

DISCOVERY OF THE ACCRETION-POWERED MILLISECOND PULSAR SWIFT J1756.9–2508 WITH A LOW-MASS COMPANION

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ABSTRACT

We report on the discovery by the *Swift Gamma-Ray Burst Explorer* of the eighth known transient accretion-powered millisecond pulsar, SWIFT J1756.9–2508, as part of routine observations with the *Swift* Burst Alert Telescope hard X-ray transient monitor. The pulsar was subsequently observed by both the X-Ray Telescope on *Swift* and the *Rossi X-Ray Timing Explorer* Proportional Counter Array. It has a spin frequency of 182 Hz (5.5 ms) and an orbital period of 54.7 minutes. The minimum companion mass is between 0.0067 and 0.0086 M_{\odot} , depending on the mass of the neutron star, and the upper limit on the mass is 0.030 M_{\odot} (95% confidence level). Such a low mass is inconsistent with brown dwarf models, and comparison with white dwarf models suggests that the companion is a He-dominated donor whose thermal cooling has been at least modestly slowed by irradiation from the accretion flux. No X-ray bursts, dips, eclipses, or quasi-periodic oscillations were detected. The current outburst lasted ≈ 13 days, and no earlier outbursts were found in archival data.

Subject headings: binaries: close — pulsars: general — stars: neutron — white dwarfs — X-rays: binaries

1. INTRODUCTION

It is now well established that a subset of neutron stars in low-mass X-ray binary systems are millisecond pulsars powered by accretion from either a brown dwarf or degenerate companion in a very close binary configuration. Such systems are believed to be progenitors of millisecond radio pulsars (Alpar et al. 1982). Before the discovery of SWIFT J1756.9–2508, there were seven known accretion-powered millisecond pulsars (see Kaaret et al. 2006 for a full listing). The sample is now becoming large enough to start to study classes of these objects, which can in turn lead to a better understanding of the properties of neutron star X-ray binaries and their relationship to millisecond radio pulsars. The known accretion-powered millisecond pulsars all appear to be in very close orbits (with orbital periods ranging between 41.1 and 256.5 minutes). For all but one such system, XTE J1814–338 (Bhattacharyya et al. 2005), orbital constraints rule out a hydrogen-burning main-sequence companion, and for four systems, even a brown dwarf companion is not consistent with the orbital parameters. The application of recent models of white dwarf donors in ultracompact binaries (e.g., Deloye & Bildsten 2003; Deloye et al. 2007) shows that, within certain temperature ranges, white dwarfs are a viable model for the companion in the most compact of the X-ray millisecond pulsar systems.

In this Letter, we report the discovery of SWIFT J1756.9–2508 (hereafter J1756), with a frequency of 182.07 Hz. It was discovered by the *Swift Gamma-Ray Burst Explorer* Burst

Alert Telescope (BAT) as part of the hard X-ray transient monitor program. In § 2, we discuss the observations made with *Swift* and the *Rossi X-Ray Timing Explorer* (RXTE). In §§ 3 and 4, we discuss the timing and spectral analysis, and in § 5 we derive limits on the companion mass and compare this source to the other known accreting X-ray millisecond pulsars. All errors are at the 90% confidence level unless otherwise stated.

2. OBSERVATIONS AND DATA REDUCTION

2.1. *Swift*

J1756 was discovered using the *Swift* BAT hard X-ray transient monitor (Krimm et al. 2006). The BAT is a large field-of-view instrument that continually monitors the sky to provide the gamma-ray burst trigger for *Swift*. On average, more than 70% of the sky is observed on a daily basis. Sky images are processed in near real time to detect astrophysical sources in the 15–50 keV energy band.¹²

J1756 was detected in the daily full-sky mosaics that are produced in the transient monitor. It first rose to detectability (Fig. 1) on 2007 June 7 (MJD 54,258) with an average rate in *Swift* BAT of 0.0083 ± 0.0015 counts $s^{-1} cm^{-2}$, corresponding to approximately 37 mcrab. The discovery was announced on June 13 (Krimm et al. 2007). The BAT data for all pointings when the source was fully coded (visible to all BAT detectors) were processed using the *Swift* analysis tools, which included cleaning out the diffuse background and flux from other bright sources in the field of view.

