

HYPERVELOCITY A AND B STARS SHOULD BE SLOW ROTATORS

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ABSTRACT

The most commonly accepted explanation for the origin of hypervelocity stars in the halo of the Milky Way is that they are the result of tidal disruption of binaries by the massive black hole at the center of the Galaxy. We show that if this scenario is correct and if the original binary properties are similar to those in the local stellar neighborhood, then the hypervelocity stars should rotate with velocities measurably lower than those for field stars of similar spectral type. This may prove to be a more direct test of the model than trying to predict the position and velocity distributions.

Subject headings: Galaxy: center — Galaxy: kinematics and dynamics — scattering — stars: rotation

1. INTRODUCTION

The discovery of a population of young, high-velocity stars in the halo of the Galaxy (Brown et al. 2005), termed hypervelocity stars, is quite naturally explained in terms of ejection from the region around the massive black hole at the center (e.g., Hills 1988). The original proposal by Hills was that stellar binaries on grazing orbits are tidally disrupted by the black hole, leaving one star bound and one unbound with a high ejection velocity. The connection is further enhanced by the fact that the hypervelocity stars are late B-type stars and thus of similar spectral type to the population of peculiarly young stars present in the central parsec (Ghez et al. 2003; Eisenhauer et al. 2005), although this is partly the result of sample selection (Brown et al. 2006).

Examining the dynamical link between the hypervelocity stars and the so-called Sgr A* cluster may help to explain the puzzling origins of this latter population. The modern realization of Hills’s scenario invokes the scattering (Perets, Hopman & Alexander 2007) of young stellar binaries onto nearly radial orbits by massive perturbers (usually giant molecular clouds) from Galactocentric radii of roughly several parsecs (where such stars can form more naturally than in the strong tidal field of the central black hole). The dynamical disruption of these binaries then produces the observed hypervelocity stars and a bound population that results in the Sgr A* cluster (Gould & Quillen 2003; Ginsburg & Loeb 2006). However, some issues remain. The apastra of the cluster stars do not extend as far out as the Hills scenario would suggest, and the scenario does not naturally explain the second population of young stars in the Galactic center, the so-called He emission-line stars (e.g., Paumard et al. 2006), as these stars are too far from the black hole to be the result of tidal disruption.

At least two alternative scenarios exist for the origin of the hypervelocity stars. They may also be ejected through two-body encounters with massive, dark objects orbiting the central black hole (Yu & Tremaine 2003). Indeed, the existence of a second, intermediate-mass, black hole ($\sim 10^3$ – $10^4 M_\odot$) in the Galactic center has been suggested (Hansen & Milosavljevic 2003) as a mechanism for explaining the presence of the young stars as the tidally stripped remnants of an inspiralling stellar cluster (Gerhard 2001). In this scenario there is no direct one-to-one link between hypervelocity stars and Sgr A* cluster

stars, although both are assumed to be drawn from the same parent population.

In a related fashion, stars may be ejected by two-body encounters with stellar-mass black holes $\sim 10 M_\odot$ (O’Leary & Loeb 2006), a population of which is expected to form a cusp around the central black hole as a result of mass segregation in the Galactic potential (Morris 1993; Miralda-Escudé & Gould 2003). The efficiency of this process is lower than scattering by an intermediate-mass black hole, and more subject to stellar destruction by tidal disruption or physical collisions (Ginsburg & Loeb 2007). However, it may still be significant and may be applicable if the observed young stars had their origins in a gravitationally unstable accretion disk (Levin & Beloborodov 2003) rather than a tidally disrupted cluster.

The original scenario due to Hill requires a binary, whereas the scenarios in which stars scatter off black holes do not. Furthermore, the orbital separations of the binaries that give rise to hypervelocity stars are small enough (Hills 1988; Gualandris, Portegies Zwart, & Sipior 2005; Bromley et al. 2006) that the stellar rotation of the star is affected by the rotation of hypervelocity stars to test whether they did indeed spend some part of their lifetime in a binary. If the rotation velocities of hypervelocity stars are similar to those of binary stars of the same spectral type and measurably different to equivalent single stars, then we may interpret this as evidence in favor of the binary disruption origin of hypervelocity stars.

