

WHEN A STANDARD CANDLE FLICKERS

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ABSTRACT

The Crab Nebula is the only hard X-ray source in the sky that is both bright enough and steady enough to be easily used as a standard candle. As a result, it has been used as a normalization standard by most X-ray/gamma-ray telescopes. Although small-scale variations in the nebula are well known, since the start of science operations of the *Fermi* Gamma-ray Burst Monitor (GBM) in 2008 August, a $\sim 7\%$ (70 mCrab) decline has been observed in the overall Crab Nebula flux in the 15–50 keV band, measured with the Earth occultation technique. This decline is independently confirmed in the ~ 15 –50 keV band with three other instruments: the *Swift* Burst Alert Telescope (*Swift*/BAT), the *Rossi X-ray Timing Explorer* Proportional Counter Array (*RXTE*/PCA), and the Imager on-Board the *INTEGRAL* Satellite (IBIS). A similar decline is also observed in the ~ 3 –15 keV data from the *RXTE*/PCA and in the 50–100 keV band with GBM, *Swift*/BAT, and *INTEGRAL*/IBIS. The pulsed flux measured with *RXTE*/PCA since 1999 is consistent with the pulsar spin-down, indicating that the observed changes are nebular. Correlated variations in the Crab Nebula flux on a ~ 3 year timescale are also seen independently with the PCA, BAT, and IBIS from 2005 to 2008, with a flux minimum in 2007 April. As of 2010 August, the current flux has declined below the 2007 minimum.

Key words: pulsars: individual (Crab Pulsar) – X-rays: individual (Crab Nebula)

1. INTRODUCTION

X-ray and gamma-ray astronomers frequently consider the Crab supernova remnant to be a steady standard candle suitable as a calibration source (e.g., Kirsch et al. 2005; Jourdain & Roques 2009; Weisskopf et al. 2010; Meyer et al. 2010). Jourdain & Roques (2009) presented over five years of the Spectrometer on *INTEGRAL* (SPI, 20 keV–8 MeV) observations, with fitted flux normalizations at 100 keV consistent with being constant to within the $\sim 3\%$ quoted errors. On the basis of data from *XMM-Newton*, *INTEGRAL*, *Swift*, *Chandra*, *RXTE*, and several earlier missions, Kirsch et al. (2005) have concluded that the Crab flux can be described at least up to 30 keV by the same spectrum proposed by Toor & Seward (1974) three decades earlier: $dN/dE = (9.7 \pm 1.0)E^{-(2.1 \pm 0.03)}$ photons $\text{cm}^{-2} \text{s}^{-1} \text{keV}^{-1}$, i.e., they describe the Crab as a standard candle.

Driven by the central pulsar's spin-down luminosity, the surrounding remnant consists of a cloud of expanding thermal

ejecta and a synchrotron nebula (Hester 2008) with an integrated luminosity $\sim 10^{38}$ erg s^{-1} . The pulsar provides a shocked wind that accelerates electrons and positrons to energies $\sim 10^7$ GeV and a source of kinetic energy driving turbulent motion of a ring of wisps nearly surrounding the synchrotron nebula. High-resolution observations reveal wisps and knots moving at velocities up to $0.7c$ from radio to X-ray energies (Hester et al. 1995, 2002; Greiveldinger & Aschenbach 1999; Bietenholz et al. 2001; Mori et al. 2004). A central torus and jet structure extending out from the pulsar were observed in X-rays by *Chandra* (Weisskopf et al. 2000), aligned closely with the pulsar's proper motion (Ng & Romani 2006). The nebular emission is considered to be a combination of synchrotron radiation up to ~ 100 MeV and a harder inverse Compton spectrum extending up to TeV energies (De Jager et al. 1996).

Observations of the 8 GHz nebular flux in 1985 (Aller & Reynolds 1985) showed a decrease of $0.167\% \pm 0.015\% \text{yr}^{-1}$, consistent with predictions by Reynolds & Chevalier (1984) for an expanding synchrotron-emitting cloud. At optical wavelengths, Smith (2003) reported a decrease in the

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nebula-integrated flux of $0.5\% \pm 0.2\% \text{ yr}^{-1}$ from 1987 to 2002. At X-ray energies (2–28 keV), 1996–2002 *BeppoSAX* observations described by Verrecchia et al. (2007) included a 2% systematic error to account for the observed fluctuations in time. In the 35–300 keV energy region, Ling & Wheaton (2003) reported $\sim 10\%$ variations in the flux observed with the Burst and Transient Source Experiment (BATSE) on the *Compton Gamma-Ray Observatory* (CGRO) over periods of days to weeks.

