁷ The COSMOS data set is described in Bundy et al. (2010), with rest-frame colors calculated by Ilbert et al. (2010).

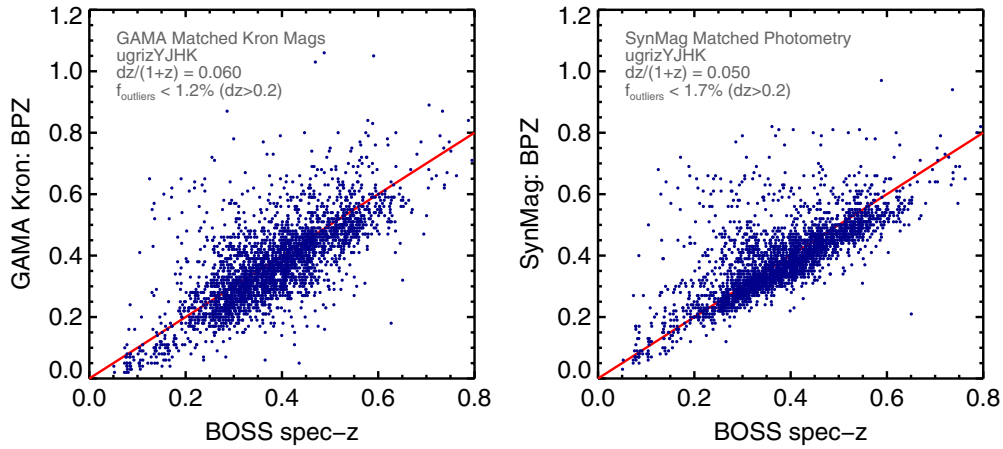


Figure 4. For the same sample of BOSS galaxies, the quality of photometric redshifts derived from FWHM-homogenized *ugrizYJHK* GAMA photometry (left panel) compared to the synthetic aperture photometry method using the same filter set (right panel). The SYN MAG approach performs better.

(A color version of this figure is available in the online journal.)

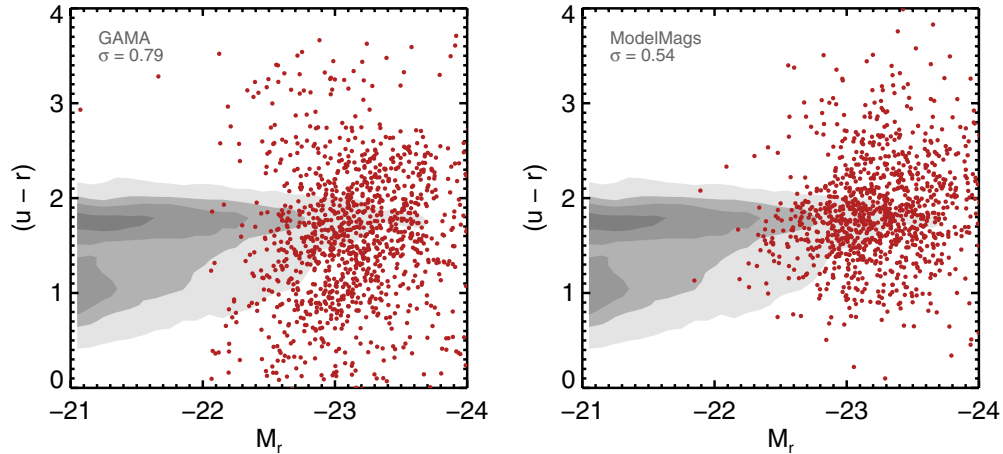


Figure 5. Rest-frame $(u-r)$ color-magnitude diagrams of (mostly) red-sequence CMASS galaxies from the BOSS survey. As with Figure 3, the observed scatter for the GAMA FWHM-homogenized matched-aperture photometry is larger than for SDSS ModelMags, demonstrating the utility of profile fitting in this context. Contours indicate the location of much less luminous COSMOS galaxies in the same redshift range.

(A color version of this figure is available in the online journal.)

photo- z s possible for this sample, but to compare the resulting photo- z quality from two different photometry techniques. When excluding outliers defined as $|z_{\text{spec}} - z_{\text{phot}}| > 0.2$, we find that the FWHM-homogenization approach used by GAMA achieves a $dz/(1+z)$ of 0.06 while the SYN MAGs deliver an improved scatter of 0.05. Indeed, the decreased scatter in the SYN MAG photo- z s is evident by eye in Figure 4. The SYN MAGs outlier fraction is slightly larger, however (1.7% instead of 1.2%).