In § 2 we examine the empirical trends of stellar rotation with binary separation, for early-type stars, and how this translates into a relation between stellar rotation and ejection velocity, via the auspices of Hills’s scenario. Observational confirmation would provide a strong argument in favor of this model, as the alternative models do not require binary stars and thus do not imply the same consequence.

2. A RELATION BETWEEN ROTATION AND EJECTION VELOCITY

The rotational properties of early-type stars in binaries have been compiled by Abt & Boonyarak (2004). Their most striking finding is that stars in binaries seem to show a measurable decrease in rotation velocity (relative to single stars of similar type) at separations well beyond the limit at which tides are expected to operate. For B-type stars, the most relevant to the observations of Brown et al., the trend of $V \sin i$ with P_{orb} is flat out to $P_{\text{orb}} \sim 10^4$ days, with a mean velocity of 83 ± 16 km s⁻¹, measurably lower than the value of 134 ± 7 km s⁻¹ for single stars (Abt, Levato, & Grosso 2002).

More powerful discriminants are possible if the observa-

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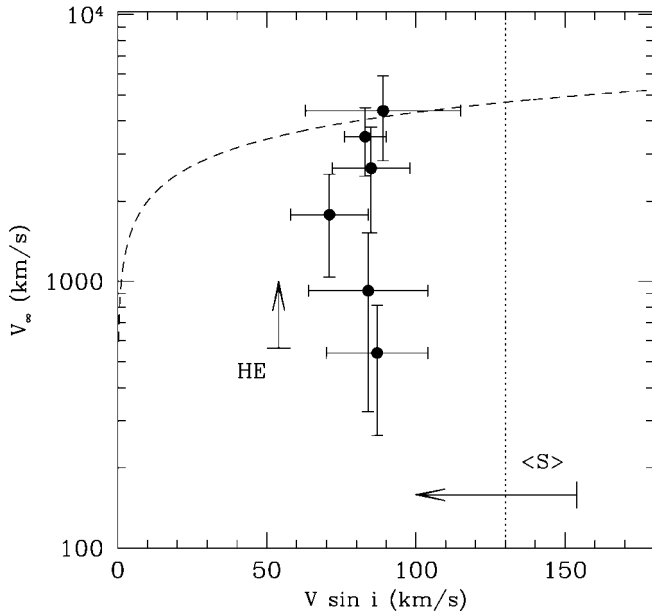


FIG. 1.—Points show the relationship between rotation velocity and ejection velocity expected, on the basis of Hills’s scenario, for B-type stars. The vertical dotted line indicates the average rotation velocity of single stars. The dashed line assumes strict synchronization of the star with the original binary rotation, with an ejection velocity given by equation (1) of Bromley et al. (2006) and randomly oriented inclinations. The arrow at the bottom right is the upper limit on the average rotation velocity of the Sgr A* cluster B-type stars from Eisenhauer et al. (2005). The arrow at left shows the parameters of the star HE 0437–5439 (Edelmann et al. 2005), where we have used the measured radial velocity as a lower limit, since it has presumably slowed as it climbed out of the Galactic potential well.

tional selection is extended to encompass stars of early A type. For stars of spectral type A0–A5, Abt & Boonyarak show that the rotation of stars in the shortest-period binaries show the expected synchronization trend, but that significantly reduced rotation extends out as far as orbital periods ~ 100 days, before returning to rates consistent with single stars at longer periods. An orbital period of 100 days corresponds to a separation of 0.77 AU (for a binary consisting of two $3 M_{\odot}$ main-sequence A stars), which yields an ejection velocity of 935 km s^{-1} using equation (1) from Bromley et al. (2006). Thus, the orbital separations of interest for hypervelocity stars are precisely those for which rotation velocities are affected by the presence of a companion.

