

THE GRAVITATIONAL LENS CFRS 03.1077¹

DAVID CRAMPTON² AND DAVID SCHADE²

Dominion Astrophysical Observatory, Herzberg Institute of Astrophysics, National Research Council of Canada, 5071 West Saanich Road, Victoria, BC V9E 2E7, Canada; David.Crampton@nrc.ca

F. HAMMER² AND A. MATZKIN

Observatoire de Paris, Section de Meudon, DAEC, F-92195 Meudon Principal Cedex, France

S. J. LILLY²

Dominion Astrophysical Observatory, Herzberg Institute of Astrophysics, National Research Council of Canada, 5071 West Saanich Road, Victoria, BC V9E 2E7, Canada

AND

O. LE FÈVRE²

Laboratoire d'Astronomie Spatiale, F-13376 Marseille Cedex 12, France

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ABSTRACT

An exquisite gravitational arc with a radius of $2''$ has been discovered around the $z = 0.938$ field elliptical galaxy CFRS 03.1077 during *HST* observations of Canada-France Redshift Survey (CFRS) fields. Spectroscopic observations of the arc show that the redshift of the resolved lensed galaxy is $z = 2.941$. This gravitational lens source system is well fitted using the position angle and ellipticity derived from the visible matter distribution and an isothermal mass profile with a mass corresponding to $\sigma = 387 \pm 5 \text{ km s}^{-1}$. Surprisingly, given the evidence for passive evolution of elliptical galaxies, this is in good agreement with an estimate based on the fundamental plane for $z = 0$ ellipticals. This, perhaps, indicates that this galaxy has not shared in the significant evolution observed for average elliptical galaxies at $z \sim 1$. A second elliptical galaxy with similar luminosity from the CFRS survey, CFRS 14.1311 at $z = 0.807$, is also a lens, but in this case the lens model gives a much smaller mass-to-light ratio; i.e., it appears to confirm the expected evolution. This suggests that this pair of field elliptical galaxies may have very different evolutionary histories, which would be a significant result if confirmed. Clearly, CFRS 03.1077 demonstrates that these “Einstein rings” are powerful probes of high-redshift galaxies.

Subject headings: galaxies: individual (CFRS 03.1077) — gravitational lensing

1. INTRODUCTION

A very beautiful arc structure, characteristic of a lensed object, was noted around the galaxy CFRS 03.1077 (Hammer et al. 1995) during analysis of *HST* images of high-redshift field galaxies. CFRS 03.1077 is a bright elliptical galaxy at $z = 0.938$ and, hence, is likely to be lensing background galaxies or quasars. Miralda-Escudé & Lehár (1992) predicted that ~ 100 of gravitationally lensed rings per square degree should be observable at optical wavelengths to $B \sim 26$, given the number of Einstein rings at radio wavelengths and the projected density of galaxies in optical surveys. Indeed, Ratnatunga, Griffiths, & Ostrander (1999) report finding 10 good galaxy lens candidates in 0.1 deg^2 , based on analyses of *HST* Medium Deep Survey images. These, however, have not yet been confirmed and remain only as candidates.

It is now recognized that galaxies that act as gravitational lenses are extremely important tools for studying a variety

of cosmological problems, including the nature and evolution of the lensing galaxies themselves (Kochanek et al. 2000; Narayan & Bartelmann 1999; Refsdal & Surdej 1994; Schneider 1996). For example, Im, Griffiths, & Ratnatunga (1997) showed that observations of seven galaxy lenses favor a nonzero cosmological constant, while Kochanek et al. (2000) and Keeton, Kochanek, & Falco (1998) present results on the properties of the lensing galaxies, most of which are early type. Since the lens models give data on the total (dark+visible) distribution of mass, investigations like the latter are extremely powerful, especially for studying the properties of galaxies at high redshifts in which traditional methods become increasingly difficult. Einstein ring lenses provide additional constraints and, hence, are even more important to the determination of mass distributions of the lensing galaxies (Kochanek, Keeton, & McLeod 2001).