Much et al. (1995) reported a $\sim 40\%$ increase in the unpulsed flux (0.75–30 MeV) measured with the CGRO Compton Telescope (COMPTEL) between 1991 April/May and 1992 August/September. At the same time, De Jager et al. (1996) reported a $\sim 50\%$ decrease in the 75–150 MeV flux and steady emission from 150 MeV to 30 GeV measured with the CGRO Energetic Gamma-Ray Experiment Telescope (EGRET) between 1991 and 1993. They interpret this as steady Compton emission >150 MeV from long-lived ~ 5 –100 GeV electrons and <150 MeV synchrotron emission from shorter-lived 100 TeV to 1 PeV electrons accelerated by time-variable small-scale shock structures. A change in this electron acceleration mechanism would drop a portion of the electrons from the range responsible for the EGRET emission to the COMPTEL range, resulting in the observed fluxes. The *Fermi* Large Area Telescope (LAT) found no variations with time in the 100 MeV–30 GeV band from 2008 August to 2009 April (Abdo et al. 2010a). Recently, *Fermi* LAT (Abdo et al. 2010b) reported flares from the Crab Nebula above 100 MeV in 2009 February and 2010 September. *AGILE* simultaneously detected the 2010 flare (Tavani et al. 2010). Reports by the High Energy Gamma-Ray Astronomy experiment (Aharonian et al. 2004), the Major Atmospheric Gamma-ray Imaging Cherenkov telescope (Albert et al. 2008), the High Energy Spectroscopic System (Aharonian et al. 2006), and the Very Energetic Radiation Imaging Telescope System (Wakely 2010) provide no evidence for time variability, consistent with expectations for higher energies.

Although it is extremely difficult to obtain absolute flux measurements with accuracy $\sim 1\%$ with a single instrument, we have analyzed independent data sets from four separate operating missions. In Section 2, we present the Crab light curves measured independently by the *Fermi* Gamma-ray Burst Monitor (*Fermi*/GBM), the *Swift* Burst Alert Telescope (*Swift*/BAT), the Imager on-Board the *INTEGRAL* Satellite (*INTEGRAL*/IBIS) and the Joint European X-ray Monitor (JEM-X), and the *Rossi X-ray Timing Explorer* Proportional Counter Array (*RXTE*/PCA). In Section 3, we summarize the results and discuss their implications.

2. OBSERVATIONS AND RESULTS

2.1. *Fermi* GBM

The GBM instrument (Meegan et al. 2009), sensitive from 8 keV to 40 MeV, provides nearly continuous full-sky coverage via the Earth occultation technique, successfully demonstrated with BATSE (Harmon et al. 2002; Ling et al. 2000). The Harmon et al. (2002) approach has been adapted for GBM (Wilson-Hodge et al. 2009; Case et al. 2010). To date, six persistent and two transient sources have been detected above 100 keV (Case et al. 2010), including the Crab.

The GBM implementation of the Earth occultation technique uses both CTIME (8 energy channels with 0.256 s resolution) and CSPEC data (128 energy channels with 4.096 s resolution). A detailed detector response model has been developed based