2.1. Velocity-Velocity Relation

In order to investigate how the rotational properties of the original binary members are related to their final fates as hypervelocity stars, we have performed our own version of the scattering calculations of Hills, Gualandris et al., and Bromley et al., using a Bulirsch-Stoer integrator with energy precision $\epsilon = 10^{-9}$ for each step size. We allow for random binary orientations and assume a logarithmic distribution of initial separations, for a range from 0.01 to 10 AU. We start the binaries at a distance of 5000 AU from the center, with velocities of 250 km s^{-1} and a range of impact parameters, weighted by equal probability per unit area, out to impact parameters of 3300 AU (which corresponds to closest approach distances of ~ 100 AU). As noted by previous authors, there is a distinct trend indicating that higher ejection velocities result from the disruption of closer binaries. However, there is significant scatter in the relation, as a result of the other variables in the

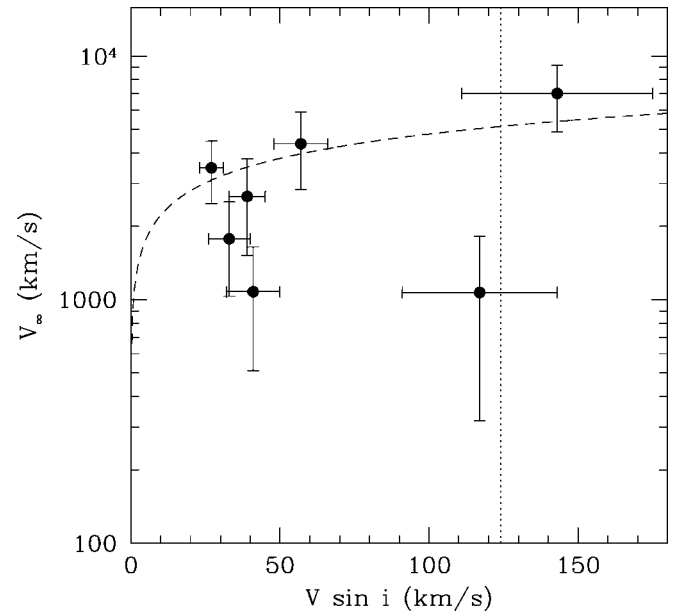


FIG. 2.—Points show the expected relationship between rotation velocity and ejection velocity, on the basis of Hills’s scenario, for early A stars. The vertical dotted line indicates the average rotation velocity of single stars and the dashed line is as described in Fig. 1.

problem, mostly related to the orientation of the binary relative to its trajectory. For each of the initial period bins used by Abt & Bonnyarek, we average over the distribution of final V_{∞} , to derive the relationship between average $V \sin i$ and average V_{∞} (emerging from the Galactic center). The resulting relationship is shown in Figure 1 for B stars and in Figure 2 for A0–A5 stars. We see that the hypervelocity B stars are expected to be systematically slower rotators than the $\sim 130 \text{ km s}^{-1}$ measured for isolated B stars (Abt et al. 2002). For the A stars, there is more structure in the relation. At the highest ejection velocities (resulting from the closest binaries), the spins follow the synchronization trend, but there is a low rotation tail that extends down to velocities $\sim 1000 \text{ km s}^{-1}$, before the rotation returns to values more characteristic of single stars. In both figures, the error bars on V_{∞} indicate the dispersion about the mean, rather than the error in the mean, as in the case of $V \sin i$.

One might wonder whether the passage of the star close to the central black hole can have an effect on the spin. However, the distances of closest approach are rarely close enough (< 1 AU) for there to be a significant tidal spin-up effect, especially since there is only one passage past the black hole. For the member of the original binary that remains bound, repeated encounters with other stars may have an effect over time (Alexander & Kumar 2001), but the ejected star should remain unaffected.

What of those stars that remain bound to the black hole? Ghez et al. (2003) report a rotation of $220 \pm 40 \text{ km s}^{-1}$ for the star SO-2, which suggests that this star is not a natural fit to the favored scenario. This is corroborated by the fact that the measured eccentricity ($e = 0.87$), while high, is not high enough ($e > 0.97$) to match the expectations of the simple binary disruption model (as noted in Ginsburg & Loeb 2006, and which our simulations confirm). For the other S stars, Eisenhauer et al. (2005) find an upper limit to the average rotation velocity of the Sgr A* cluster stars of 154 km s^{-1} , by co-adding the spectra. This shows that they are, at least, not