Even though Einstein ring lenses are predicted to be much more common at optical wavelengths than at radio wavelengths, only one optically identified system has been confirmed thus far: 0047–2808 (Warren et al. 1998, 1999). Rather than being initially recognized from an image as a “ring,” this lens was detected through superposition of emission lines from the lensed object on the spectrum of the foreground galaxy. Hewett et al. (2000) and Hall et al. (2000) recently found additional candidate lenses in this manner, but these remain to be confirmed. These authors argue that spectroscopic observation of distant early-type

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² Visiting Astronomer, Canada-France-Hawaii Telescope, operated by the National Research Council of Canada, the Centre de la Recherche Scientifique de France, and the University of Hawaii.

galaxies is one of the best methods of detecting galaxy lenses, particularly because the presence of spectroscopic features from the source implies that the crucial redshifts of both the source and deflector can be determined. In practice, however, this method is only viable for background galaxies with strong emission lines. Furthermore, galaxy lens candidates are quite easily recognizable with images from *HST* or ground-based adaptive optic systems, and disentangling the spectral features arising from the source and deflector galaxies is no longer as difficult with the improved spatial resolution being achieved on modern large ground-based telescopes.

In this paper, we report on a detailed analysis of the *HST* images of the lens and source galaxies related to CFRS 03.1077. We also present spectra of these galaxies obtained with the Canada-France-Hawaii Telescope (CFHT) and derive the redshift of the source—a galaxy without strong emission lines. Finally, a simple model of the lens source system is described.

2. OBSERVATIONS

2.1. *HST* Imaging

HST Wide Field and Planetary Camera 2 (WFPC2) observations of the CFRS 0300+00 field were obtained on 1996 January 10 (Lilly et al. 1995) during a survey of the morphology of high-redshift field galaxies in CFRS fields. The total exposure time for the F814W images of the CFRS 0300+00 field was 6700 s. Even though CFRS 03.1077 was not located in the higher resolution PC part of the WFPC2 field, the arc structure near the galaxy was immediately obvious while preprocessing the data (see Fig. 1a) and was clear in the individual frames before stacking.

2.1.1. *The Lensing Galaxy*

The J2000.0 coordinates of CFRS 03.1077 are R.A. = $03^{\text{h}}02^{\text{m}}30^{\text{s}}.9$ and decl. = $00^{\circ}06'02''.1$. The arc has a radius of $2''.1 \pm 0''.1$ and is accurately centered (to within $\pm 0''.05$) on the galaxy.

Special attention was paid to deriving the parameters of CFRS 03.1077. A high signal-to-noise ratio (S/N) point-spread function (PSF) was determined from six stars located at similar locations on images of other fields (since no suitable nearby star was present), and several other fainter stars were used to assess possible errors. Fitting was carried out on a “symmetrized” image (see Schade et al. 1995, 1996 for

details) of the galaxy so that the procedure was not perturbed by the presence of the arc or by any other images. The best-fit value of the reduced χ^2 was 0.9988 for a fit radius of 20 pixels. Subtraction of the model galaxy (see Fig. 1b) demonstrates that the resulting fit is an extremely good representation of the galaxy’s luminosity distribution.

The parameters of the galaxy were measured from the combined F814W image, with a total integration time of 6700 s. The errors in the parameters were estimated from the fits to the five individual *HST* images with exposure times of 1100, 1200, or 2100 s. The dispersions among these fits were $0''.13$ in R_e , 0.025 in axial ratio (b/a), $2''.8$ in position angle, 0.04 mag in total magnitude, and 0.06 mag in central surface brightness. The errors in the stacked image are expected to be roughly $\sqrt{5}$ times smaller because of better S/N. These errors are also likely to be systematically overestimated because cosmic rays are not properly removed in the individual images. Possible systematic errors due to imperfect knowledge of the PSF were estimated by repeating the fits with several different PSFs from different observations and different positions on the chip. The results show that these systematic errors are a few percent or less in all parameters. The derived parameters depend weakly on the choice of the fitting radius (the radius over which χ^2 was computed), and these were estimated from simulations of this galaxy with realistic backgrounds. For reasonable values of fitting radius, the errors were smaller than 3% but reached extreme values of 6% for fitting radii much smaller than the half-light radius. We adopt errors corresponding to those from the individual image fits divided by $\sqrt{5}$ and then add 5% in quadrature to allow for possible systematic errors. Details of the parameters of the lensing galaxy are summarized in Table 1.