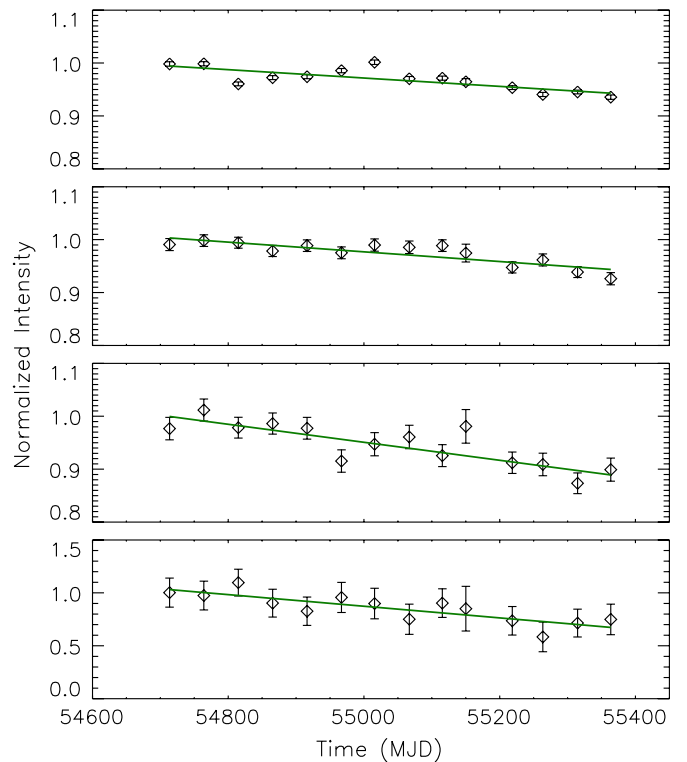


Figure 1. From top to bottom: 50 day average GBM Crab measurements for 12–50, 50–100, 100–300, and 300–500 keV. Solid lines are fits used to compute the percent change in rate.

on Geant4 simulations confirmed by extensive ground testing in order to determine the response as a function of orientation (Hoover et al. 2008; Bissaldi et al. 2009). In flight, fits to background lines (e.g., 511 keV) over time show a stable gain and energy resolution in all the GBM detectors and electronics, with lines typically within 1% of their expected position.

The Crab light curve measured in four energy bands with GBM from 2008 August 12 through 2010 July 13 (MJD 54690–55390) is shown in Figure 1. With respect to the rate on MJD 54690, the Crab rate appears to have decreased steadily by more than 5%: the decrease is $5.4\% \pm 0.4\%$, $6.6\% \pm 1.0\%$, $12\% \pm 2\%$, and $39\% \pm 13\%$ in the 12–50, 50–100, 100–300, 300–500 keV bands, respectively. Inclusion of a linear decline in the 12–50 keV band improves reduced χ^2 to 605.8/130 = 4.66 from 956.3/131 = 7.30 for a constant Crab.

2.2. *RXTE* PCA

Frequent observations with the *RXTE* PCA were made to monitor the radio-X-ray phase of the Crab pulsed emission (Rots et al. 2004) and for calibration purposes (Jahoda et al. 2006; Shaposhnikov 2010). In the PCA, the Crab is bright (~ 2500 counts s^{-1} detector $^{-1}$). Unrejected background from all sources amounts to about 1 mCrab. The PCA is a relatively simple instrument, with commanded changes in operating conditions limited to the high voltage. Data since the last high voltage change in 1999 for PCU 2, 3, and 4 are used in this Letter.

The PCA response has two small time-dependent effects, both accounted for in the response matrices. First, xenon is slowly accumulating in the front veto layer (nominally filled with propane) and reducing the low energy sensitivity with time. Second, there is a small energy drift in the pulse height

channel boundaries, so that a constant channel selection samples a slowly varying energy band. Both effects can influence the rate, though flux determinations (i.e., conversion of count rate to flux) account for this. In particular, the correction for changing opacity of the front veto layer is negligible in the 15–50 keV band. Our observed changes in the Crab rate (see Figure 2) are more than five times larger than these effects combined.

Figure 2 shows total Crab rates for individual *RXTE* pointed observations. From MJD 54690–55435 the Crab rate in PCU 2 declines by $5.1\% \pm 0.2\%$ and $6.8\% \pm 0.3\%$ in the 2–15 and 15–50 keV bands, respectively, relative to MJD 54690. Similar results, variations of 2%–7%, are seen if the bands are further subdivided. In spectral fits to individual PCA observations, the power-law index softens and the normalization and absorption column gradually increase with time, with no clear correlation with flux. These light curves were produced using *RXTE*/PCA standard 2 data (129 energy channel, 16 s) that were extracted, background subtracted, dead-time corrected using standard *RXTE* recipes,¹⁸ and corrected for the known time dependence of the response.