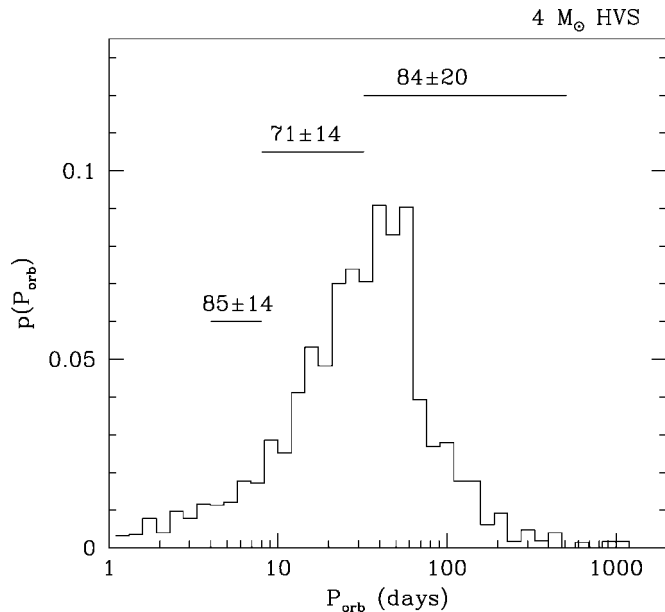


FIG. 3.—Histogram of the original binary orbital periods for the observed hypervelocity B stars. Also indicated are the mean rotation velocities measured by Abt & Boonyarek for field binaries in each of the designated period ranges. We see that the bulk of the hypervelocity stars should be slow rotators.

significantly spun-up by tidal interactions, either in very close binaries or by the central black hole.

2.2. Passage into the Halo

We must also account for the deceleration of the ejected stars as they climb out of the potential well to their present observed positions. To answer this, we generate specific realizations, following the approach of Bromley et al. We sample the distribution of ejected velocities shown above directly, integrating the orbits outward in the Galactic potential model of Bromley et al., with values $\alpha = 2$, $a_c = 8$ pc, and $\rho_0 = 1.27 \times 10^4 M_\odot \text{pc}^{-3}$. We integrate for a random value of the age between 0 and the turnoff age. For $3 M_\odot$, we take this to be 350 Myr, and for $4 M_\odot$ we adopt 160 Myr. We select stars that lie between 10 and 100 kpc (the approximate observational volume).

Figure 3 shows the distribution of original orbital periods of the binaries that give rise to the observable hypervelocity B stars. Interestingly, we see a strong peak at ~ 40 – 60 days and that the vast majority of the observed stars had original orbital periods < 100 days. Thus, we can make the strong prediction that the observed rotation velocities of the stars observed by Brown et al. should be ~ 70 – 90 km s^{-1} .

Figure 4 shows the same distribution in the case of $3 M_\odot$, which we take to represent the early A stars. We see that we can make an even stronger prediction regarding the spins of hypervelocity early A stars, namely, that the great majority should be very slowly rotating, with $V \sin i \sim 30$ – 40 km s^{-1} . Stars with rotation velocities $> 100 \text{ km s}^{-1}$ (common among single A stars) should be quite rare amongst the hypervelocity stars.

3. DISCUSSION

The currently most widely accepted model for the origin of hypervelocity stars is that they correspond to the break-up of binaries that interact with the massive black hole at the Galactic center. We have shown that, based on the empirically deter-

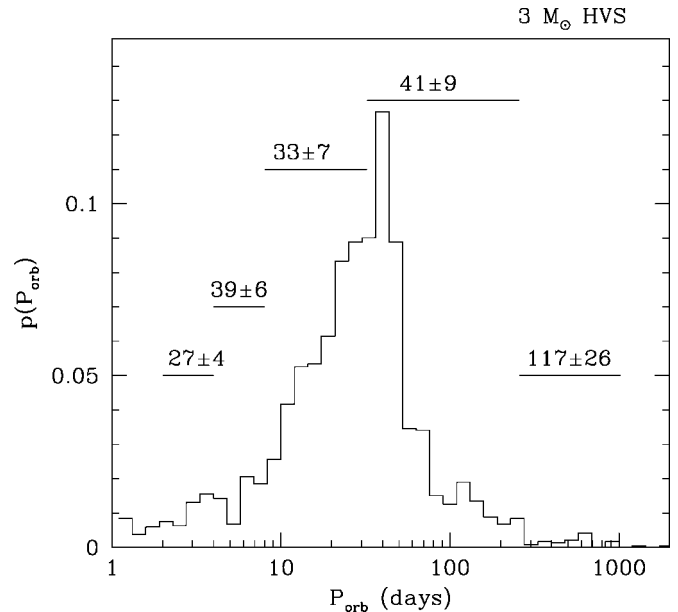


FIG. 4.—Histogram of the probability distribution of the original binary orbital periods for hypervelocity A stars in the observational volume. Also indicated are the mean rotation velocities measured by Abt & Boonyarek for field binaries in each of the designated period ranges. We see that the bulk of the hypervelocity stars should be slow rotators.