This result was anticipated by Figure 3; the tighter colors do appear to yield better photo- z s. This is highly encouraging for the utility of the SYN MAG approach. Even if future improvements to the GAMA photometry deliver better BPZ performance, the SYN MAGs require far less work and computation time. The photometry computation for the comparison set of ~ 3000 galaxies here takes less than one minute on a modern desktop (8-core CPUs with 10 GB memory). Combined with the red-sequence test, these results indicate that the scatter in SYN MAG colors may outperform PSF homogenization. The photo- z test also tells us that SYN MAGs have not introduced strong systematic offsets (greater than 0.05 mag) between filters because these would decrease the photo- z precision. The more

general question of whether SYN MAGs or any matching technique introduces biases in the observed colors requires defining what is meant by “correct” colors, against which biases could be measured. As we discuss in Section 4, no single definition is appropriate and this argues for adopting different approaches tied to specific scientific goals.

3.4. Why SYN MAGs Perform Well

A comparison of the red-sequence colors (Figure 3) and photometric redshifts (Figure 4) suggests that sets of photometry constructed with SYN MAGs perform better than FWHM-homogenization matched-aperture photometry. Putting aside the possibility that the GAMA matched-aperture photometry may improve in future releases, it is worth asking why SYN MAGs may perform better in these tests. Two explanations are apparent. To begin with, SYN MAGs do not require a degradation of the image quality to the lowest-common-denominator PSF. For many data sets, this is typically $\text{FWHM} = 2''$, which is on the same order as both the size of distant galaxies and the aperture size over which their flux is measured. The benefit of not degrading the PSF is demonstrated by the improved scatter of the SYN MAG $(r-J)$ colors.

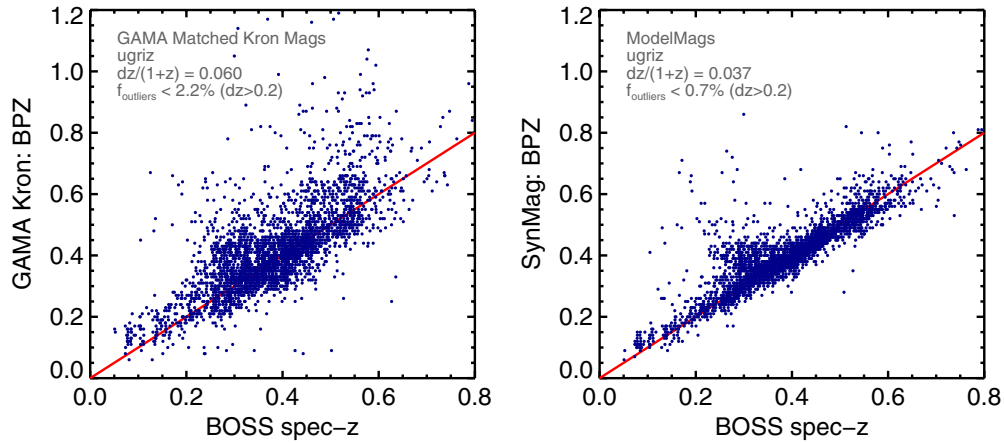


Figure 6. Same photometric redshift comparison as in Figure 4, but restricting only to SDSS *ugriz* photometry. SYNMAg colors in this case are equivalent to SDSS ModelMags.

(A color version of this figure is available in the online journal.)

The second explanation is that we have constructed a set of photometry that takes advantage of profile-fitting photometry when available. The photo-*z*s presented above are based on a mix of SDSS ModelMag colors from profile fitting in the optical and SYNMAg colors defined by aperture photometry from UKIDSS in the near-IR. We can examine the value of the profile-fitting colors alone by ignoring the near-IR SYNMAgs for the moment and comparing colors and photo-*z*s between GAMA and SDSS in the optical only. Figure 5 presents the rest-frame (*u* − *r*) color–magnitude diagrams derived for CMASS galaxies from GAMA (left panel) compared to the SDSS profile-fitting ModelMags (right panel), and Figure 6 shows the same test with *ugriz* photo-*z*s. Both show significantly reduced scatter from the use of ModelMags.¹⁸ The SYNMAgs themselves cannot claim credit for the reduced scatter in this case, but it is clearly important to make use of profile-fitting photometry if it is available. It is likely, for example, that a mix of ModelMags in the optical and FWHM-homogenized colors in the near-IR would yield tighter colors and photo-*z*s than FWHM-homogenized colors in all bands.

4. CONCLUSIONS

We have presented a technique for using catalog-level data products to match the photometry of disparate imaging data sets. Our SYNMAg approach uses MGEs to represent both intrinsic profiles and the PSF, allowing for extremely fast convolutions. The result is a tool that can be used to quickly obtain matched photometry for large data sets without re-analyzing pixel data.