From visual inspection of the *RXTE* light curve, three evident peaks suggest a periodic or quasi-periodic variation with a period of 1000–1500 days. To quantify these impressions we constructed a power spectrum, shown in the third panel of Figure 2, and conducted a search for periodic signals. We averaged the corrected 15–50 keV PCU 2 rates within uniformly spaced bins, using three bins per year, with the yearly interval where Crab cannot be observed because of Sun constraints occurring in the center of every third bin. A linear trend, which passed through the first and last binned rate, was subtracted from the rates, to limit the bleeding of low frequency power into higher frequency bands. The power spectrum was then created from the Fourier transform of the binned rates. The lower five points in the plot are from individual Fourier amplitudes, with the remainder rebinned to reduce errors. A maximum likelihood fit to the unbinned power spectrum was made using a power-law model. The best-fit model is shown, which has a power-law index of 2.1 ± 0.4 .

Standard pulse search methods such as the Lomb test are inappropriate because of the underlying red noise power spectrum. The test statistic we adopted is the improvement in χ^2 between fitting the binned rates to a quadratic and to a quadratic plus a sinusoid. The quadratic accounts for the low frequency trend in the rates. Since the source power dominates the counting statistics, we use uniform errors in the fits, setting $\sigma^2 = P/\Delta t$, where P is the power spectrum model at the middle of the region where a periodicity may be present ($8.5 \times 10^{-4} \text{ day}^{-1}$) and Δt is the bin width. As seen in the bottom panel of Figure 2, a peak in the $\Delta\chi^2$ is seen at $(8.5 \pm 0.7) \times 10^{-4} \text{ day}^{-1}$, corresponding to a period of 1180 ± 100 days. However, its significance is only 2σ . A longer history of the Crab flux will be needed to determine if this feature is a property of the source, or only a statistical fluctuation. Interestingly, this peak value is consistent with twice the period of 568 ± 10 days found in Crab radio timing noise from 1982 to 1989 (Scott et al. 2003).

The Crab pulsed flux measured using PCU 2 event mode data (250 μs , 129 energy channels, top layer) is shown in Figure 3. Although the pulsed flux (upper panel) steadily decreases at $\sim 0.2\% \text{ yr}^{-1}$, consistent with the pulsar spin-down, the larger (several percent per year) variation in the signal is not seen in the pulsed emission and clearly seems to be nebular in origin.