At this redshift, the lensing galaxy has a half-light radius of $R_e = 14.7 \pm 1$ kpc and $M_{\text{AB}}(B) = -22.77 \pm 0.1$ for $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $q_0 = 0.5$. The $V-I$ color of the galaxy is typical of galaxies with $z \sim 1$ in the CFRS sample. The computed rest-frame color is $(U-V)_{\text{AB}} = 1.53$.

2.1.2. *The Arc and Counterimages*

The measured properties of the arc image are given in Table 2. The arc has a length of $2''.1$ and is centered on the galaxy to within $\pm 0''.05$ (the precision of the measurement). The observed surface brightness is $21.2 \text{ mag arcsec}^{-2}$ in the brightest pixel, which corresponds to 16.73 in rest-frame AB(2030) magnitudes. This surface brightness is modulated by pixellation and by the convolution with the *HST* PSF. For a compact exponential disk, the implied central surface

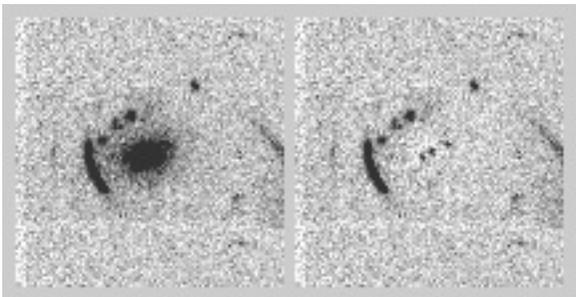


FIG. 1.—*Left*: *HST* F814W image of CFRS 03.1077 showing the bright arc to the east. *Right*: Same image after the subtraction of the “symmetrized” galaxy, as discussed in the text. The images are 10 arcsec square; north is up, and east is to the left.

TABLE 1
PARAMETERS OF THE LENS CFRS 03.1077

Parameter	Value
R_e (arcsec).....	1.74 ± 0.1
b/a	0.67 ± 0.03
P.A. (deg).....	104.4 ± 3
I_{AB} (mag).....	20.36 ± 0.05
$(V-I)_{\text{AB}}$ (mag).....	2.13 ± 0.15
$M_{\text{AB}}(B)$ (mag).....	-22.77 ± 0.1
$(U-V)_{0,\text{AB}}$ (mag).....	1.53 ± 0.2

NOTE.—Assumes $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $q_0 = 0.5$.

TABLE 2
PARAMETERS OF THE ARC NEAR CFRS 03.1077

Parameter	Value
z	2.941
Length (arcsec).....	2.1
Width (arcsec).....	0.17 ± 0.02
$SB(I_{AB,max})$ (mag).....	21.2
I_{AB} (mag).....	21.94
$M(B)_{AB}$ (mag).....	-22.5

NOTE.—Assumes $H_0 = 50 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $q_0 = 0.5$. $M(B)_{AB}$ is calculated assuming a lensing magnification factor of 10. It also assumes a very blue spectral energy distribution and corresponds to a lower limit on the B -band luminosity. The absolute magnitude at $AB = 2030 \text{ \AA}$ is 0.41 mag fainter.

brightness would be about 2 mag brighter than the observed surface brightness, and for a compact $R^{1/4}$ bulge the implied central surface brightness would be higher by ~ 7 mag. The exact values depend on the details of the luminosity profiles.

The arc is resolved across its narrowest dimension; i.e., it has nonzero intrinsic width, compared with the PSF of a star on the same frame. One-dimensional Gaussian fits across the arc at various points along its length give a median observed width (FWHM) of $0''.23$ and a dispersion of $0''.025$. For comparison, a fit to a star gives $0''.18$, which implies an equivalent Gaussian width of $\sim 0''.17$ for the arc itself. Figure 2 shows the brightness profile along the length of the arc (and integrated across the profile). The arc centroid position was traced and the flux extracted from the image as if it were a two-dimensional spectrum. The error bars are based on the number of electrons in the image and on the observed sky noise. This figure confirms the visual impression that there is real variation in surface brightness along the length of the arc and the peak of the profile occurs at the center, which is consistent with a normal galaxy luminosity profile.