We find that SYNMAgs yield colors and derived photo-*z*s with scatter that is as good if not better than FWHM-homogenized aperture photometry. We have also shown that the resulting scatter is minimized by constructing photometric data sets that employ profile-fitting colors when available. These tests, however, *assume* that the “best” color estimators minimize the scatter. This may not be the case, especially given that most galaxies are two-component bulge + disk systems, often with different colors for each component. In fact, the choice of aperture size or assumed profile shape should depend on

the goal for which the photometry is being analyzed. Such “application-oriented photometry” would suggest, for example, that for photo-*z*s, fitting profiles that emphasizes the typically redder light from the bulge would strengthen the 4000 Å break in the derived SED and yield tighter photo-*z* estimates. The same profile shape would not be optimal for studying the SEDs of galaxy disks, however.

Applying SYNMAgs to future survey data would benefit from having some specific measurements recorded in publicly available catalogs. Even if profile-fitting photometry is performed, this basic information is valuable for cross-checks, troubleshooting, and the application of SYNMAgs to future overlapping data. First, it would be useful to have aperture photometry performed in a series of circular apertures that range from roughly one-half of the typical PSF FWHM to *R* = 6″–7″ (integer values like radii of 2″, 3″, 4″, etc., are common). The *R* = 6″ corresponds to ∼11 kpc/*h*₇₀ at *z* = 0.1, sufficient for galaxy photometry at *z* > 0.1. Several larger apertures up to *R* = 20″ would provide adequate sampling for most of the largest galaxies in SDSS (*z* < 0.1). No corrections or adjustments (e.g., PSF corrections) should be applied to the aperture fluxes beyond background subtraction. Reliable uncertainties on each flux measurement are also valuable. Special care is needed for smaller apertures that subtend only a few pixels, otherwise it may be better to remove small aperture measurements from the catalogs. It is also useful to consider the definition of null values for measured magnitudes. The user should be able to distinguish between the case when no flux was detected (e.g., *m* = 99) and when imaging for a source in a particular band is missing (e.g., *m* = −99), either because it was not observed or was somehow compromised.

SYNMAgs also require information about the PSF. Ideally, each catalog entry would provide a way to determine the PSF shape at the location of the measured object given the conditions under which it was observed. Often a two-component Gaussian description is sufficient although the modeled shape could be more sophisticated and could be modeled as a function of both detector position and observing conditions or exposure number. The position angle of off-axis or non-axisymmetric components also needs to be recorded.

We thank Masahiro Takada, Sogo Mineo, Chiake Higake, and Daniel Thomas for useful discussions that provided new

¹⁸ It is perhaps surprising that the loss of near-IR photometry does not degrade the photo-*z*s. In fact, for the ModelMags, the photo-*z*s significantly *improve* without the near-IR. This is in large part a result of uncertainties in rest-frame near-IR templates, a widespread problem that has motivated template error weighting in recent photo-*z* codes (e.g., Brammer et al. 2008).

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APPENDIX

COOKBOOK: USING THE SYN MAG SOFTWARE

This Appendix presents instructions for using the SYN MAG software.¹⁹ We describe the key steps in the process and point to the routine `synmag_sdss` for the specific case of matching to SDSS data.

Before using the SYN MAG routines, some photometric information from the data sets involved is required. This should include parameters from profile fitting in at least one band, aperture photometry results for the “target” data set where profile fits are not available, and seeing estimates for this data set. With these information in hand, the procedure is as follows.

1. *Decompose the intrinsic profile shape into an MGE.* This only needs to be done once for each profile shape that is used. For standard de Vaucouleurs and exponential profiles, or their SDSS variants, the Gaussian coefficients can be read off of Table 1. Assuming more generally that the profile is a Sérsic, the `mgser` routine can be used to determine the MGE. It utilizes a 1D fitting routine in the `MGE_FIT_SECTORS` IDL package presented in Cappellari (2002). As discussed in Section 2.4.1, however, a least-squares MGE fit to the 2D profile provides a better approximation (see Hogg & Lang 2012). The PSF of the

target catalog can also be decomposed in this way if desired.

2. *Compute synthetic photometry.* Use `synmag` to compute the synthetic aperture photometry in the profile-defining band for the PSF and aperture size measured in the target catalog. Multiple apertures can be requested and the PSF parameters can be specified separately for each source to account for variations across the image or from differing observing epochs.
3. *Assemble a matched set of photometry.* Repeat the `synmag` procedure for each filter band in the target catalog to account for PSF variations and missing data from one band to the next. Then subtract the resulting colors from the total flux measured in the profile-defining band in order to construct a set of matched flux total flux measurements (i.e., $m_{i,\text{tot}} = m_{p,\text{tot}} - C_{pi}$, see Section 2.1).

The PSF-convolved MGEs can also be returned by `synmag`. The routine `mgstr` computes the effective radius and axis ratio of any MGE (see Section 2.4.3).

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¹⁹ <http://member.ipmu.jp/kevin.bundy/synmag>