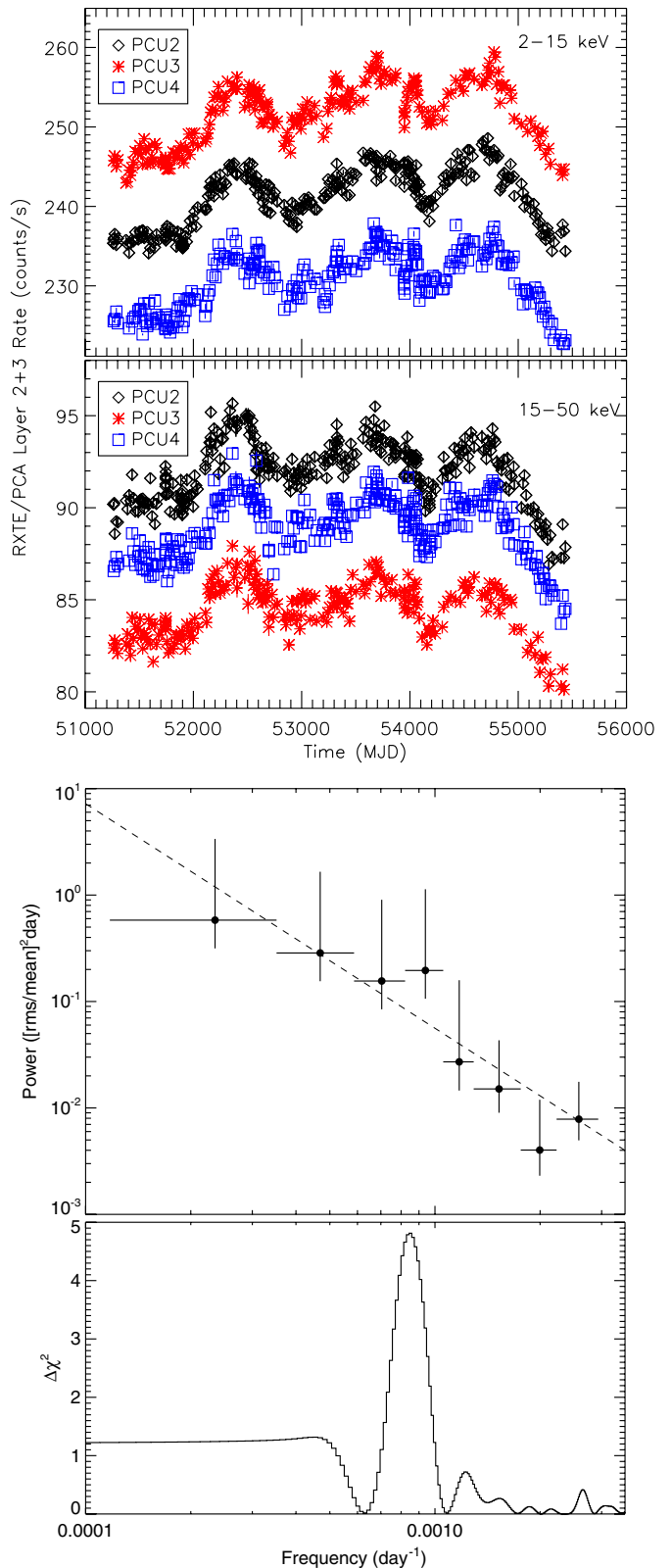


Figure 2. The upper two panels show total Crab rates (nebula + pulsar) for layers 2+3 of PCU 2 (diamonds), PCU 3 (asterisks), and PCU 4 (squares) in the 2–15 and 15–50 keV bands, respectively. The third panel shows the power spectrum of the *RXTE* 15–50 keV rates. The error bars give 68% confidence intervals. The dashed line is the best-fit power law. The bottom panel shows the test statistic for a search for periodic signals.

¹⁸ <http://heasarc.gsfc.nasa.gov/docs/xte/recipes/>

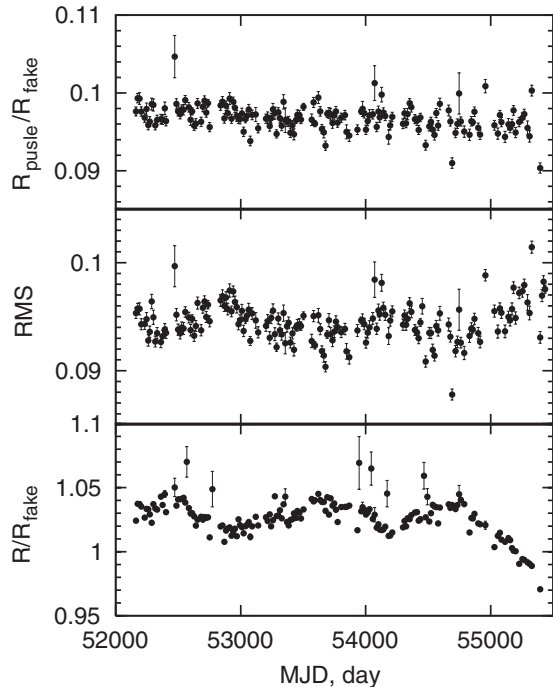


Figure 3. Top: *RXTE*/PCA pulsed flux (3.2–35 keV). Center: fractional rms amplitude for the first two harmonics of the pulse period. Bottom: total Crab rate in PCU 2. Rates in the top and bottom panels are normalized by the response-predicted count rate R_{fake} in the 3.2–35 keV band.

2.3. *INTEGRAL* IBIS and JEM-X

Here, we present results from the JEM-X (3–35 keV; Lund et al. 2003) and the *INTEGRAL* Soft Gamma-Ray Imager layer of IBIS (ISGRI, 15 keV–10 MeV; Ubertini et al. 2003) on board *INTEGRAL* (Winkler et al. 2003). The Crab has been observed every spring and fall with *INTEGRAL* since 2002, mainly for calibration purposes. To reduce systematic effects as much as possible, we have selected on-axis ($<0.25^\circ$) observations for JEM-X2 and observations within 10° of on-axis for IBIS/ISGRI. We include only JEM-X data using the latest on board software (since MJD 53068).