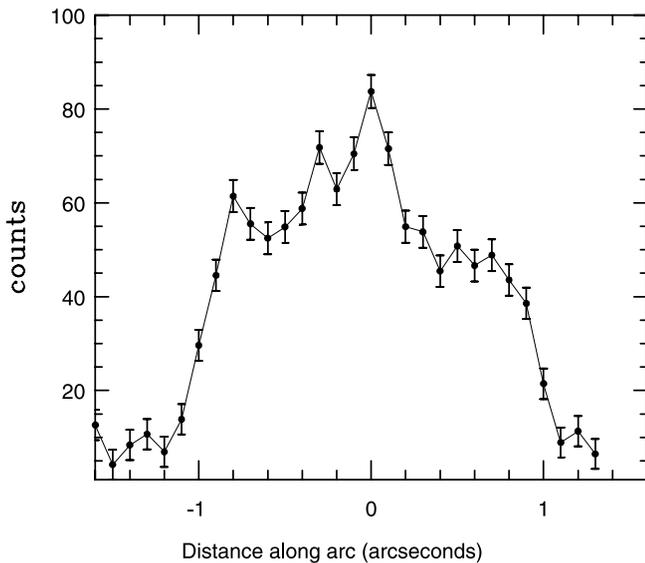


FIG. 2.—Distribution of intensity along the arc (integrated across its width).

In addition to the arc itself, other faint images are evident, especially in the subtracted image (Fig. 1*b*). Large magnification events such as arcs are generally due to the presence of a source near the caustic lines (e.g., Refsdal & Surdej 1994). In simple lens geometry configurations (spherical or elliptical), they are usually accompanied by additional images on the opposite side of the lens. Models of the CFRS 03.1077 lens indicate that faint images near the position of the lensing galaxy (visible near the center of the lens in Fig. 1*b*) are likely to be counterimages to the arc. However, they could also be small objects that are embedded within the lensing galaxy or simply a projection of foreground/background objects (with the only reservation being that if they are unlensed objects, then they should then be at a redshift that is significantly lower than that of the arc source). Unfortunately, we do not have color information of these images. Since their location suggests that they are counterimages of the arc, and in the absence of other data, we assume in the following that they are, in fact, lensed images.

2.2. CFHT Spectroscopy

A spectrum of the lensing galaxy CFRS 03.1077 from the original CFRS survey is represented in Figure 3. Based mostly on the location of the 4000 \AA break and its overall spectral energy distribution, Hammer et al. (1995) assigned a redshift of $z = 0.938$ with a high confidence level (class 3; see Le Fèvre et al. 1995). The strong 4000 \AA break and absence of any $[\text{O III}]$ emission indicates that the galaxy is likely to be an elliptical. The slit was positioned east-west, and its width was $1''.75$, so part of the arc was in the Multiple Object Spectrograph (MOS) slit, and indeed, the spectrum of a neighboring object, i.e., the arc, was noted by the CFRS team, but no redshift was derived. After recognizing that this was probably a lensed galaxy, this spectrum was rereduced and analyzed, yet no redshift determination was possible, partly because the spectrum of the arc was overwhelmed by that of the bright galaxy. (This is unlike the case of CFRS 14.1311, in which Crampton et al. [1996] were able to subsequently detect strong emission lines of the background quasar.)