mined properties of early-type stars in binaries and the kinematic requirements of the ejection scenario, one should expect the hypervelocity stars to rotate slower than the average for single stars of their type. This is the result of the fact that, in order for a star to attain the velocities observed, it must come from a system whose initial orbital separation is small enough that the rotation is affected by the presence of a binary companion.

In quantitative terms, we have shown that the rotation velocities for the observed late-type B stars are expected to be $\sim 80 \text{ km s}^{-1}$, compared with $\sim 130 \text{ km s}^{-1}$ one finds for most single B-type stars. If the observational selection can be pushed into the regime of the early-type A stars, the prediction gets even stronger, with expected rotation velocities $\sim 40 \text{ km s}^{-1}$. Furthermore, it is worth noting that the discovery of hypervelocity Am stars would be a direct signature of initial binarity, since the peculiar spectral signatures of Am stars are the result of radiative levitation in atmospheres that are very quiescent due to their slow rotation. This is why Am stars are almost always found in binaries. (Abt 1965, 2000)

The preceding assertions assume that the properties of the initial binaries are accurately reflected by the studies of those in the solar neighborhood. However, this is implicit in Hills's scenario, in that it invokes stars that formed far enough from the black hole that there is no reason to expect the star formation properties to be any different (unlike many scenarios for the origins of the young stars in the central parsec, where the physics is likely quite different). Another underlying assumption that one might question is whether the binaries have had enough time to circularize before being disrupted (i.e., are the solar-neighborhood binaries, in some sense, older)? To answer this, we note Abt & Boonyarek's empirical result that the regime of quasi-synchronization extends to orbital periods well beyond that expected by simple tidal theory. Thus, it is likely the result of processes at work during the formation of the binary and which then imprint the rotational properties of the stars well before the binary is disrupted. Even in the case in

which the synchronization is the result of classical tidal interaction, for early-type stars the tidal forces are quite sensitive to the stellar structure, and most of the tidal synchronization takes place in the early stages of the star's evolution (Zahn 1977).

Thus, we consider it a robust prediction that hypervelocity stars resulting from Hills's scenario should be slowly rotating and measurably different from the rotational properties of field stars of similar spectral type. If hypervelocity stars are instead found to be rotating like normal field B or A stars, then this would favor one of the other ejection scenarios, such as scattering by an intermediate-mass black hole (Levin 2006; Baumgardt, Gualandris, & Portegies Zwart 2006; Sesana, Haardt, & Madau 2006). Such a measurement should be within reach of current generations of telescopes. Indeed, the spectra of Brown et al. have a velocity resolution of $\sim 100 \text{ km s}^{-1}$, so only a modest increase of a factor of 2 in spectral resolution is required to start putting some real constraints on the ejection scenario. In fact, the hypervelocity star HE 0437–5439 (Edelmann et al. 2005) already has a measured rotation velocity of

$V \sin i = 54 \pm 4 \text{ km s}^{-1}$. This low value is nicely consistent with the expectations of the binary disruption scenario, although its possible origin in the LMC is a complicating factor.

In almost all cases, if one binary member is ejected, the other remains bound to the black hole in a highly eccentric orbit. Thus, the above considerations apply to these stars as well. Indeed, Hills's scenario has been invoked to explain the origin of the so-called Sgr A* cluster. However, the best-studied member of this group, the star SO-2, does not agree well with the predicted properties from Hills's scenario. The eccentricity is not large enough to be the result of a binary disruption, and it is rotating normally, in contrast to the above prediction. In principle, the presence of other (possibly dark) bodies in the same vicinity may lead to additional perturbations that might explain the present observed parameters. However, the short age of the star renders this explanation problematical. As always, we await further observation.

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