In Figure 4, ISGRI and JEM-X2 count rates from individual pointings averaged over the 3 day *INTEGRAL* orbit with rms errors are shown. The ISGRI data were analyzed with the Off-line Analysis (OSA) package (Courvoisier et al. 2003) version 9 with the settings used for the *INTEGRAL* Galactic Bulge monitoring program¹⁹ light curves (Kuulkers et al. 2007). Using images from individual pointings, the point-spread function of ISGRI is fitted. These images are integrated in a given energy band after gain, offset, and charge loss corrections are performed for each event. A time-dependent effective area correction, usually performed by assuming that the Crab flux is constant, has been excluded from these data, meaning that not all systematic effects are taken into account. Known effects include residuals in gain and charge loss corrections, present with an amplitude of $\sim 1\%$ – 2% , varying on month–years timescales. Similarly, for JEM-X, the ad hoc piecewise linear correction (added to OSA to reduce time trends in the Crab flux) was excluded from the standard OSA analysis. JEM-X consists of two identical units: JEM-X1 and JEM-X2. During the period of interest, JEM-X2 has mostly been in standby mode and JEM-X1 the active unit. A gradual decrease has been observed in the sensitivity of

JEM-X1, so only JEM-X2 is shown. The scatter in the JEM-X data is large compared to the observed Crab variations, especially below 10 keV. From MJD 54690–55390, the ISGRI 15–50, 50–100, and 100–300 keV flux decreases by $8.2\% \pm 1.1\%$, $8.3\% \pm 1.1\%$, and $5.7\% \pm 1.0\%$, respectively, relative to MJD 54690.

2.4. *Swift* BAT

Swift/BAT is a coded aperture telescope operating in the 14–150 keV range (Barthelmy et al. 2005). The *Swift*/BAT 14–50 and 50–100 keV light curves (see Figure 5) are based on publicly available 58 month light curves²⁰ from the *Swift*/BAT all-sky hard X-ray survey (Tueller et al. 2010; Baumgartner et al. 2010) extended to 2010 May 30 by the BAT team. We binned data from individual *Swift* pointings in 50 day intervals, eliminating pointings of less than 200 s duration and those in which less than 15% of the BAT detectors were illuminated by the Crab. The statistical errors on each data point are small (0.1%) and are dominated by systematic errors. We estimate the systematic errors to be $\sim 0.75\%$ by assuming that the long-term variations in the light curve are due to real variations in the Crab, and that the shorter term variations around that trend are representative of the systematic error. The BAT data show variations in the Crab flux at the level of $\sim 3\% \text{ yr}^{-1}$. From MJD 54690–55340, BAT observes a decrease of $6.7\% \pm 0.7\%$ and $10.4\% \pm 0.8\%$ in the 15–50 and 50–100 keV bands, respectively, relative to the rate on MJD 54690, similar to the decrease seen by GBM in the same energy range.

3. DISCUSSION AND SUMMARY

Figure 5 shows composite light curves combining the overlapping results from *RXTE*, *INTEGRAL*, *Swift*, and *Fermi*/GBM. All instruments agree well from 2008 to 2010, with all instruments registering a decline in the Crab 15–50 keV flux of $\sim 7\%$ (70 mCrab) over the 2 years starting at MJD 54690, with a similar decline in the 50–100 keV band. PCA and BAT continue to agree back to the start of the *Swift* mission. For *RXTE*, *Swift*, and *INTEGRAL*/ISGRI the latest measurements shown are significantly below previous minima. *INTEGRAL*/ISGRI shows evidence for the dip near MJD 54100–54200 and the increase before \sim MJD 53700, with similar but less significant variations seen in JEM-X2. Prior to this time, the PCA measurements show continued variations extending back to \sim MJD 52000, which are not seen with ISGRI in the 20–50 keV band. We investigated the effect of a change in the default dithering pattern since 2006 March, but found that this cannot explain the observed difference. Known systematic errors in ISGRI energy reconstruction are expected to account for $\sim 1\%$ – 2% deviations. Beginning at \sim MJD 54000, there is a strong correlation among the results from the four independent instruments with very different signal-to-noise characteristics and observing techniques: Earth occultation, coded-mask imaging, and collimated detectors. The range of techniques strengthens the case that the variation is intrinsic to the Crab. We found no apparent correlations between these variations and variations in the *INTEGRAL*/SPI anti-coincidence detector count rates or GBM count rates, disfavoring local background condition changes as a possible origin, and further supporting a Crab origin. The pulsed flux stability suggests that the observed variations are nebular.