Spectra of the arc alone were obtained with the Optionally Stabilized Imager and Spectrometer (OSIS) on CFHT on 1997 August 26–28. Unfortunately, the seeing was only average during these exposures ($0''.5\text{--}0''.8$), and the OSIS fast-guiding system, which would have improved the image quality somewhat, was not available. A $0''.75$ slit was aligned along the brightest part of the arc at P.A. = 20° on two of the nights, but a $1''.0$ slit had to be used during poorer seeing on 1997 August 27. Two spectra were obtained on 1997 August 26; three were obtained on August 27, and five were obtained on August 28. All exposures were 1800 s, so the total exposure is equivalent to 5 hr. Small shifts of the arc along the slit were made between exposures to improve the flat-fielding and, hence, S/N of the final spectra. The spectra were recorded with the Space Telescope Imaging Spectrograph (STIS) 2 CCD with $21 \mu\text{m}$ pixels that correspond to $3.4 \text{ \AA pixel}^{-1}$ with the V150 grism and $0''.14 \text{ pixel}^{-1}$ in the spatial dimension. The wavelength resolution, set by the $0''.75$ slit, is 17 \AA , and the usable wavelength range is $4000\text{--}9000 \text{ \AA}$.

The two-dimensional spectra were bias-subtracted and flat-fielded using a normalized flat derived from spectra of internal calibration lamps. Two-dimensional sections were

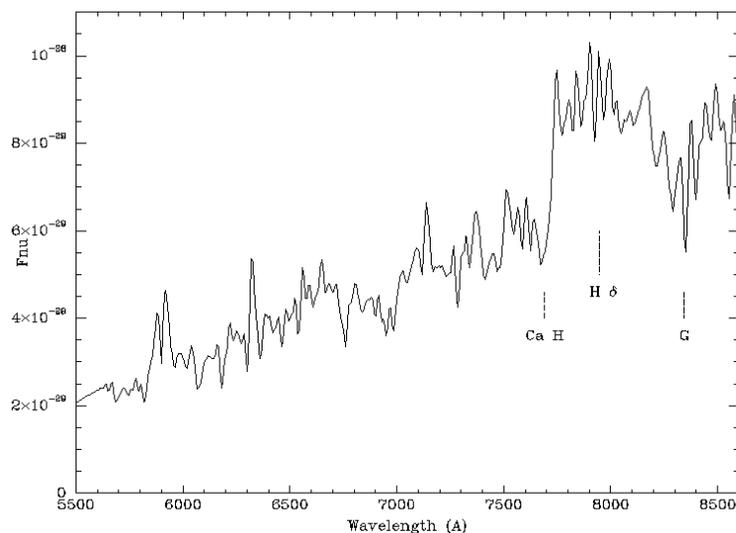


FIG. 3.—CFHT MOS spectrum of the lensing galaxy CFRS 03.1077. The locations of some typical features redshifted to $z = 0.938$ are marked, although the large 4000 Å break is the most distinctive feature.

then extracted from each image, which were centered on the spectrum of the arc. These were shifted to compensate for any wavelength shifts (determined from the night-sky lines), combined, and then a one-dimensional spectrum was extracted. The background-subtracted two-dimensional spectrum was inspected to see if there was any indication of spatial variations of the spectral features along the arc, but none was detectable. The spectrum width is ~ 14 pixels, or $\sim 2''$, which is consistent with length of the arc (see Figs. 1 and 2).

A portion of the spectrum of the arc is represented in Figure 4. A spectrum of Feige 24 was used to provide flux calibration, although the absolute calibration is uncertain because of the narrow slit width. The spectrum was cross-correlated against *HST* ultraviolet spectra of NGC 1741

and NGC 4214, which are two prototypical starburst galaxies (Conti, Leitherer, & Vacca 1996; Leitherer et al. 1996). Since no emission lines are obvious in our spectrum of the arc, the strong $\text{Ly}\alpha$ emission lines were first excised from these starburst spectra. Both of the resulting templates give a strong cross-correlation peak at $z = 2.94$, which is the same redshift as that derived from a visual examination of the spectrum of the arc. The rest wavelengths of the most prominent features familiar in nearby starburst galaxies are identified in Figure 4, with a spectrum of NGC 4214 shown for comparison. In common with many of the $z \sim 3$ galaxies observed, for example, by Steidel et al. (1996) and Pettini et al. (1997), no $\text{Ly}\alpha$ emission is visible; the continuum is flat and drops shortward of $\text{Ly}\alpha$, and numerous stellar and interstellar lines of C, O, Si, and Al are visible. As with all