The discovery of J1756 triggered a target of opportunity (ToO) observation in the *Swift* X-Ray Telescope (XRT) beginning at 15:25:32 UT, 2007 June 13. It was intermittently observed over 13 days for a total exposure of 60 ks. The XRT data were processed with standard procedures (xrtpipeline ver. 0.10.6 within FTOOLS in the Heasoft package ver. 6.1.2), filtering, screening, and grade selection criteria (Burrows et al. 2005). The photon counting data collected during the first three segments were corrected for pileup and for the presence of single-reflection

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¹² See <http://swift.gsfc.nasa.gov/docs/swift/results/transients> for details on the transient monitor.

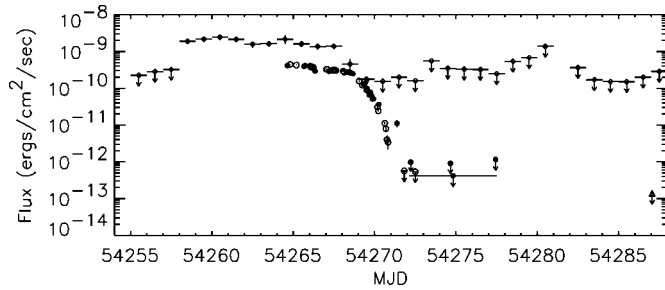


FIG. 1.—Combined light curve of SWIFT J1756.9–2508 from *Swift* BAT (diamonds), *Swift* XRT (filled circles), *RXTE* PCA (open circles), and *Chandra* (triangle; Papitto et al. 2007). The BAT points show flux in the 15–50 keV band, *Chandra* the 0.3–8.0 keV band, and the other points the 2–10 keV band. The PCA data is contaminated by emission from the Galactic ridge. To correct for this, a constant baseline of 1.24×10^{-10} ergs cm $^{-2}$ s $^{-1}$ was empirically determined and subtracted from the PCA data. Data points with $<2\sigma$ significance are shown as upper limits. For comparison with the archival upper limits quoted in § 2.2, for the PCA (BAT) 1 mcrab $\approx 2.4(5.2) \times 10^{-11}$ ergs cm $^{-2}$ s $^{-1}$.

rings due to the nearby bright source GX 5–1.¹³ For the windowed timing data, we extracted source events in a square region with a side of 20 pixels. Ancillary response files were generated with *xrtmkarf* and account for different extraction regions, vignetting, and point-spread function (PSF) corrections. The light curve (Fig. 1) was corrected for PSF losses and for vignetting. The best-fit XRT position is shown in Table 1.

2.2. Rossi X-Ray Timing Explorer

The discovery of J1756 also triggered a ToO observation in the *RXTE*, starting at 18:42:00 UT, 2007 June 13. Because GX 5–1 is within the $\approx 1^\circ$ field of view of the *RXTE* Proportional Counter Array (PCA) collimator, the PCA pointing direction was offset by about $18'$, primarily eastward of the true pulsar position. The PCA light curve (Fig. 1) has been corrected to effective on-axis values using the collimator response function, and spectral response functions account for the offset as well. The offset pointing direction is within $3'$ of the Galactic ridge, and so contamination from diffuse emission is present. The amount of contamination was found by examining other nearby PCA observations, when J1756 was quiescent, and fitting with a simple phenomenological model (and assuming it is constant). The contaminating flux was about ≈ 3 mcrab. The total exposure was ≈ 104 ks in two observing programs.

Archived monitor data from *RXTE* and *Swift* were searched for earlier outbursts of J1756. The PCA performs regular monitoring observations of the Galactic bulge (Swank & Markwardt 2001), and a search of archived PCA bulge scan data shows that the source has not been detected in the 2–10 keV band over the past 8.4 yr, with an upper limit of about 5 mcrab (2–10 keV).¹⁴ This limit is likely to be systematics driven, based on the nearby bright source GX 5–1. Archived daily BAT monitor data (15–50 keV) were searched back 2.5 yr, and the *RXTE* All-Sky Monitor (ASM) light curves (2–10 keV) were searched back 11.4 yr with no evidence of the source in either instrument with 3σ upper limits of ≈ 15 mcrab (BAT) and ≈ 50 mcrab (ASM). Due to observing constraints (mostly for Sun avoidance), the overall good coverage of the three detectors is $\approx 87\%$. With these gaps and the fact that even the 2007 outburst was not seen strongly

¹³ XRT observations after MJD 54,278 are not usable due to severe contamination from GX 5–1.