The observed time variability may be explained by models of the Crab pulsar wind flow. In some models (e.g., Camus et al.

¹⁹ <http://integral.esac.esa.int/BULGE/>

²⁰ <http://swift.gsfc.nasa.gov/docs/swift/results/>

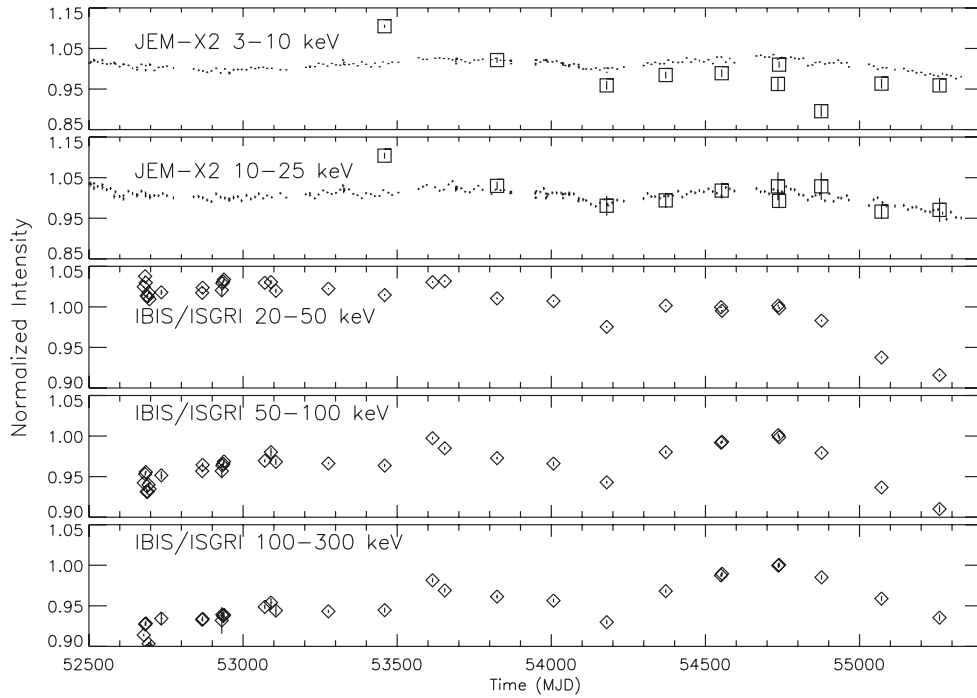


Figure 4. *INTEGRAL* 3 day averaged light curves of the Crab measured in the 3–10 and 10–25 keV bands with JEM-X, and the 20–50, 50–100, and 100–300 keV bands with ISGRI. Normalized *RXTE* PCU2 rates in the 2–15 and 15–50 keV bands are overlotted with the JEM-X data for comparison.