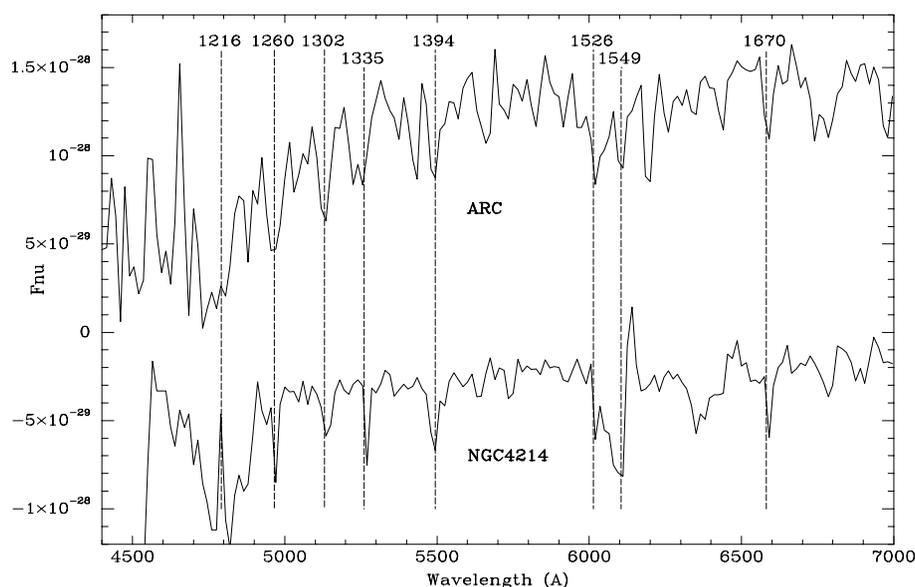


FIG. 4.—Spectrum of the arc (*top*) with a redshifted spectrum of the starburst galaxy NGC 4214 (Leitherer et al. 1996) shown for comparison. The flux scale of the latter was arbitrarily scaled and shifted for display purposes. The dashed lines indicate the nominal positions of common absorption lines, assuming $z = 2.941$ (their rest-frame wavelengths are indicated at the top of the figure).

such galaxies (Pettini et al. 1997), there is considerable scatter among the velocities derived from the individual features, and we adopt $z = 2.941$ with an estimated error of 0.008. We note that, in retrospect, some of the features identified in Figure 4 are visible on our original MOS spectrum of the arc, which lends additional weight to our redshift determination.

3. LENS MODELING

The lens CFRS 03.1077 was modeled assuming an isothermal mass profile. Gravitational bending angles have been computed using the code AIGLE (see Hammer et al. 1997) for an elliptical lens with an isothermal profile ($\rho = \rho_0[1 + (r/r_c)^2]^{-1}$), where r_c is the core radius. Mass distributions are used rather than potentials to avoid any further assumptions in mass estimations. The arc itself can be perfectly reproduced by the model, assuming a uniform circular source (radius $r_s = 0''.1$) at $z = 2.941$ and a lens with parameters provided by the light distribution (see Table 1). However, since the true lensing parameter is the mass enclosed within a circle defined by the arc, there is a clear degeneracy between the core radius and the line-of-sight (LOS) velocity dispersion (σ_{los}). It is beyond the scope of this paper to accurately evaluate the available range of acceptable parameters, but an initial evaluation gives $\sigma_{\text{los}} = 390 \pm 70 \text{ km s}^{-1}$ and small core radii (r_c from $0''.1$ to $0''.3$). It is more interesting to note that for a large fraction of the (r_c, σ_{los}) configurations, the model predicts additional images on the other side of the lensing galaxy. These fall within $0''.03$ of the locations of the two faint blobs on the opposite side of the lensing galaxy (discussed in § 2.1.2) if a slight offset of the lens position angle, from 104° to 122° , is assumed. Figure 5a shows a simulation of the three images (I0: the arc, I1, and I2: the blobs on the other side of the lens near its center). Figure 5b shows the location of the source relative to the area of high magnification. The source falls at the edge of the elliptical caustic lines, with most of its surface within the three-image area. The magnification factor (10 for the arc alone) provides a magnitude for the unlensed source of 24.4.