¹⁴ The 2007 outburst of J1756 was not seen in the PCA bulge scan because *RXTE* spacecraft slewing constraints prevented observations in early June. The most recent observation before the outburst was 2007 June 6.

TABLE 1
TIMING PARAMETERS OF SWIFT J1756.9–2508

Parameter	Value
Right ascension, α (J2000.0) ^a	17 ^h 56 ^m 57.35 ^s
Declination, δ (J2000.0) ^a	25°06′27.8″
Barycentric pulse frequency, f_p (Hz)	182.065804253(72) ^b
Pulsar frequency derivative, $ \dot{f} $ (Hz s $^{-1}$)	$<1 \times 10^{-12c}$
Projected semimajor axis, $a_s \sin i$ (lt-ms)	5.942(27)
Binary orbital period, P_b (s)	3282.104(83)
Time of ascending node, T_{asc}^d	54265.28707(6)
Orbital eccentricity, e^e	$<0.026^e$
Pulsar mass function, f_x ($10^{-7} M_\odot$) ^e	1.56(3)
Minimum companion mass, M_c ($10^{-3} M_\odot$) ^e	6.7–9.2 ^f
Maximum power, Z_{max}^2	2061

^a The estimated position uncertainty is $3.5''$. Position is from the XRT and held fixed as a timing parameter.

^b Uncertainties are 1σ in the last quoted digits.

^c 95% upper limit.

^d Modified Julian days, referred to barycentric dynamical time.

^e Derived parameter.

^f For neutron star masses of 1.4 – $2.2 M_\odot$.

in the ASM, it is possible that an earlier outburst was missed. Thus, we set a tentative recurrence timescale of ≥ 10 yr.

2.3. Other Observations

The discovery announcement led to several multiwavelength observations of J1756. A possible fading near-IR counterpart was announced with $K_s = 19.7$ on MJD 54,270 and $K_s = 21.0$ on MJD 54,282 (Burderi et al. 2007). However, an observation with the Westerbork Synthesis Radio Telescope during the X-ray outburst (1.4 GHz; MJD 54,266.103; Hessels & Stappers 2007) showed no sign of coherent radio pulsations. Furthermore, J1756 was not detected in observations after the outburst in either 8.7 GHz radio (MJD 54,277.371; Possenti et al. 2007) or 0.3–8.0 keV X-ray (*Chandra* upper limit of 1.4×10^{-13} ergs cm $^{-2}$ s $^{-1}$, starting MJD 54,287.059; Papitto et al. 2007) observations.

3. TIMING ANALYSIS

Figure 1 shows that the outburst of J1756 had a very rapid rise (≤ 1 day). After this the BAT flux was nearly constant for the next several days. For the first several days of XRT and PCA observations (up to MJD $\approx 54,268.6$), the flux in both instruments showed a slow decline, which can be fit to a nearly exponential decay, $e^{-t/\tau}$, where $\tau = 7.3 \pm 0.4$ days. This matches fairly closely the decay in the BAT light curve. After this date, the flux in all three instruments began a steeper decline ($\tau = 0.6 \pm 0.3$ days) and was undetectable even in the PCA after MJD 54,271, a mere 13 days after the start of the outburst. This light curve is similar in form to those of three other X-ray millisecond pulsars: SAX J1808.4–3658 (Bildsten & Chakrabarty 2001), IGR J00291+5934 (Galloway et al. 2005), and in particular XTE J1751–305 (Markwardt et al. 2002), which showed very similar values of τ (Gierliński & Putanen 2005). The three other objects have had more than one outburst, on recurrence timescales of ~ 2 – 3 yr, and all outbursts of XTE J1751–305 and IGR J00291+5934 and most outbursts of SAX J1808.4–3658 have been short (≤ 15 days), although notably the 2000 (Wijnands et al. 2001) and 2002 (Markwardt et al. 2005) outbursts of SAX J1808.4–3658 were longer and more erratic. Given the nondetection of earlier outbursts, J1756 is alone among the fast-decaying millisecond pulsar systems in not showing recurrence.