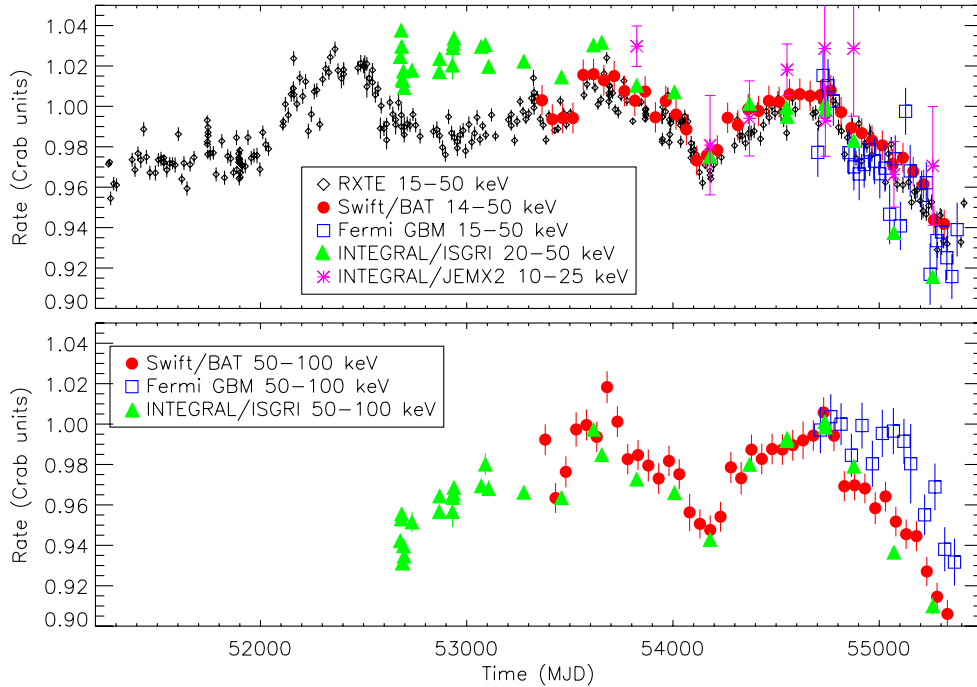


Figure 5. Composite Crab light curves for *RXTE*/PCA (15–50 keV: black diamonds), *Swift*/BAT (top: 14–50 keV, bottom: 50–100 keV: red filled circles), *Fermi*/GBM (top: 15–50 keV, bottom: 50–100 keV: open blue squares), *INTEGRAL*/ISGRI (top: 20–50 keV, Bottom: 50–100 keV: green triangles), and *INTEGRAL*/JEM-X2 (10–25 keV). Each data set has been normalized to its mean rate in the time interval MJD 54690–54790. All error bars include only statistical errors.

2009), a radial plasma flow in the equatorial plane decelerates downstream of a termination shock located at a radius of about 0.5 lt-yr and near the inner ring observed in X-rays (Weisskopf et al. 2000). Due to adiabatic and synchrotron losses in the fluid the flow becomes inhomogeneous with large variations in local magnetic field strength. These magnetosonic waves are relativistic and the variability timescale is roughly the fluid crossing time across the shock diameter or 1–2 years. Alternatively (Spitkovsky & Arons 2004), variability on scales

of the ion Larmor radius may result from cycles of compression of the electron–positron plasma induced by magnetosonic waves caused by the cyclotron instability in the ion orbits.

Chandra (Weisskopf et al. 2000; Mori et al. 2004) and *XMM-Newton* (Kirsch et al. 2006) observations of the Crab suffer from pileup effects, making it difficult to monitor absolute fluxes at the level of a few percent. No *Chandra* ACIS observations of the Crab were performed from MJD 54135–55466. Nevertheless, both instruments have shown that the spectrum of the

synchrotron X-rays grows distinctly softer as distance from the pulsar increases. Since higher energy electrons have shorter synchrotron lifetimes, the spectrum becomes softer as the particles move outward and synchrotron losses grow. Alternatively, the site of the main particle acceleration or the spectral steepening as a function of distance from the pulsar and shock region may vary with time.

The differential photon spectrum dN/dE produced by synchrotron-emitting electrons depends on magnetic field strength B and photon energy E as $dN/dE \sim B^\gamma E^{-\gamma}$, where γ is the power-law photon energy index (Felten & Morrison 1966), suggesting that the observed change in flux could be produced either by a change in the accelerated electron population or a change in the nebular magnetic field of a few percent.

In summary, the widely held assumption that the Crab can be used as a standard candle, suitable for normalizing instrument response functions and for calibrating X-ray instruments, should be treated with caution. Although obtaining absolute calibrations and instrument normalizations at $\sim 1\%$ is difficult, the results presented here from four independent spacecraft demonstrate that in fact the nebular X-ray/gamma-ray emission from the Crab varies at a level of $\sim 3.5\% \text{ yr}^{-1}$. The variation is seen in the nebular emission, and so apparently results from changes in the shock acceleration or the nebular magnetic field. We cannot predict if the present decline will continue or if the ~ 3 year pattern will persist. Longer baselines and multi-wavelength observations are needed to answer these questions.

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