The requirement to properly reproduce this three-image configuration tightly constrains the lensing parameters.



FIG. 5.—*Left*: Model of the arc (I0) and the two counterimages (I1 and I2). The model assumes $\sigma_{\text{los}} = 390 \text{ km s}^{-1}$, $r_c = 0''.19$, ellipticity = 0.24, and P.A. = 122° for the lensing galaxy. The magnification factor is 12, and the ratio between the arc and counterimage luminosities is 0.17. *Right*: High-magnification areas (caustics) in the source plane. This has been computed assuming a source at $z = 2.941$, with $r = 0''.1$ (same parameters as for the lens in Fig. 5a). The faint elliptical boundary delimits the 1–3 image areas, and the diamond shape the 3–5 image areas. A source with an impact parameter of ($x = -0''.47$ and $y = -0''.2$) would provide three images (I0, I1, and I2) in the image plane, as shown in the panel on the left.

Indeed, the actual number of constraints is eight (six for the image positions and two for the intensity ratios), while the number of parameters is only six (r_c, σ_{los} , ellipticity, position angle, and the source location). This leads us to believe that I1 and I2 are likely counterimages of the arc, with the result being that the mass estimates of the lens are very accurate. Actually, the lensing model predicts an accuracy as small as 2 km s^{-1} for the lens velocity dispersion. The cosmological parameter (q_0 or Λ) is another factor of uncertainty. However, the gravitational bending angle depends on the ratio of the angular diameter distances between the lens and the source of the lens, and so the dependency on the cosmological parameters is small. For the present example, $q_0 = 0.5$ leads to $\sigma_{\text{los}} = 390 \pm 2 \text{ km s}^{-1}$, and $q_0 = 0$ leads to $\sigma_{\text{los}} = 384 \pm 2 \text{ km s}^{-1}$. Our determination of the mass parameter can be summarized by $\sigma_{\text{los}} = 387 \pm 5 \text{ km s}^{-1}$.

3.1. Evolution of Early-Type Galaxies

The photometric model fitting to the luminosity profile of the lensing galaxy that was described in § 2.1.1 gives the central surface brightness μ_0 and half-light radius R_e , which corresponds to an $R^{1/4}$ law. These values allow us to derive an expected velocity dispersion from the fundamental plane relations between size, surface brightness, and velocity dispersion. Jorgensen, Franx, & Kjaergaard (1996) give the required relations. The observed central surface brightness is $\mu_0 = 16.20$ in I_{AB} . At a redshift of $z = 0.938$, this corresponds to $\mu_{0,B}(\text{rest frame}) = \mu_{0,I} - 7.5 \log(1+z) - 0.13$, where the latter term is the correction from observed wavelength to B band. This value, $\mu_{0,B}(\text{rest frame}) = 13.92$, converts to $\mu_{0,r_{\text{sun}}} = 13.24$ and surface brightness at $R = R_e, \langle \mu_e \rangle_r = 20.18$. Using the relations given by Jorgensen et al. (1996) yields a velocity dispersion of $\sigma = 380_{-60}^{+65}$. This is in very good agreement with the mass parameter derived from the lensing model above. The fact that it agrees so well, however, poses another problem in that it apparently represents a contradiction to the passive evolution model for elliptical galaxies. For example, Schade et al. (1999) claim evolution of nearly 1 mag in luminosity (or, equivalently, surface brightness) for an elliptical at this redshift. In fact, the lensing galaxy is one of the sample of elliptical galaxies in the CFRS/low-dispersion survey spectrograph (LDSS) imaging survey that was used to derive the evolution. In order to estimate mass from the photometric parameters of this galaxy, the observed surface brightness should be “deevolved”; that is, the central surface brightness should be made 1 mag fainter before it is compared to the local ($z = 0$) fundamental plane. If this is done, then a mass estimate of $\sigma = 214_{-30}^{+35}$ is obtained, which is obviously very significantly lower than the mass estimate from the lens model.