We searched for high-frequency variability by constructing

a fast Fourier transform–based power spectrum from the PCA data. The spectrum showed a significant, narrow excess near 182 Hz (Markwardt et al. 2007), which was also found in following observations. The feature was several millihertz wide, which suggested an X-ray pulsar with orbital Doppler broadening. Examination of power spectra on short (≈ 300 s) intervals confirmed a sinusoidal orbital modulation, making J1756 the eighth known accreting millisecond X-ray pulsar and the one with the longest rotation period.

We performed precision timing analysis on the PCA data using the techniques described in Markwardt et al. (2002). X-ray event data were recorded in the E_125us_64M_0_1s mode, and event times were corrected to the solar system barycenter using the position determined from XRT imaging (see Table 1). The Z^2 statistic (or Rayleigh statistic) was computed, and the orbital/timing parameters were adjusted to produce the largest Z^2 power of the pulse fundamental. This procedure is equivalent to radio or X-ray pulse time-of-arrival fitting (Buccheri et al. 1983) but with no binning. After adjustment, we found no long-term or orbital systematic trends. The binary orbit model was the ELL1 from TEMPO software version 11.010,¹⁵ which is designed for nearly circular orbits. The orbital parameters are shown in Table 1 and confirm that this is a very compact and circular orbit. The pulsation frequencies from MJD 54,264.78 to 54,267.07, folded on a trial orbital period, clearly establish the orbital frequency modulation. The pulsed semiamplitude is a nearly constant $\approx 6\%$ throughout the observations. No pulsed signal was seen in the 15–50 keV band from 250 s of BAT events starting at MJD 54,266.065, with a 2σ limit of $\leq 13\%$. With the present data, we can place only weak upper limits on the pulse frequency derivative (\dot{f}), since the time baseline is quite short.

The data from both the PCA and XRT (which cover 100% of the orbit) were searched for modulation at the orbital period. The PCA light curve (2–10 keV) shows no modulation at the orbital period, after the long-term trend is subtracted. The 2σ upper limit to orbital modulations is $<1.2\%$ from the PCA data. There is also no significant modulation seen in the XRT data. Furthermore, there are no signs of eclipses, dips, or bursts. We searched for rapid variability by constructing power spectra of 64 s segments of event data, with a Nyquist frequency of 2048 Hz, and averaged contiguous power spectra. While some variability is present, especially below ~ 200 Hz, a detailed timing study is beyond the scope of this Letter. We do not find any features above 200 Hz that are obvious kilohertz quasi-periodic oscillations (QPOs). For the summed outburst, the 95% upper limits on the fractional rms are about 6% (9%) for 16 Hz (128 Hz) wide QPOs and a factor of ~ 2 larger for a single observation.

4. SPECTRAL ANALYSIS

We fit the *Swift* XRT spectrum from each segment in the 0.5–9 keV energy range with different models, which included a single absorbed power law (PL; Wilms et al. 2000) and an absorbed blackbody (BB).

A simple comparison of null hypothesis probabilities clearly favors the PL with respect to the BB model (20% vs. 4%, respectively), with the BB fit producing systematic residuals both at low and high energies. The derived absorbing column density is always in excess of the Galactic value ($N_{\text{H}}^{\text{G}} = 1.47 \times 10^{22} \text{ cm}^{-2}$; Dickey & Lockman 1990), and there is no indication of a softening of the spectrum with time. Representative spectral fits are shown in Table 2. The source was de-

TABLE 2
SAMPLE SPECTRAL FITS FOR SWIFT J1756.9–2508

Instrument	MJD	Model ^a	N_{H}	Γ [kT] ^b	$\chi^2_{\text{red}}/\text{dof}$
XRT	54,266.06	PL	5.44 ± 0.41	2.08 ± 0.14	1.081/200
	54,266.06	BB	2.82 ± 0.23	$[1.30 \pm 0.05]$	1.185/200
PCA	54,264.79	PL	5.44^c	2.00 ± 0.14	0.878/47
	54,267.46	PL	5.44^c	1.95 ± 0.08	0.514/48
BAT	54,258	PL	...	2.3 ± 0.7	0.20/6
	54,266	PL	...	2.8 ± 0.5	0.38/6

^a PL = power law; BB = absorbed blackbody.