This result for CFRS 03.1077 is contrary to results from studies of the evolution of the fundamental plane at moderate redshift (Kelson et al. 2000; van Dokkum et al. 2001) and of lensing galaxies on average. Kochanek et al. (2000) present data from 30 lenses and conclude that the fundamental plane for both field galaxies and cluster galaxies are similar and that they are representative of passively evolving populations formed at $z > 2$. That view is consistent with the fundamental plane studies. The reason that CFRS 03.1077 does not appear to share this behavior is not obvious. Is there substantial variance in the evolution history of massive elliptical galaxies? An examination of

Figure 3 from Schade et al. (1999) shows that this particular galaxy lies within its 1σ error bar of the no-evolution locus for elliptical galaxies in the $M_B(\text{AB})-\log R_e$ plane and is one of the most deviant points from the evolving relation. The only other very luminous elliptical galaxy in the study by Schade et al. (1999; CFRS 14.1311 at $z = 0.807$) also appears in Figure 3 of that paper and lies exactly on the evolving $M_B(\text{AB})-\log R_e$ relation, which implies an evolution in luminosity (or, equivalently, in surface brightness) of ~ 1 mag. Remarkably, that galaxy is an Einstein cross lens discovered by Ratnatunga et al. (1995) and studied by Crampton et al. (1996), who report a velocity dispersion of 230 km s^{-1} , as implied by the lens model. If the observed values of R_e and $\langle\mu_e\rangle$ for CFRS 14.1311 are converted into a velocity dispersion via the local fundamental plane, then a value of $\sigma = 537 \text{ km s}^{-1}$ is obtained. If an evolution of 1 mag in surface brightness is then applied, then this is reduced to $\sigma = 290 \text{ km s}^{-1}$, which is roughly consistent with the value derived from the lens model. In other words, CFRS 14.1311 behaves as expected, but CFRS 03.1077 does not. It should be noted that neither of these galaxies has detectable levels of [O II] 3727 emission that would indicate star formation, and furthermore, the galaxy with lower surface brightness (CFRS 03.1077) is the bluer of the two galaxies, which is the opposite of what might be expected if a recent burst of star formation was a factor.

In summary, we have two independent methods of estimating the velocity dispersion of these two lens galaxies. The photometric method uses size and surface brightness to estimate the expected velocity dispersion, assuming that these galaxies lie on either the local fundamental plane or on an evolved fundamental plane. The lens method is independent of the galaxy light and depends only on the lens geometry. The results of these two methods can be made to be consistent only if CFRS 03.1077 lies on the unevolved

fundamental plane, whereas CFRS 14.1311 has evolved by ~ 1 mag at $z = 0.807$. In other words, the results suggest that these two highly luminous field elliptical galaxies may have had very different evolutionary histories.

4. CONCLUSION

Our spectroscopic observations confirm that the arc surrounding the $z = 0.938$ elliptical galaxy CFRS 03.1077 is indeed a lensed image of a background galaxy. The redshift of this galaxy is $z = 2.941$. Standard lens models easily reproduce the observed arc structure and also suggest that two faint objects observed near the lensing galaxy on the opposite side to the arc are lensed images. Observations at other wavelengths should be obtained to determine the colors of these objects. If they are the same as the arc, then this would be further evidence that they are lensed images and can be confidently used to constrain the lens geometry. Multiwavelength observations of CFRS 03.1077 could also be used to examine whether the internal colors of the galaxy itself are normal or show strong variations indicative of recent star formation (cf. Menanteau, Abraham, & Ellis 2001). If the internal colors are not homogeneous, then this may help explain why the LOS velocity dispersion determined from the lens model is higher than expected from fundamental plane considerations, assuming passive evolution since $z \sim 1$.

CFRS 03.1077 demonstrates the potential offered by a detailed study of Einstein ring lenses. With the advent of 8 m class telescopes, especially those equipped with integral field spectrographs, the spectroscopic data reported here can now be very significantly improved. Images at other wavelengths should be obtained to establish or identify additional lensed images in order to more tightly constrain the lens model.

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