^b Numbers are Γ for PL fits and kT (in square brackets) for the BB fit.

^c Parameter held fixed.

tected in the PCA up to energies ≈ 30 keV, and the spectrum is consistent with an absorbed PL with absorption fixed at the best-fit value of XRT segment 003. The BAT data from 14 to 195 keV were fit by a simple PL model. The BAT spectrum is softer than the spectra found from fits to the PCA alone or XRT alone, suggestive of a cutoff to the PL. However, joint fits to the BAT, XRT, and PCA spectra are not improved with a cutoff PL model.

Given the uncertainty in the spectrum at high energies, we estimate the integrated outburst fluence to be $\approx (4.5 \pm 0.8) \times 10^{-4} \text{ ergs cm}^{-2}$ (1–10,000 keV). Using the angular proximity of the source to the Galactic center to estimate the distance, we can, following Bildsten & Chakrabarty (2001), estimate the time-averaged mass accretion rate, $\dot{M}_{\text{x}} = (9.3 \pm 1.6) \times 10^{-13} M_{\odot} \text{ yr}^{-1} d_8^2 m_{1.4}^{-1} T_{10}^{-1}$, where d_8 is the distance in units of 8 kpc, $m_{1.4} = M_{\text{x}}/(1.4 M_{\odot})$, and we assume a neutron star radius of 10 km. The parameter T_{10} is the recurrence time in units of 10 yr, and only $T_{10} \geq 1$ is allowed by the observations.

5. DISCUSSION

In combination, the measured mass function and binary orbital period P_b constrain the donor’s mass-radius ($M_{\text{C}}-R_{\text{C}}$) relation, assuming the donor fills its Roche lobe. For J1756, $P_b = 54.7$ minutes, a value near the median of the known X-ray millisecond pulsars.

If we assume a particular value for M_{x} , we can calculate possible values of M_{C} from the mass function, $f_{\text{x}} = (M_{\text{C}} \sin i)^3 / (M_{\text{x}} + M_{\text{C}})^2$, where i is the binary inclination to our line of sight and M_{x} and M_{C} are the masses of the pulsar and companion, respectively. The lack of eclipses allows us to set a limit $i < 85^\circ$ (Paczynski 1971). For this inclination and $M_{\text{x}} = 1.4 M_{\odot}$, we have $M_{\text{C}, \text{min}} = 0.0067 M_{\odot}$. Assuming a uniform distribution of possible values of $\cos i$, we can set a 95% confidence level (CL) upper limit on the mass of $M_{\text{C}} < 0.022 M_{\odot}$ (for $M_{\text{x}} = 1.4 M_{\odot}$). The range of possible values of M_{C} up to this limit and corresponding companion radius R_{C} are shown in Figure 2. If we consider a larger neutron star of $M_{\text{x}} = 2.2 M_{\odot}$, the limits (95% CL) are $M_{\text{C}} < 0.030 M_{\odot}$ and $R_{\text{C}} < 0.069 R_{\odot}$. This radius is less than the minimum radius for low-mass hydrogen-rich brown dwarfs of any age (Chabrier et al. 2000; Bildsten & Chakrabarty 2001; Fig. 2), as expected for a binary period of ≤ 80 minutes (Paczynski & Sienkiewicz 1981; Nelson et al. 1986). Thus, we restrict our attention to white dwarf (WD) donor models.

The dotted and dashed curves in Figure 2 show evolutionary sequences of ultracompact binaries with He donor models originally developed for a study of the AM CVn (double WD) class of ultracompact binaries (Deloye et al. 2007). Along these curves, regions of negative slope correspond to adiabatic donor evolution. In regions of positive slope, the donor is cooling and contracting toward the fully degenerate He WD $M_{\text{C}}-R_{\text{C}}$ relation (see

¹⁵ See <http://www.atnf.csiro.au/research/pulsar/tempo>.

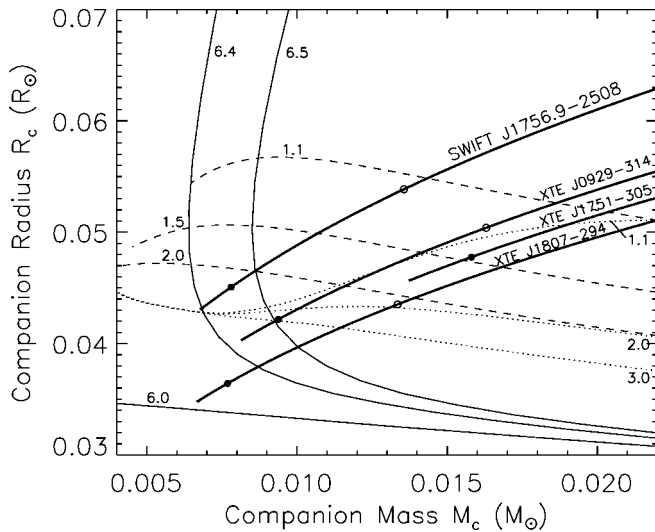


FIG. 2.—Companion radius (R_c) vs. mass (M_c) plane, showing the Roche lobe constraints for SWIFT J1756.9–2508 and, for comparison, three other X-ray millisecond pulsars with low-mass companions. The source curves trace possible values of the inclination angle, i . The filled circles indicate $i = 60^\circ$, and the open circles indicate $i = 30^\circ$. The other curves are various white dwarf models. The dotted and dashed lines show evolutionary sequences with He donor models calculated as in Deloye et al. (2007) (see text for full explanation). The dotted curves show cases where no accretion flux is reprocessed. The dashed curves show the case where 10% of accretion flux, as seen by the donor, is thermalized in the donor’s atmosphere. The numbers next to these curves indicate $\log \psi_c$, where higher values of the central degeneracy parameter ψ_c (defined in Deloye et al. 2007) indicate a more degenerate and hence compact donor. Within each set of models, evolution proceeds to lower M_c and less degenerate donors evolve at larger R_c . The solid curves are for a pure C composition from the models of Deloye & Bildsten (2003). The number next to each solid curve indicates the log of the temperature, in units of kelvins, used in the model.

Deloye et al. 2007), essentially traced by the lowest dotted line in Figure 2. The donor’s cooling is affected by the (very uncertain) efficiency with which irradiating flux generated by the accretion—which always dominates the donor’s own flux in these ultracompact binaries—is thermalized in the donor’s upper

atmosphere. The impact of irradiation is to extend the donor’s adiabatic evolution phase and to slow its rate of cooling afterward.

Since calculations similar to the Deloye et al. (2007) He donor study do not yet exist for C/O donors, we compare the J1756 constraints to a sample of pure C models of Deloye & Bildsten (2003, solid curves in Fig. 2); realistic C/O WDs will lie at somewhat smaller R_c . These models assume fully convective interior profiles and do not explicitly treat thermal transport. However, the physics that produces the transition from adiabatic evolution to cooling in He donors does not depend on composition and C/O donors will also experience a cooling phase. This means a substantial fraction of the phase space for “hot” C/O donors shown in Figure 2 will likely not be evolutionarily accessible (e.g., the nearly vertical regions of the $\log(T_c/\text{K}) > 6.4$ “isotherms”), making a C/O donor in J1756 rather unlikely. Thus, J1756 likely harbors a He-dominated donor whose thermal cooling has been at least modestly slowed. Irradiation provides a natural mechanism that is able to affect the necessary degree of slowing, as discussed above. This interpretation places J1756 in the company of at least two other X-ray millisecond pulsars with very low mass companions: XTE J0929–314 (Galloway et al. 2002) and XTE J1807–294 (Falanga et al. 2005); with its longer orbital period, J1756 provides a unique probe among the known ultracompact low-mass X-ray binaries and appears to indicate that irradiation plays at least a modest role in the evolution of the donors in these systems.

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ERRATUM: “DISCOVERY OF THE ACCRETION-POWERED MILLISECOND PULSAR SWIFT J1756.9–2508
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Table 1 contained one error. The time of ascending node T_{asc} was incorrect due to a numerical conversion error in the preparation of the original table text. The correct value is $T_{\text{asc}} = 54264.78707(6)$. The correct value was used in all of the analysis leading up to the paper. As T_{asc} is a purely fiducial reference time; the scientific conclusions of the paper are unchanged.

We thank A. Patruno for bringing this error to our attention.